

Multi-wavelength observations of the GRB 080319B afterglow and the modeling constraints*

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ABSTRACT

Context. We present observations of the GRB 080319B afterglow at optical, mm, and radio frequencies between a few hours and 67 days after the burst.

Aims. We attempt to understand the nature of this extraordinarily bright explosion based on the observed properties and its comparison with afterglow models.

Methods. Our observations and other published multiwavelength data were used to reconstruct the light curves and spectral energy distributions of the burst afterglow.

Results. Our results indicate that the observed features of the afterglow agrees equally well with the inter stellar matter and the stellar wind density profiles of the circumburst medium. In the case of both density profiles, the maximum synchrotron frequency ν_m is below optical value and the cooling break frequency ν_c is below X-rays, $\sim 10^4$ s after the burst. The derived value of the Lorentz factor at the time of naked-eye brightness is also ~ 300 for a corresponding blast-wave size of $\sim 10^{18}$ cm.

Conclusions. The numerical fit to the multiwavelength afterglow data constraints the values of physical parameters and the emission mechanism of the burst.

Key words. gamma rays: bursts

1. Introduction

The very early-time observations of gamma-ray bursts (GRBs) during both prompt emission or early afterglow phase have been one of the major contributions during the *Swift* era

* Based on observations obtained with the 0.22 m telescope at Russia the 0.7 m telescope at of Kharkov University, Ukraine, the 0.8 m telescope at Observatorio del Teide (IAC-80), Spain the 1.2 m Mercator telescope at La Palma, Spain, the 1.5 m telescope of Maidanak observatory Uzbekistan, the 2.0 m IGO Telescope at IUCAA Pune, India, the 2.5 m NOT, the PdB millimeter interferometric array France, the RATAN-600 Radio Telescope at Russia and the RT-22 radio telescope of CrAO, Ukraine.

(Zhang 2007). These observations play key role in determining the physical mechanisms underlying these energetic cosmic explosions (Mešařov 2006). The early-time multiwavelength observations are also very useful in constraining the afterglow models, and hence the nature of the possible progenitors and the ambient media surrounding the GRBs (Castro-Tirado et al. 1999; Piran 1999). Among the well-known examples of early observations in the pre-*Swift* era are GRB 990123 (Akerlof et al. 1999) and GRB 041219 (Vestrand et al. 2005) whereas GRB 050820A (Vestrand et al. 2006), GRB 060111B (Klotz et al. 2006), GRB 060210 (Stanek et al. 2007) and GRB 071010B (Wang et al. 2008) are good examples in the post-*Swift* era. The statistics of these examples of long-duration GRBs with very early-time

afterglow observations have improved because of the precise on-board localization by *Swift* and co-ordinated observations with the ground-based robotic optical telescopes (Gehrels et al. 2004; Gomboc et al. 2006). Out of these known examples, the observed optical and γ -ray prompt emission is contemporaneous for GRB 041219 and GRB 050820A, whereas for GRB 990123 and GRB 050401, the peak of the prompt optical observations were recorded after the γ -ray emission phase.

According to the standard relativistic “fireball” model (Rees & Mešzařos 1992; Mešzařos & Rees 1997; Panaiteescu et al. 1998), the GRB prompt emission is produced by internal shocks (Narayan et al. 1992; Rees & Mešzařos 1994; Sari & Piran 1997a,b). The observed steep decay in the early afterglow emission, which occurs between the prompt emission and the afterglow, are also explained in terms of “high-latitude emission” (Nousek et al. 2006) in the case of *Swift* GRBs, irrespective of the type of shock or the radiative process involved. The very early-time optical observations of several GRB afterglows are explained in terms of having a reverse shock (Kobayashi 2000) or/and forward shock (Rees & Mešzařos 1992; Katz 1994) origin. The overall afterglow behavior of long-duration GRBs, including the aforementioned examples, are explained in terms of the “collapse of very massive stars” i.e., so-called “collapsars” as the most favored progenitor (MacFadyen & Woosley 1999). By now, the majority of long-duration GRB afterglows have been explained in terms of constant ambient density i.e., inter stellar medium (*ISM*, $\rho \propto r^0$) models (Sari, et al. 1998; Wijers & Galama 1999; Panaiteescu & Kumar 2002), although a Stellar Wind Medium (*WM*, $\rho \propto r^{-2}$) profile (Chevalier & Lee 2000a,b; Li & Chevalier 2001) is the natural result of massive-star environments (Zhang 2007), where ρ and r are the ambient density and the distance from the center of the progenitor star, respectively. The value of ambient density is constrained by the parameters “number density” n and the “wind parameter” A_* , respectively for the *ISM* and *WM* models. There are also certain cases of long-duration GRB afterglows that have been explained in terms of both density profiles in the form of the transition from *WM* density profile at early times to *ISM* density profile at later epochs of observations, e.g., GRB 050904 (Gendre et al. 2007) and GRB 050319 (Kamble et al. 2007).

