

# A much lower density for the transiting extrasolar planet WASP-7★ (Research Note)

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## ABSTRACT

We present the first high-precision photometry of the transiting extrasolar planetary system WASP-7, obtained using telescope defocussing techniques and reaching a scatter of 0.68 mmag per point. We find that the transit depth is greater and that the host star is more evolved than previously thought. The planet has a significantly larger radius ( $1.330 \pm 0.093 R_{\text{Jup}}$  versus  $0.915^{+0.046}_{-0.040} R_{\text{Jup}}$ ) and much lower density ( $0.41 \pm 0.10 \rho_{\text{Jup}}$  versus  $1.26^{+0.25}_{-0.21} \rho_{\text{Jup}}$ ) and surface gravity ( $13.4 \pm 2.6 \text{ m s}^{-2}$  versus  $26.4^{+4.4}_{-4.0} \text{ m s}^{-2}$ ) than previous measurements showed. Based on the revised properties it is no longer an outlier in planetary mass–radius and period–gravity diagrams. We also obtain a more precise transit ephemeris for the WASP-7 system.

**Key words.** stars: planetary systems — stars: individual: WASP-7

## 1. Introduction

The transiting extrasolar planet (TEP) WASP-7 b was discovered by the WASP consortium (Hellier et al. 2009, hereafter H09) through the detection of transits in front of its F5 V parent star. It is a challenging target for acquiring high-precision transit photometry, due to the brightness of the parent star ( $V = 9.5$ ), the paucity of good nearby comparison stars, the transit duration (3.8 hr), and the orbital period (4.95 d) which is both comparatively long and close to an integer number of days. The charac-

terisation of WASP-7 therefore relied upon the photometry obtained by the WASP-South telescope (Pollacco et al. 2006).

The relatively large scatter in the discovery data meant that the transit shape was poorly delineated. Because of this, the analysis by H09 included an additional constraint in the form of a main sequence mass–radius relation for the host star (e.g. Anderson et al. 2010). The radius, surface gravity and density of the planet resulting from their analysis are  $R_b = 0.91^{+0.046}_{-0.040} R_{\text{Jup}}$ ,  $g_b = 26.4^{+4.4}_{-4.0} \text{ m s}^{-2}$  and  $\rho_b = 1.26^{+0.25}_{-0.21} \rho_{\text{Jup}}$ , respectively. These values placed WASP-7 b in an outlier position in the mass–radius diagram of TEPs, having one of the largest densities within the main planet population (masses  $\lesssim 2 M_{\text{Jup}}$ ). This was interpreted

\* Based on data collected by MiNDSTeP with the Danish 1.54m telescope at the ESO La Silla Observatory.

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by H09 as evidence that WASP-7 b has a massive heavy-element core.

In this work we present the first follow-up photometric observations obtained for WASP-7. The high precision of our observations (0.68 mmag scatter) allows us to obtain the physical properties of the transiting system without needing to impose any constraints on the parameters of the parent star. We find a substantially larger radius, and therefore a lower density and surface gravity. We also greatly improve the orbital ephemeris for the system, so transit midpoints in the 2011 observing season can be predicted to within 45 s instead of 27 min.

## 2. Observations and data reduction

WASP-7 was observed from 00:14 to 07:18 UT on the date 2010/09/07. We used the 1.54 m Danish telescope at ESO La Silla, equipped with the DFOSC focal-reducing CCD imager. A total of 217 integrations were obtained through a Gunn *I* filter (ESO filter #425) and with an exposure time of 60 s. The CCD was windowed down in order to reduce readout time, resulting in a dead time of 50 s between individual observations. The telescope was defocussed to a point spread function diameter of 38 pixels (15'') in order to average out flat-fielding noise and avoid saturating the CCD.

The data were reduced using a pipeline which implements the DAOPHOT aperture photometry routine (Stetson 1987). The images were debiased and flat-fielded, and slight pointing variations were detected and accounted for by cross-correlating against a reference image. A differential-magnitude light curve was constructed by simultaneously fitting a straight line to the out-of-transit observations plus weights to the reference stars used to create the ensemble comparison star. We find that the sizes of the software apertures and the choice of comparison stars has very little effect on the shape of the observed transit. For our final photometry we adopt the light curve which displays the lowest scatter with respect to a fitted transit model (see below). The root-mean-square (rms) of the scatter in this light curve is 0.68 mmag per point. For further details on the approach used to obtain and reduce the data see Southworth et al. (2009, 2010).