GRB 080319B was triggered (trigger = 306757) by *Swift*-BAT (15–350 keV) at $T_0 = 06:12:49$ UT on March 19, 2008 (Racusin et al. 2008a) and was simultaneously detected by the Konus-Wind (20 keV–15 MeV) satellite (Golenetskii et al. 2008). The optical emission started 2.75 ± 5 s after the BAT trigger, and was captured by the wide-field robotic telescope “Pi of the Sky” (Cwiok et al. 2007, 2008) and also TORTORA (Molinari et al. 2006; Racusin 2008b) and Raptor-Q (Wozniak et al. 2009) at later epochs. For a few seconds after the burst, the observed prompt optical flash of GRB 080319B was visible even to the unaided eye in dark skies (peaking at a visual magnitude of 5.3 around 18 s after the onset of the burst, observed by TORTORA), breaking the record of a handful of known cases in which bright optical-*NIR* prompt emissions have been observed (Jelinek et al. 2007). *Swift*-UVOT slewed towards the burst 51 s after the trigger (Holland & Racusin 2008) and later many other ground and space-based multiwavelength facilities joined in, as summarized in detail by Bloom et al. (2009), Racusin et al. (2008b), and references therein.

The measured redshift value of the burst ($z = 0.937$, Vreeswijk et al. 2008) corresponds to a luminosity distance of $d_L = 6.01 \times 10^3$ Mpc (for the cosmological parameters $H_0 = 71 \text{ km s}^{-1}\text{Mpc}^{-1}$, $\Omega_M = 0.27$, $\Omega_\Lambda = 0.73$) and an equivalent isotropic energy in γ -rays of $E_{\gamma,\text{iso}} = 1.4 \times 10^{54}$ erg

Table 1. Log of milli meter-wave and radio observations of the GRB 080319B afterglow.

Start time (UT)	End time (UT)	Frequency (GHz)	Flux center (mJy)	Telescope
2008 Mar. 20.0	Mar. 20.5	97.0	$+0.41 \pm 0.12$	IRAM
2008 Mar. 23.0	Mar. 23.5	97.0	$+0.35 \pm 0.19$	IRAM
2008 Mar. 27.0	Mar. 27.5	97.0	$+0.20 \pm 0.09$	IRAM
2008 Mar. 28.0	Mar. 28.5	4.8	<4.0	RATAN
2008 Mar. 30.0	Mar. 30.5	4.8	<4.0	RATAN
2008 Apr. 08.0	Apr. 08.5	4.8	<3.0	RATAN
2008 May 25.0	May 25.5	2.0	<3.0	RT-22
2008 May 25.0	May 25.5	8.0	<3.0	RT-22

(20 keV–7 MeV), which is among the highest energy ever measured for GRBs. The inferred high luminosity of the burst after correcting to the Galactic and host galaxy extinction in the burst direction (Wozniak et al. 2009) implies that such an event could easily have been detected at redshift $z = 17$ with meter-class telescopes in the *K*-band (Bloom et al. 2009; Sagar 2005). The observed value of the optical flash flux density of GRB 080319B (~ 20 Jy) is also about 4 orders of magnitude higher than the peak flux density in γ -rays (~ 14 mJy), the highest ever observed for GRBs (Yu et al. 2008). The significant excess of prompt optical flux compared to the extrapolated γ -ray spectrum implies that different emission components related to the two observed frequencies (Racusin et al. 2008b).

Thus, it is clear that multiwavelength afterglow observations of GRB 080319B provide unique opportunity to study the nature of this energetic cosmic explosion in detail. The observed spectral and temporal coverage of the afterglow imply that the burst is one of the most well-known examples for testing theoretical afterglow models (Pandey et al. 2003; Resmi et al. 2005). In this paper, we summarize the radio, milli meter-wave (mm), and optical observations of the afterglow in Sect. 2. The afterglow properties are discussed in Sect. 3. The modeling of the multiwavelength afterglow data and the derived parameters are described in Sect. 4. In the last section, we summarize our results in the context of observed and modeled parameters of the burst.

2. Observations

2.1. Milli meter-wave and radio observations

Observations of GRB 080319B afterglow at millimeter frequencies were performed with the Plateau de Bure (PdB) Interferometer (Guilloteau et al. 1992) in a six-antenna extended configuration on the dates listed in Table 1. The afterglow was detected in the first round of observations starting 0.5 days after the burst and upper limits were established at two later epochs. Radio observations were also performed using RATAN-600 at 4.8 GHz between 9 and 19 days after the burst and the afterglow was not detected. The afterglow was also monitored using the radio telescope RT-22/CraO at 2.0 and 8.0 GHz on 25th May 2008. The upper limits (3σ) inferred from RATAN-600 and RT-22/CraO observations are also listed in Table 1.

2.2. Optical observations

Our optical observations were performed using several telescopes (0.22 m SR-22; 0.7 m AZT-8; 0.8 m IAC; 1.2 m Mercator, 1.5 m Maidanak (AZT-22), 2.0 m IGO, 2.5 m NOT(+ALFOSC)) in the far east of Russia, Europe, Middle Asia and India starting from 0.5 day to 19 days after the burst. IRAF and DAOPHOT softwares were used to perform data reduction using standard techniques. The B , V , R , and I data were calibrated using the nearby secondary standards as given in Henden (2008). The Gunn r data were calibrated using nearby calibrators. The log of our optical observations along with the details are given in Table 2.

3. Properties of the burst

In this section, we discuss the observed temporal properties of the burst during the prompt emission and during the afterglow phase.

3.1. Prompt emission phase

The prompt optical light curve appears to be correlated with the γ -ray light curve, because of the high time-resolution (Racusin et al. 2008b). This correlation suggests that radiation at both frequencies originate in the same physical source based on the assumptions of internal shock model (Mešzařos & Rees 1999; Kumar & Panaiteescu 2008). There is no considerable evolution in the X-ray hardness ratio from very early to late times of the observations with an average photon index $\Gamma = -1.81 \pm 0.04$ (Bloom et al. 2009). There are no multi-band optical observations during the prompt emission phase of the burst. The observed spectral indices around 430 s and 875 s after the burst also show negligible evolution at optical frequencies (Bloom et al. 2009). The reverse shock or forward shock origin of the observed prompt optical emission is excluded based on the observed values of temporal decay indices from being too steep than expected (Sari, et al. 1998; Kobayashi 2000). The observed constancy of the optical pulse width with time during the prompt emission phase (Ramirez-Ruiz & Fenimore 2000) also does not support the reverse shock origin. The temporal coincidence between the prompt optical and γ -rays of the burst also indicate that they might have originated in the same emitting region or two regions sharing the same dynamical behavior (Yu et al. 2008). The correlation of spectral lag evolution with the observed prompt optical emission of GRB 080319B also shows that they originate from the same astrophysical origin and that the respective radiation mechanisms were dynamically coupled (Stamatikos et al. 2009). Furthermore, the significant excess of the prompt optical flux in comparison to the γ -ray spectrum extrapolated to optical frequencies, is possibly indicative of two different emission components for the two observed frequencies (but see Kumar & Panaiteescu 2008).

3.2. Multi-frequency evolution of the afterglow

According to the standard fireball model, the GRB afterglow is expected to be synchrotron radiation with the observed flux $f \propto \nu^\beta t^\alpha$, for the regions without spectral breaks, where the values of power-law temporal decay index (α) and spectral index (β) are related to each other and can be used to understand the afterglow evolution (Sari, et al. 1998; Sari, et al. 1999). The rich multi-frequency data of the afterglow of GRB 080319B available in the literature and the data from the present work are

used to measure the indices α and β of the afterglow. The values of the temporal decay indices have been derived using empirical broken power-law relations by minimization of χ^2 as described in Granot & Sari (2002) (their Eqs. (1) and (4)). The bright optical flash ($10 < t < 100$ s) rises to its peak brightness ($V \sim 5.32 \pm 0.04$) with $\alpha_{V1} = 4.64 \pm 0.67$, and after about 50 s decays with $\alpha_{V2} = -4.41 \pm 0.10$. After ~ 100 s, the decay of the optical afterglow could be described as a broken power-law with $\alpha_{V3} = -2.33 \pm 0.23$, $\alpha_{V4} = -1.31 \pm 0.02$ with the break occurring at $\sim 784.9 \pm 304.4$ s after the burst. The X-ray afterglow light curve is described by a triple broken power-law with $\alpha_{X1} = -1.40 \pm 0.01$, $\alpha_{X2} = -1.94 \pm 0.12$, $\alpha_{X3} = -1.14 \pm 0.09$, and $\alpha_{X4} = -2.67 \pm 0.74$ and the breaks happening at $t_{b,X1} = 2583.39 \pm 871$ s, $t_{b,X2} = (3.99 \pm 0.89) \times 10^4$ s, and $t_{b,X3} = (1.0 \pm 0.29) \times 10^6$ s with $\chi^2/dof = 975/710 = 1.37$. The derived value of the X-ray temporal index α_{X4} and the corresponding break time are in agreement with those derived by Tanvir et al. (2009) using late-time X-ray and optical data, which provide evidence in favor of jet-break around $\sim 10^6$ s after the burst.

4. Discussion

In this section, we discuss about the burst ambient media and the numerical fit to the afterglow data.

4.1. Surrounding medium

The closure relations between α and β can also be used to infer the density profile of the circumburst medium or to distinguish between theoretical afterglow models such as *ISM* and *WM* (Price et al. 2002; Starling et al. 2008). For radiation produced by a shock interacting with the *ISM* circumburst medium, the expected closure relation is $\alpha - 1.5\beta = 0.0$ in a spectral regime $\nu_m < \nu < \nu_c$, where ν_m and ν_c are the maximum synchrotron and cooling frequencies, respectively. In the case of the *WM* circumburst medium and a spectral regime $\nu_m < \nu < \nu_c$, the closure relation is $\alpha - 1.5\beta = 0.5$. If $\nu_m < \nu_c < \nu$, then $\alpha - 1.5\beta = -0.5$ is expected irrespective of whether the density profile is *ISM* or *WM*. At early times ($t \leq 2000$ s), the optical afterglow is dominated by emission caused by the reverse shock, and at late times ($t \geq 10^5$ s) it is clearly dominated by that due to the forward shock. The values of β estimated by Bloom et al. (2009) and the values of α inferred from optical light curves at $t > 10^5$ s are consistent with the closure relations for both the density profiles, *ISM* and *WM* in the spectral regime $\nu_m < \nu_{\text{optical}} < \nu_c$.