## 3. Light curve analysis

The light curve of WASP-7 was modelled using the JKTEBOP code<sup>1</sup>, which in turn is based on the EBOP program (Popper & Etzel 1981; Nelson & Davis 1972). The main parameters of the fit were the sum of the fractional radii of the star and planet,  $r_A + r_b = \frac{R_A + R_b}{a}$  where  $a$  is the semimajor axis, their ratio,  $k = \frac{r_b}{r_A}$  and the orbital inclination,  $i$ . After a preliminary fit was obtained, the measurement errors were rescaled to give a reduced  $\chi^2$  of 1.0. In order to provide an accurate measure of the orbital period we included the time of mid-transit quoted by H09, using the approach given by Southworth et al. (2007a). This results in an orbital ephemeris of:

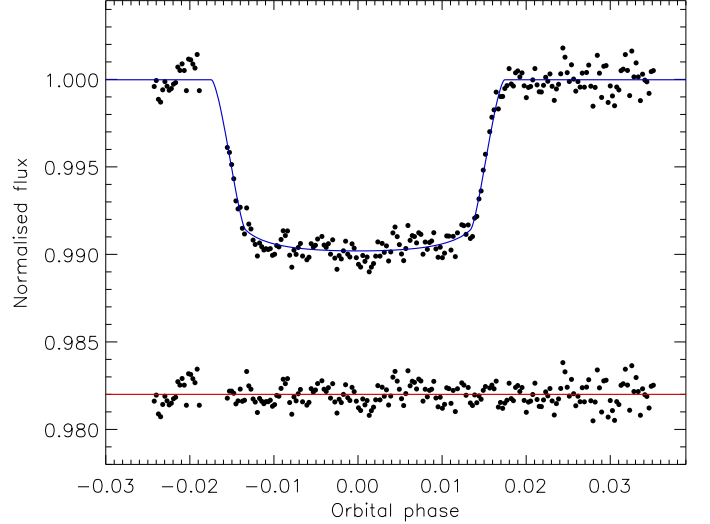
$$T_0 = \text{BJD(TDB)} 2\,455\,446.63493(30) + 4.9546416(35) \times E$$

where  $E$  is the number of orbital cycles after the reference epoch and the bracketed quantities denote the uncertainty in the final digit of the preceding number. This ephemeris is on the BJD(TDB) timescale (Eastman et al. 2010). The timestamps in our data were manually verified to be correct to within  $\pm 2$  s.

<sup>1</sup> JKTEBOP is written in FORTRAN77 and the source code is available at <http://www.astro.keele.ac.uk/~jkt/codes/jktebop.html>

**Table 2.** Final parameters of the fits to the light curve of WASP-7, compared to the results from H09.

	This work	H09
$r_A + r_b$	$0.1207 \pm 0.0068$	0.1001
$k$	$0.0956 \pm 0.0016$	$0.0761 \pm 0.0008$
$i$ (°)	$87.03 \pm 0.93$	$89.6^{+0.4}_{-0.9}$
$r_A$	$0.1102 \pm 0.0061$	0.0930
$r_b$	$0.01053 \pm 0.00070$	0.00708



**Fig. 1.** Phased light curve of WASP-7 compared to the best fit found using JKTEBOP. The residuals of the fits are plotted at the base of the figure, offset from zero.

Limb darkening (LD) was included and its uncertainty accounted for using five different functional laws, and solutions were obtained for the cases when both coefficients were fixed to theoretical values, the linear coefficient was fitted for, and both coefficients were fitted for. Theoretical values were calculated for the measured effective temperature and surface gravity of the star by linearly interpolating within the tables of Van Hamme (1993) and Claret (2000). The results for each of these modelling runs are given in Table 1. We find that the results for different LD laws are generally in good agreement, although the amount of LD is lower than theoretically predicted.