It is also clear that the temporal decay index of the optical afterglow between 100 s and 1000 s ($\alpha \sim -2.3$) is higher than expected for a forward shock scenario ($\alpha \sim -0.8$) interacting with the circumburst medium (Sari, et al. 1998). However, this steep decay would be expected in the case of a reverse shock interacting with the ejected shell. The radiation produced by the interaction of the forward shock with the circumburst medium would then dominate the afterglow after the radiation from reverse shock has diminished considerably, which in this case could happen after 10^5 s. Thus, this optical afterglow could be explained using the reverse-shock and forward-shock interactions with the circumburst medium.

4.2. Numerical model fits to the data

We fitted the afterglow using both a reverse shock (RS) model and a forward shock (FS) model, and assumed that the

Table 2. Log of optical-NIR observations of the afterglow of GRB 080319B.

<i>UT</i> start [d]	<i>T</i> – <i>T</i> ₀ [d]	<i>T</i> _{exp}	Filter	mag	err _{mag}	Telescope
2008 Mar. 19.4616	0.2027	1200 s	none	>15.5		0.22 m SR-22
2008 Mar. 19.7642	0.5052	600 s	<i>R</i>	19.07	0.19	0.7 m AZT-8
2008 Mar. 19.7721	0.5132	600 s	<i>R</i>	19.26	0.27	0.7 m AZT-8
2008 Mar. 19.7800	0.5211	600 s	<i>R</i>	19.74	0.50	0.7 m AZT-8
2008 Mar. 19.7911	0.5322	600 s	<i>R</i>	19.96	0.39	0.7 m AZT-8
2008 Mar. 19.7974	0.5385	600 s	<i>R</i>	19.60	0.34	0.7 m AZT-8
2008 Mar. 19.8055	0.5466	900 s	<i>R</i>	19.60	0.20	0.7 m AZT-8
2008 Mar. 19.8174	0.5585	900 s	<i>R</i>	19.40	0.15	0.7 m AZT-8
2008 Mar. 19.8285	0.5696	900 s	<i>R</i>	19.89	0.28	0.7 m AZT-8
2008 Mar. 19.8470	0.5881	900 s	<i>R</i>	19.82	0.23	0.7 m AZT-8
2008 Mar. 19.8557	0.5968	900 s	<i>R</i>	20.06	0.19	0.7 m AZT-8
2008 Mar. 19.8630	0.6041	900 s	<i>R</i>	19.88	0.23	0.7 m AZT-8
2008 Mar. 19.8888	0.6299	720 s	<i>R</i>	20.00	0.27	0.7 m AZT-8
2008 Mar. 19.8980	0.6391	720 s	<i>R</i>	19.52	0.16	0.7 m AZT-8
2008 Mar. 19.9071	0.6482	720 s	<i>R</i>	20.09	0.28	0.7 m AZT-8
2008 Mar. 19.9139	0.6550	720 s	<i>R</i>	19.80	0.24	0.7 m AZT-8
2008 Mar. 19.9234	0.6645	900 s	<i>R</i>	19.88	0.27	0.7 m AZT-8
2008 Mar. 19.9354	0.6765	900 s	<i>R</i>	20.19	0.40	0.7 m AZT-8
2008 Mar. 19.9497	0.6908	900 s	<i>R</i>	19.54	0.15	0.7 m AZT-8
2008 Mar. 19.9625	0.7036	900 s	<i>R</i>	20.16	0.44	0.7 m AZT-8
2008 Mar. 19.9752	0.7163	720 s	<i>R</i>	20.47	0.25	0.7 m AZT-8
2008 Mar. 19.9843	0.7254	960 s	<i>R</i>	20.00	0.37	0.7 m AZT-8
2008 Mar. 19.9959	0.7370	900 s	<i>R</i>	20.03	0.37	0.7 m AZT-8
2008 Mar. 20.0127	0.7538	720 s	<i>R</i>	20.01	0.22	0.7 m AZT-8
2008 Mar. 20.0196	0.7607	720 s	<i>R</i>	19.83	0.21	0.7 m AZT-8
2008 Mar. 20.0664	0.8075	720 s	<i>R</i>	19.94	0.23	0.7 m AZT-8
2008 Mar. 20.0755	0.8166	720 s	<i>R</i>	20.22	0.26	0.7 m AZT-8
2008 Mar. 20.0848	0.8259	900 s	<i>R</i>	20.28	0.24	0.7 m AZT-8
2008 Mar. 20.0959	0.8370	900 s	<i>R</i>	19.74	0.16	0.7 m AZT-8
2008 Mar. 20.1069	0.8480	600 s	<i>R</i>	19.56	0.20	0.7 m AZT-8
2008 Mar. 19.8319	0.5730	600 s	<i>V</i>	19.60	0.70	0.7 m AZT-8
2008 Mar. 20.0870	0.8351	1200 s	<i>I</i>	19.57	0.08	0.8 m IAC 80
2008 Mar. 19.9924	0.7457	4 × 300 s	<i>R</i>	19.82	0.07	0.8 m IAC 80
2008 Mar. 20.0557	0.8003	600 s	<i>R</i>	19.94	0.11	0.8 m IAC 80
2008 Mar. 20.1067	0.8548	2 × 600 s	<i>R</i>	20.15	0.08	0.8 m IAC 80
2008 Mar. 20.1294	0.8778	2 × 600 s	<i>R</i>	20.14	0.09	0.8 m IAC 80
2008 Mar. 20.2454	0.9972	3 × 600 s	<i>R</i>	20.50	0.08	0.8 m IAC 80
2008 Mar. 22.0730	2.8355	6 × 600 s	<i>R</i>	22.03	0.27	0.8 m IAC 80
2008 Mar. 20.0252	0.7814	(1200+600) s	<i>V</i>	20.39	0.10	0.8 m IAC 80
2008 Mar. 20.0648	0.8165	3 × 600 s	<i>B</i>	20.89	0.12	0.8 m IAC 80
2008 Mar. 20.1703	0.9149	600 s	<i>B</i>	20.75	0.20	0.8 m IAC 80
2008 Mar. 19.9361	0.6985	120 s	<i>r</i>	19.56	0.17	1.2 m Mercator
2008 Mar. 19.9613	0.7031	120 s	<i>r</i>	19.28	0.19	1.2 m Mercator
2008 Mar. 19.9631	0.7049	120 s	<i>r</i>	19.74	0.15	1.2 m Mercator
2008 Mar. 19.9649	0.7067	120 s	<i>r</i>	19.37	0.11	1.2 m Mercator
2008 Mar. 19.9667	0.7085	120 s	<i>r</i>	19.93	0.15	1.2 m Mercator
2008 Mar. 19.9681	0.7095	60 s	<i>r</i>	19.53	0.11	1.2 m Mercator
2008 Mar. 19.9691	0.7106	60 s	<i>r</i>	20.07	0.21	1.2 m Mercator
2008 Mar. 19.9702	0.7117	60 s	<i>r</i>	19.41	0.10	1.2 m Mercator
2008 Mar. 19.9713	0.7127	60 s	<i>r</i>	19.66	0.20	1.2 m Mercator
2008 Mar. 19.9723	0.7138	60 s	<i>r</i>	19.82	0.15	1.2 m Mercator
2008 Mar. 19.9734	0.7148	60 s	<i>r</i>	19.68	0.09	1.2 m Mercator
2008 Mar. 19.9745	0.7159	60 s	<i>r</i>	19.61	0.09	1.2 m Mercator
2008 Mar. 19.9759	0.7177	120 s	<i>r</i>	19.57	0.12	1.2 m Mercator
2008 Mar. 19.9777	0.7195	120 s	<i>r</i>	19.54	0.09	1.2 m Mercator
2008 Mar. 19.9794	0.7212	120 s	<i>r</i>	19.62	0.07	1.2 m Mercator
2008 Mar. 19.9812	0.7230	120 s	<i>r</i>	19.83	0.09	1.2 m Mercator
2008 Mar. 19.9830	0.7248	120 s	<i>r</i>	19.58	0.06	1.2 m Mercator
2008 Mar. 19.9848	0.7266	120 s	<i>r</i>	19.65	0.07	1.2 m Mercator
2008 Mar. 19.9865	0.7283	120 s	<i>r</i>	19.67	0.07	1.2 m Mercator

Table 2. continued.

UT start [d]	$T - T_0$ [d]	T_{exp}	Filter	mag	err _{mag}	Telescope
2008 Mar. 19.9883	0.7301	120 s	<i>r</i>	19.78	0.08	1.2 m Mercator
2008 Mar. 19.9900	0.7318	120 s	<i>r</i>	19.51	0.07	1.2 m Mercator
2008 Mar. 19.9918	0.7336	120 s	<i>r</i>	19.66	0.08	1.2 m Mercator
2008 Mar. 19.9936	0.7354	120 s	<i>r</i>	19.47	0.05	1.2 m Mercator
2008 Mar. 19.9953	0.7371	120 s	<i>r</i>	19.62	0.06	1.2 m Mercator
2008 Mar. 19.9971	0.7389	120 s	<i>r</i>	19.46	0.05	1.2 m Mercator
2008 Mar. 19.9989	0.7407	120 s	<i>r</i>	19.62	0.06	1.2 m Mercator
2008 Mar. 20.1446	0.8871	120 s	<i>r</i>	19.76	0.07	1.2 m Mercator
2008 Mar. 20.1464	0.8888	120 s	<i>r</i>	19.95	0.08	1.2 m Mercator
2008 Mar. 20.1481	0.8906	120 s	<i>r</i>	20.14	0.10	1.2 m Mercator
2008 Mar. 20.1500	0.8925	120 s	<i>r</i>	19.78	0.07	1.2 m Mercator
2008 Mar. 20.1517	0.8942	120 s	<i>r</i>	20.04	0.11	1.2 m Mercator
2008 Mar. 20.1535	0.8960	120 s	<i>r</i>	19.98	0.10	1.2 m Mercator
2008 Mar. 20.1554	0.8979	120 s	<i>r</i>	19.85	0.08	1.2 m Mercator
2008 Mar. 20.1572	0.8997	120 s	<i>r</i>	19.85	0.12	1.2 m Mercator
2008 Mar. 20.1589	0.9014	120 s	<i>r</i>	19.87	0.13	1.2 m Mercator
2008 Mar. 20.1607	0.9032	120 s	<i>r</i>	19.87	0.11	1.2 m Mercator
2008 Mar. 20.1625	0.9049	120 s	<i>r</i>	19.92	0.09	1.2 m Mercator
2008 Mar. 20.1642	0.9067	120 s	<i>r</i>	19.96	0.11	1.2 m Mercator
2008 Mar. 20.1660	0.9085	120 s	<i>r</i>	20.04	0.10	1.2 m Mercator
2008 Mar. 20.1678	0.9103	120 s	<i>r</i>	19.84	0.08	1.2 m Mercator
2008 Mar. 20.1695	0.9120	120 s	<i>r</i>	19.73	0.09	1.2 m Mercator
2008 Mar. 20.1713	0.9138	120 s	<i>r</i>	19.84	0.11	1.2 m Mercator
2008 Mar. 20.1731	0.9156	120 s	<i>r</i>	19.92	0.10	1.2 m Mercator
2008 Mar. 20.1748	0.9173	120 s	<i>r</i>	20.01	0.12	1.2 m Mercator
2008 Mar. 20.1766	0.9191	120 s	<i>r</i>	19.91	0.10	1.2 m Mercator
2008 Mar. 20.1784	0.9208	120 s	<i>r</i>	20.08	0.09	1.2 m Mercator
2008 Mar. 20.1801	0.9226	120 s	<i>r</i>	19.71	0.07	1.2 m Mercator
2008 Mar. 20.1819	0.9244	120 s	<i>r</i>	19.87	0.09	1.2 m Mercator
2008 Mar. 20.1837	0.9261	120 s	<i>r</i>	19.83	0.07	1.2 m Mercator
2008 Mar. 20.1854	0.9279	120 s	<i>r</i>	19.84	0.09	1.2 m Mercator
2008 Mar. 21.8618	2.6029	5760 s	<i>I</i>	21.22	0.15	1.5 m AZT-22
2008 Mar. 21.8781	2.6192	5610 s	<i>R</i>	21.81	0.26	1.5 m AZT-22
2008 Apr. 07.0086	18.7479	7200 s	<i>R</i>	>23.6		1.5 m AZT-22
2008 Mar. 19.9179	0.6590	3 × 600 s	<i>R</i>	19.66	0.09	2.0 m IGO
2008 Mar. 27.9511	8.6524	28 × 120 s	<i>R</i>	>23.0		2.5 m NOT