Uncertainties in the fitted parameters were obtained using Monte Carlo and residual-permutation simulations (Southworth 2008; Jenkins et al. 2002). We are able to reject the solutions with all LD coefficients fixed to theoretical values, as they provide a clearly worse fit to the data. The solutions where two LD coefficients are fitted are marginally the best ones, so for the final parameters we adopt their mean values from these four solutions. We adopt the parameter uncertainties from the residual permutation algorithm, which are roughly 1.5 times larger than the Monte Carlo alternatives. This implies that correlated noise is significant in our data, as expected from visual inspection of the plot of the best fit (Fig. 1). The final parameter values and uncertainties are given in Table 2. Compared to those from H09, we find a solution with a lower orbital inclination and larger fractional radius for both the star and the planet. Our results were obtained using the methods of Southworth (2008) and are homogeneous with the results in that work.

**Table 1.** Parameters of the JKTEBOP best fits of the light curve of WASP-7, using different approaches to limb darkening (LD).

	Linear LD law	Quadratic LD law	Square-root LD law	Logarithmic LD law	Cubic LD law
All LD coefficients fixed:					
$r_A + r_b$	$0.1255 \pm 0.0052$	$0.1180 \pm 0.0049$	$0.1194 \pm 0.0048$	$0.1184 \pm 0.0051$	$0.1223 \pm 0.0046$
$k$	$0.09598 \pm 0.00089$	$0.09480 \pm 0.00075$	$0.09524 \pm 0.00075$	$0.09475 \pm 0.00082$	$0.09633 \pm 0.00060$
$i$ (°)	$86.61 \pm 0.52$	$87.43 \pm 0.59$	$87.23 \pm 0.56$	$87.39 \pm 0.63$	$86.80 \pm 0.47$
$u_A$	0.43 fixed	0.20 fixed	0.00 fixed	0.55 fixed	0.20 fixed
$v_A$		0.30 fixed	0.60 fixed	0.25 fixed	0.15 fixed
$r_A$	$0.1145 \pm 0.0047$	$0.1078 \pm 0.0044$	$0.1090 \pm 0.0044$	$0.1081 \pm 0.0046$	$0.1116 \pm 0.0041$
$r_b$	$0.01099 \pm 0.00053$	$0.01022 \pm 0.00048$	$0.01038 \pm 0.00048$	$0.01025 \pm 0.00051$	$0.01075 \pm 0.00045$
$\sigma$ (mmag)	0.7136	0.6977	0.6946	0.7015	0.6828
Fitting for the linear LD coefficient and perturbing the nonlinear LD coefficient:					
$r_A + r_b$	$0.1224 \pm 0.0043$	$0.1192 \pm 0.0047$	$0.1227 \pm 0.0046$	$0.1221 \pm 0.0047$	$0.1200 \pm 0.0044$
$k$	$0.09675 \pm 0.00069$	$0.09587 \pm 0.00072$	$0.09646 \pm 0.00065$	$0.09628 \pm 0.00069$	$0.09623 \pm 0.00071$
$i$ (°)	$86.77 \pm 0.44$	$87.15 \pm 0.55$	$86.77 \pm 0.48$	$86.83 \pm 0.49$	$87.05 \pm 0.49$
$u_A$	$0.245 \pm 0.047$	$0.059 \pm 0.064$	$-0.134 \pm 0.059$	$0.387 \pm 0.064$	$0.190 \pm 0.051$
$v_A$		0.30 perturbed	0.60 perturbed	0.25 perturbed	0.15 perturbed
$r_A$	$0.1116 \pm 0.0039$	$0.1088 \pm 0.0043$	$0.1119 \pm 0.0042$	$0.1113 \pm 0.0043$	$0.1095 \pm 0.0039$
$r_b$	$0.01080 \pm 0.00043$	$0.01043 \pm 0.00047$	$0.01079 \pm 0.00045$	$0.01072 \pm 0.00046$	$0.01054 \pm 0.00042$
$\sigma$ (mmag)	0.6841	0.6829	0.6826	0.6825	0.6827
Fitting for both LD coefficients:					
$r_A + r_b$		$0.1216 \pm 0.0050$	$0.1199 \pm 0.0049$	$0.1221 \pm 0.0054$	$0.1196 \pm 0.0050$
$k$		$0.09583 \pm 0.00115$	$0.09537 \pm 0.00145$	$0.09575 \pm 0.00148$	$0.09527 \pm 0.00151$
$i$ (°)		$86.91 \pm 0.56$	$87.15 \pm 0.60$	$86.89 \pm 0.63$	$87.18 \pm 0.63$
$u_A$		$-0.021 \pm 0.256$	$-0.866 \pm 0.979$	$0.578 \pm 0.391$	$0.066 \pm 0.131$
$v_A$		$0.42 \pm 0.41$	$1.83 \pm 1.60$	$0.56 \pm 0.60$	$0.54 \pm 0.45$
$r_A$		$0.1110 \pm 0.0045$	$0.1095 \pm 0.0045$	$0.1114 \pm 0.0049$	$0.1092 \pm 0.0045$
$r_b$		$0.01064 \pm 0.00051$	$0.01044 \pm 0.00051$	$0.01067 \pm 0.00055$	$0.01040 \pm 0.00052$
$\sigma$ (mmag)		0.6826	0.6818	0.6820	0.6817