radiation mechanism is synchrotron following the standard Fireball scenario (Rees & Mešzařos 1992; Mešzařos & Rees 1994; Kobayashi 2000). As shown in Fig. 1, the multi-band optical afterglow is reproduced well by this model, whereas X-ray observations cannot be reproduced by the model. Our best-fit model is consistent with no jet break in the optical afterglow, and introducing a jet break at $t \sim 10^6$ s to coincide with the break in the X-ray afterglow does not change the fits significantly.

The complicated behavior of the X-ray afterglow was explained by Racusin et al. (2008b) using a two-component jet model – a central narrow jet surrounded by a coaxial wider jet. In this model, the X-ray afterglow from the central narrow jet dominates at early times, until about 2×10^4 s, when it fades below the brightness of the surrounding wide jet. In this model, the breaks in the X-ray afterglow light curves at times $t_{\text{b},\text{X}1} \sim 2 \times 10^3$ s and $t_{\text{b},\text{X}3} \sim 10^6$ s correspond to the jet breaks caused by the lateral spread of narrow and wide jets, respectively. We also attempted to reproduce the entire data set using this two-component-jet model, but found that with this model the optical radiation from the narrow jet would be brighter than the observed optical afterglow (see Fig. 2). There is no possible way to suppress this optical radiation from the narrow jet. The dominance of RS over

FS contribution at early times of the light curves are also shown for optical-NIR frequencies in Fig. 3 for both the models.

Alternatively, it is possible that the X-ray and optical afterglows of GRB 080319B are not related to each other and may have independent origins. The X-ray afterglow predicted by the RS-FS model is also fainter than the observed one and requires an additional component.

Assuming that the shock expands into the circum-burst medium, which has a *WM* density profile, our best-fit model spectral parameters at epoch $t = 10^4$ s imply that the peak of the synchrotron spectrum is below the optical bands ($\nu_m < 1.3 \times 10^{14}$ Hz) with corresponding normalization flux ($F_{\nu_m} < 607 \mu\text{Jy}$) and the self-absorption frequency, $\nu_a > 10^8$ Hz. The index of electron energy distribution $p \sim 2.07$ is lower than the canonical value $p \sim 2.3$, which has often been observed (Panaitescu & Kumar 2001, 2002; Starling et al. 2008). The location of the cooling break frequency can be constrained to be below the X-ray band $\nu_c \leq 10^{18}$ Hz, which would be consistent with the observed X-ray afterglow at late times ($t > 10^5$ s). However, if the origin of the X-ray afterglow is independent of the optical afterglow, then $\nu_c \ll 10^{18}$ Hz.

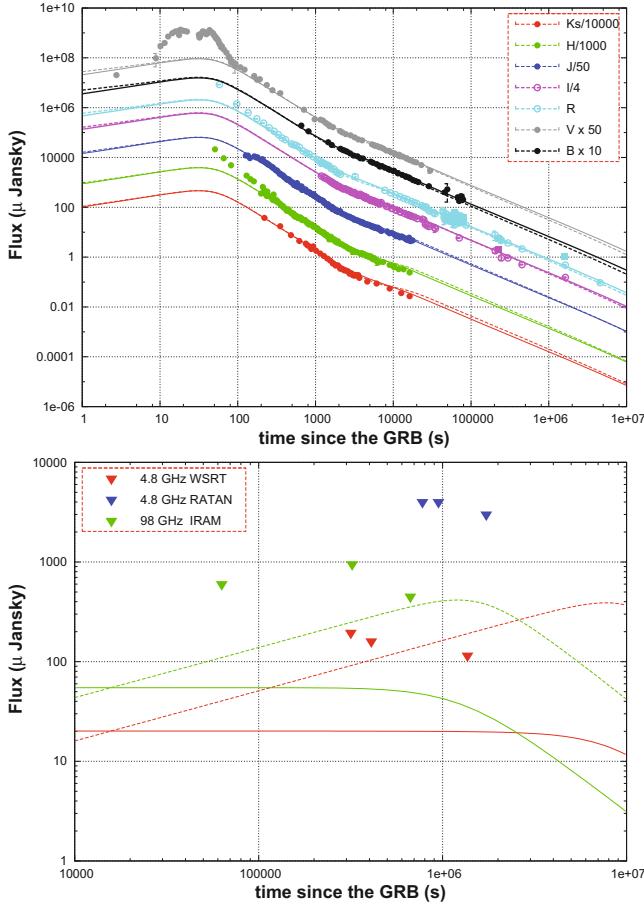


Fig. 1. Light curves of the GRB 080319B afterglow: the optical-NIR and X-ray afterglow is shown in the *top panel* and the radio data in the *bottom panel*. The present optical observations are plotted along with other published optical and X-ray data. The FS-RS model is used to fit the afterglow light curves. The solid lines in the plots represent the resultant best fit *WM* model whereas dotted line show the best fit *ISM* model.

The estimated isotropic equivalent afterglow kinetic energy (E_K^{iso}) released in this explosion turns out to be $<5.5 \times 10^{53}$ erg. For the assumed *WM* circumburst medium, we find the parameter value $A_* > 0.01$. The fraction of the total energy given to accelerating electrons (ϵ_e) and to the magnetic fields (ϵ_B) turns out to be about >0.41 and $<3 \times 10^{-3}$, respectively. The absence of any jet break until about 10^6 s imply a jet opening angle >1.4 degrees, and hence the true amount of released energy must be $>8 \times 10^{49}$ erg. The radius of the fireball at 10^4 s was $<5 \times 10^{18}$ cm, which we extrapolated back in time to estimate the physical size of the emitting region during the optical flash ($t \sim 50$ s) and found to be $<3.6 \times 10^{17}$ cm. Similarly, the Lorentz factor of the blast-wave was <64 at 10^4 s after the burst, which corresponds to the Lorentz factor of about <240 at the time of the optical flash. These findings, which are only based on the evolution in the afterglow, are comparable with those of Racusin et al. (2008b) and Kumar & Panaiteescu (2008), who used arguments based on prompt emission to reach their conclusions. It is reassuring that the different approaches measured similar fireball sizes and blast-wave Lorentz factors.

The afterglow can also be explained by assuming *ISM* density profile of the circumburst medium. In this case, the best-fit model spectral parameters at epoch $t = 10^4$ s were found to be $\nu_m < 2.1 \times 10^{14}$ Hz with corresponding flux $F_{\nu_m} < 566 \mu\text{Jy}$,

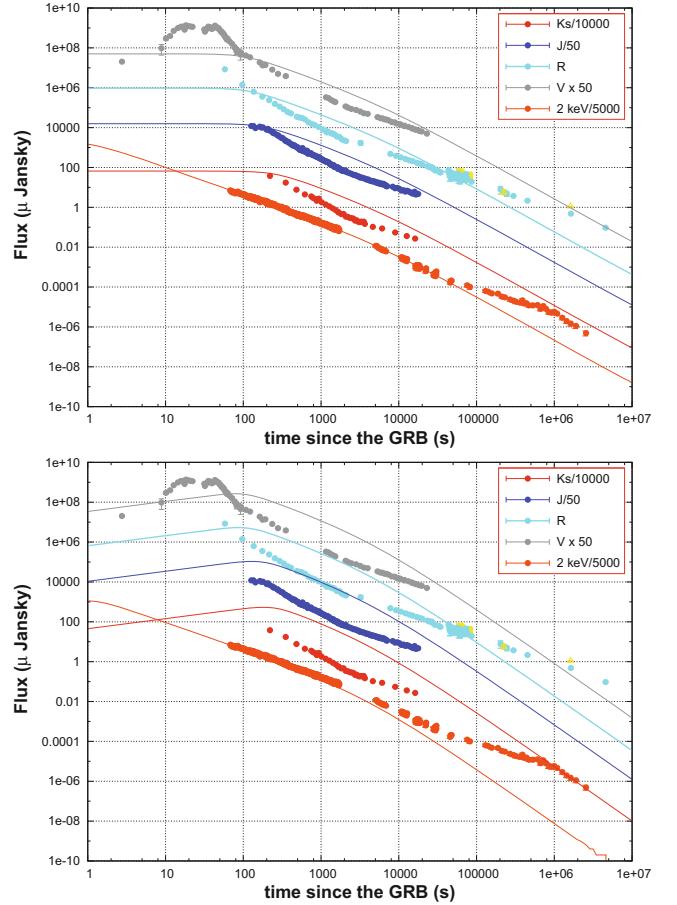


Fig. 2. The afterglow of GRB 080319B – a comparison with the double-jet model: the model light curves shown in *top and bottom panels* are plotted for the *WM* and *ISM* cases, respectively. The model here assumes the double-jet scenario. The contribution to the afterglow only due to the narrow jet is shown in the figure. The narrow jet is assumed to be responsible for the early X-ray afterglow.