**Table 3.** Derived physical properties of WASP-7 using five different sets of theoretical stellar models.

	(Claret models)	(Y <sup>2</sup> models)	(Teramo models)	(VRSS models)	(DSEP models)
$K_b$ (km s <sup>-1</sup> )	$135.6 \pm 2.1$	$135.6 \pm 1.9$	$135.1 \pm 2.2$	$135.4 \pm 1.9$	$134.4 \pm 2.2$
$M_A$ (M <sub>⊙</sub> )	$1.285 \pm 0.059$	$1.287 \pm 0.054$	$1.273 \pm 0.061$	$1.280 \pm 0.052$	$1.254 \pm 0.061$
$R_A$ (R <sub>⊙</sub> )	$1.436 \pm 0.092$	$1.436 \pm 0.092$	$1.431 \pm 0.092$	$1.434 \pm 0.090$	$1.424 \pm 0.092$
$\log g_A$ (cgs)	$4.233 \pm 0.047$	$4.233 \pm 0.047$	$4.231 \pm 0.047$	$4.232 \pm 0.047$	$4.229 \pm 0.047$
$M_b$ (M <sub>Jup</sub> )	$0.96 \pm 0.13$	$0.96 \pm 0.13$	$0.96 \pm 0.13$	$0.96 \pm 0.13$	$0.95 \pm 0.13$
$R_b$ (R <sub>Jup</sub> )	$1.333 \pm 0.093$	$1.334 \pm 0.093$	$1.329 \pm 0.093$	$1.331 \pm 0.092$	$1.322 \pm 0.092$
$\rho_b$ (ρ <sub>Jup</sub> )	$0.407 \pm 0.101$	$0.406 \pm 0.101$	$0.408 \pm 0.101$	$0.407 \pm 0.101$	$0.410 \pm 0.102$
$\Theta$	$0.069 \pm 0.011$	$0.069 \pm 0.010$	$0.070 \pm 0.011$	$0.070 \pm 0.011$	$0.070 \pm 0.011$
$a$ (AU)	$0.06185 \pm 0.00095$	$0.06188 \pm 0.00086$	$0.06166 \pm 0.00098$	$0.06177 \pm 0.00084$	$0.06134 \pm 0.00100$
Age (Gyr)	$2.5^{+0.5}_{-1.0}$	$2.5^{+0.2}_{-1.0}$	$2.3^{+0.6}_{-0.8}$	$2.0^{+0.7}_{-0.6}$	$2.7^{+0.8}_{-0.7}$

#### 4. The physical properties of WASP-7

The light curve parameters alone do not allow the physical properties of the WASP-7 system to be calculated. Additional observed properties are available from H09: the star’s orbital velocity amplitude ( $K_A = 97 \pm 13$  m s<sup>-1</sup>), effective temperature ( $T_{\text{eff}} = 6400 \pm 100$  K), surface gravity ( $\log g = 4.3 \pm 0.2$ ) and metallicity ( $[\frac{\text{Fe}}{\text{H}}] = 0.0 \pm 0.1$ ). Of these, the surface gravity is quite imprecise so is only useful as a consistency check.

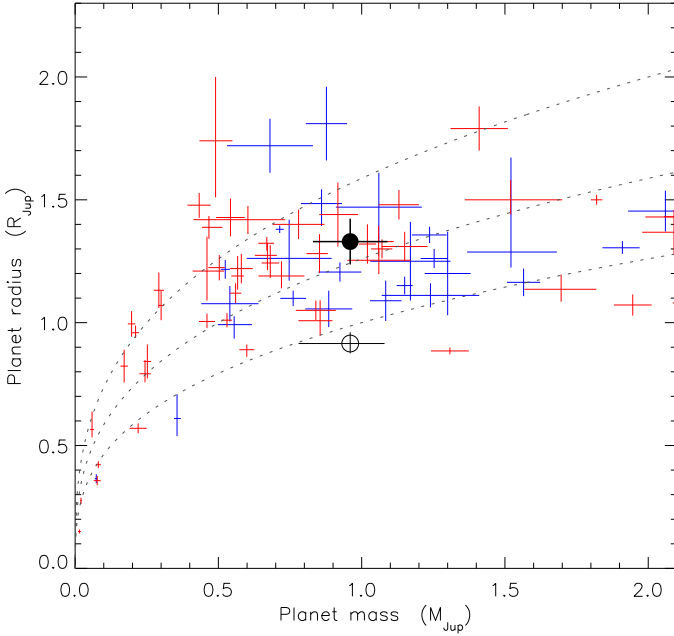
In order to calculate the physical properties of WASP-7 we use the same approach as in Southworth (2009). Firstly, we adopt tabulated predictions from theoretical models of low-mass stars. Secondly, we guess an initial value of the velocity amplitude of the planet,  $K_b$ . We then use the known  $r_A$ ,  $r_b$ ,  $i$ ,  $K_A$  and orbital period to calculate the physical properties of the system. The observed  $T_{\text{eff}}$  and calculated  $R_A$  are then compared to the model-predicted values for a star of the calculated mass ( $M_A$ ), and  $K_b$  is adjusted until the best match is found. This is performed for a grid of ages covering 0.1 to 20 Gyr in 0.1 Gyr steps, by which

the overall best physical properties and age are found. The uncertainty in each of the input parameters is propagated by repeating this procedure with the parameter perturbed, resulting in a detailed error budget for every output parameter.

The above procedure is performed five times, using independent sets of theoretical model predictions. These are: *Claret* (Claret 2004), *Y<sup>2</sup>* (Demarque et al. 2004), *Teramo* (Pietrinferni et al. 2004), *VRSS* (VandenBerg et al. 2006) and *DSEP* (Dotter et al. 2008). The final value for each of the physical properties is taken to be the unweighted mean of the values found using the five different model sets. The random error is taken to be the largest of the individual errors found by the perturbation analysis. The systematic error is the standard deviation of the values from the five different model sets. The results for each model set are given in Table 3. The final physical properties of WASP-7 are collected in Table 4 and compared to those found by H09. In this table the quantity  $T'_{\text{eq}}$  is the equilibrium temperature of the planet excluding the energy redistribution factor

**Table 4.** Final physical properties of WASP-7 compared to those found by H09. The second errorbars represent systematic errors.

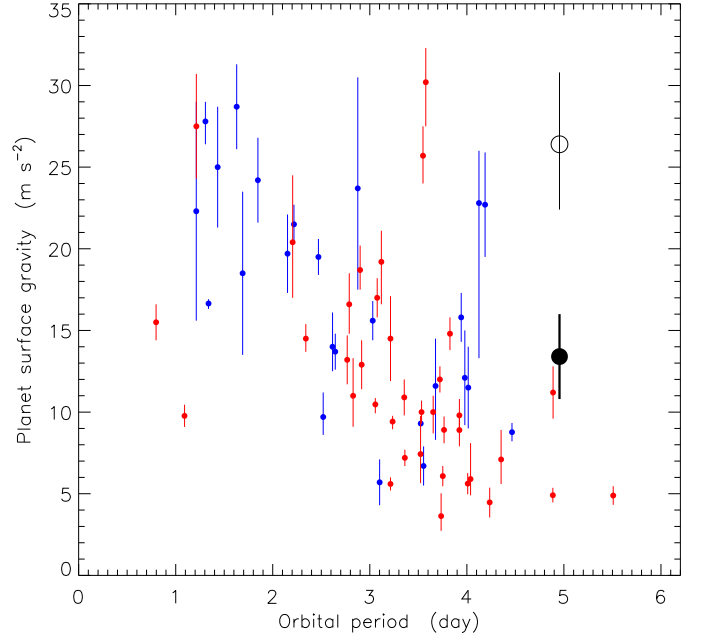
	This work (final)	H09
$M_A$ ( $M_\odot$ )	$1.276 \pm 0.061 \pm 0.022$	$1.28^{+0.09}_{-0.19}$
$R_A$ ( $R_\odot$ )	$1.432 \pm 0.092 \pm 0.008$	$1.236^{+0.059}_{-0.046}$
$\log g_A$ (cgs)	$4.232 \pm 0.047 \pm 0.003$	$4.363^{+0.010}_{-0.047}$
$\rho_A$ ( $\rho_\odot$ )	$0.434 \pm 0.074$	
$M_b$ ( $M_{\text{Jup}}$ )	$0.96 \pm 0.13 \pm 0.01$	$0.96^{+0.12}_{-0.18}$
$R_b$ ( $R_{\text{Jup}}$ )	$1.330 \pm 0.093 \pm 0.008$	$0.915^{+0.046}_{-0.040}$
$g_b$ ( $\text{m s}^{-1}$ )	$13.4 \pm 2.6$	$26.4^{+4.4}_{-4.0}$
$\rho_b$ ( $\rho_{\text{Jup}}$ )	$0.41 \pm 0.10$	$1.26^{+0.25}_{-0.21}$
$T'_{\text{eq}}$ (K)	$1487 \pm 48$	$1379^{+35}_{-23}$
$\Theta$	$0.070 \pm 0.011 \pm 0.000$	
$a$ (AU)	$0.0617 \pm 0.0010 \pm 0.0004$	$0.0618^{+0.0014}_{-0.0033}$
Age (Gyr)	$2.4^{+0.8+0.3}_{-1.0-0.4}$	

**Fig. 2.** Plot of the masses and radii of the known TEPs. The blue symbols denote values from the homogeneous analysis of Southworth (2010) and the red symbols results for the other known TEPs. WASP-7 is shown in black with an open circle (H09) and a filled circle (this work). Grey dotted lines show where density is 1.0, 0.5 and 0.25  $\rho_{\text{Jup}}$ .

(Southworth 2009) and  $\Theta$  is the Safronov (1972) number. We find that the star is slightly evolved, allowing a reasonable estimate of its age ( $2.4 \pm 1.0$  Gyr).

## 5. Summary and conclusions

We have observed a transit of the WASP-7 system using defocused photometry techniques, resulting in a light curve with an rms scatter of only 0.68 mmag. These data were modelled with the JKTEBOP code, and the results used to determine the physical properties of the planet and its host star. The higher quality of our photometry has allowed us to measure the system properties without constraining the host star to follow a main-sequence mass-radius relation. Compared to H09, we find both a larger radius for the host star and a larger ratio of the radii of the planet and star. This results in a sizeable increase in the planetary ra-

**Fig. 3.** Plot of the orbital periods and surface gravities of the known TEPs. The symbols are the same as for Fig. 2.

dus, from  $0.915^{+0.046}_{-0.040} R_{\text{Jup}}$  to  $1.330 \pm 0.093 R_{\text{Jup}}$ , which in turn means a lower surface gravity and mean density.

We have collected literature measurements of the physical properties of the 96 known TEPs as of 2010/11/18. For 30 of these objects we used the results of the homogeneous analysis performed by Southworth (2010). In Fig. 2 we plot their masses and radii, plus the values from WASP-7 found in this work (filled circle) and from H09 (open circle). The H09 results for WASP-7 b indicated that it was one of the densest known TEPs below  $\sim 2 M_{\text{Jup}}$ , along with the more recently-discovered planets CoRoT-13 b (Cabrera et al. 2010) and HAT-P-15 b (Kovács et al. 2010). The outlier status has now been lifted: our results place WASP-7 b in a well-populated part of the mass-radius diagram and demonstrate that high-quality data is required to obtain reliable measurements of the properties of TEPs.

Fig. 3 shows the surface gravities of the known TEPs as a function of their orbital periods. The existing planet population shows an inverse correlation between period and surface gravity (Southworth et al. 2007b), at least for the dominant population with periods  $\lesssim 10$  d and masses  $\lesssim 3 M_{\text{Jup}}$  (see also Fressin et al. 2009). The revised properties of WASP-7 b move it from outlier status to within the sprawl of parameter space occupied by the general planet population.

The theoretical models of irradiated gas giant planets Bodenheimer et al. (2003), Fortney et al. (2007) and Baraffe et al. (2008) predict radii of no more than  $1.16 R_{\text{Jup}}$  for a  $1.0 M_{\text{Jup}}$  planet with a range of chemical compositions and core masses, and without an arbitrary additional heating source. Our upward revision of the radius of WASP-7 b means that it no longer matches the predictions of these models at the  $2\sigma$  level. These conclusions could be strengthened by the provision of more radial velocity measurements as well as the acquisition of further photometric observations.

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## References

- Anderson, D. R., Hellier, C., Gillon, M., et al. 2010, *ApJ*, 709, 159  
 Baraffe, I., Chabrier, G., & Barman, T. 2008, *A&A*, 482, 315  
 Bodenheimer, P., Laughlin, G., & Lin, D. N. C. 2003, *ApJ*, 592, 555  
 Cabrera, J., Bruntt, H., Ollivier, M., et al. 2010, *A&A*, 522, A110  
 Claret, A. 2000, *A&A*, 363, 1081  
 Claret, A. 2004, *A&A*, 424, 919  
 Demarque, P., Woo, J.-H., Kim, Y.-C., & Yi, S. K. 2004, *ApJS*, 155, 667  
 Dotter, A., Chaboyer, B., Jevremović, D., et al. 2008, *ApJS*, 178, 89  
 Eastman, J., Siverd, R., & Gaudi, B. S. 2010, *PASP*, 122, 935  
 Fortney, J. J., Marley, M. S., & Barnes, J. W. 2007, *ApJ*, 659, 1661  
 Fressin, F., Guillot, T., & Nasta, L. 2009, *A&A*, 504, 605  
 Hellier, C., Anderson, D. R., Gillon, M., et al. 2009, *ApJ*, 690, L89  
 Jenkins, J. M., Caldwell, D. A., & Borucki, W. J. 2002, *ApJ*, 564, 495  
 Kovács, G., Bakos, G. Á., Hartman, J. D., et al. 2010, *ApJ*, 724, 866  
 Nelson, B. & Davis, W. D. 1972, *ApJ*, 174, 617  
 Pietrinferni, A., Cassisi, S., Salaris, M., & Castelli, F. 2004, *ApJ*, 612, 168  
 Pollacco, D. L., Skillen, I., Cameron, A. C., et al. 2006, *PASP*, 118, 1407  
 Popper, D. M. & Etzel, P. B. 1981, *AJ*, 86, 102  
 Safronov, V. S. 1972, *Evolution of the Protoplanetary Cloud and Formation of the Earth and Planets* (Jerusalem: Israel Program for Scientific Translation)  
 Southworth, J. 2008, *MNRAS*, 386, 1644  
 Southworth, J. 2009, *MNRAS*, 394, 272  
 Southworth, J. 2010, *MNRAS*, 408, 1689  
 Southworth, J., Bruntt, H., & Buzasi, D. L. 2007a, *A&A*, 467, 1215  
 Southworth, J., Hinse, T. C., Jørgensen, U. G., et al. 2009, *MNRAS*, 396, 1023  
 Southworth, J., Mancini, L., Calchi Novati, S., et al. 2010, *MNRAS*, 408, 1680  
 Southworth, J., Wheatley, P. J., & Sams, G. 2007b, *MNRAS*, 379, L11  
 Stetson, P. B. 1987, *PASP*, 99, 191  
 Van Hamme, W. 1993, *AJ*, 106, 2096  
 VandenBerg, D. A., Bergbusch, P. A., & Dowler, P. D. 2006, *ApJS*, 162, 375