and the self-absorption frequency, $\nu_a > 10^8$ Hz. To explain the relatively steep decay of the afterglow, the required value of $p = 2.71$ is on a high side, and $\nu_c \leq 10^{18}$ Hz, as in the case of a *WM* density profile.

The physical parameters estimated using these best-fit spectral parameters are $E_K^{\text{iso}} < 2.4 \times 10^{52}$ erg, $\epsilon_e > 0.14$, $\epsilon_B < 0.02$, and $n > 7.5 \times 10^{-4}$ atoms/cc. The absence of any jet break until about 10^6 s means that the jet opening angle of $\theta_j > 4.6$ degrees, and hence the beaming corrected afterglow kinetic energy is $E_K^{\text{corr}} > 3.8 \times 10^{49}$ erg. The radius of the fireball at 10^4 s turns out to be about $<2.1 \times 10^{18}$ cm and at the time of the optical flash, $t \sim 50$ s, is about $<5.5 \times 10^{17}$ cm. The corresponding Lorentz factors are <52.4 and <382 , at 10^4 and 50 s after the burst, respectively. As before, these values are also comparable to those of Racusin et al. (2008) and Kumar & Panaiteescu (2008).

5. Conclusions

We have presented observations of the GRB 080319B afterglow at radio, mm, and optical frequencies. The simultaneous multi-band afterglow modeling is useful in constraining the nature of the burst emission mechanism and the ambient medium. We found that the GRB 080319B afterglow is consistent with the fireball expanding either into the *ISM* or into the *WM*

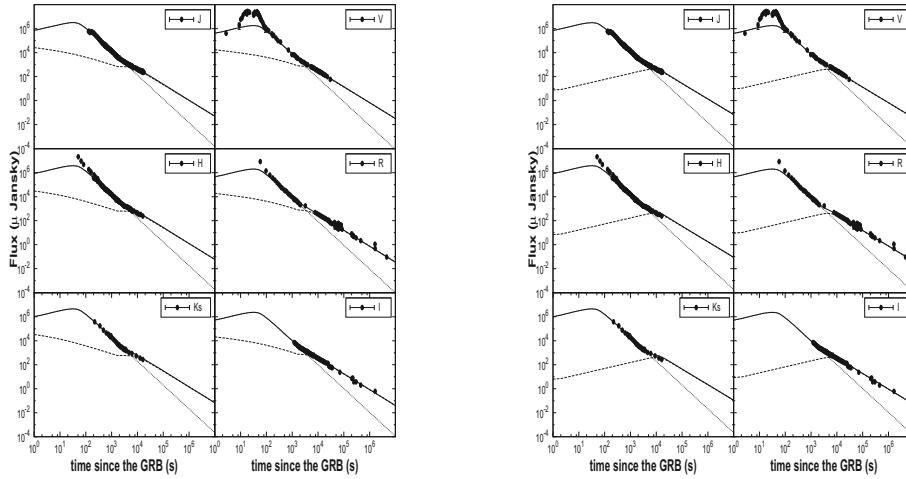


Fig. 3. The contributions by the RS and FS to the GRB 080319B afterglow at optical-NIR bands are represented by using the dashed and dotted lines, respectively. The addition of the two components is shown using the solid line. The *left and right panels* are WM and ISM models, respectively.

circumburst medium. We also excluded the double-jet model as a correct explanation of the observed multi-band behavior of the afterglow. Based on the assumptions of our model, we were able to explain the multi-band optical and radio afterglow reasonably well, but were unable to explain the X-ray afterglow. It is possible alternatively, that the optical and X-ray afterglows have different origins. The Lorentz factor of the shock was estimated from the afterglow evolution. The Lorentz factor at the time of naked-eye brightness, extrapolated from the late estimation, was found to be ~ 300 . The corresponding radius of the shock front is about 10^{17} cm. We also showed that the early peak brightness of the afterglow could not be caused the reverse shock. An additional emission mechanism, such as internal shocks, is required. Our results also indicate that existing blast-wave afterglow models should be modified in the light of the complicated behavior inferred for observed afterglows. In the future, observations of prompt optical spectra will be very useful in understanding the very early part of the afterglows of GRBs.

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References

Akerlof, C., Balsano, R., Barthelmy, S., et al. 1999, *Nature*, 398, 400
 Band, D., Matteson, J., Ford, L., et al. 1993, *ApJ*, 413, 281
 Bloom, J., Perley, D. A., Li, W., et al. 2009, *ApJ*, 691, 723
 Castro-Tirado, Alberto, J., Rosa Zapatero-Osorio, M., Caon, N., 1999, *Science*, 283, 2069
 Chevalier, R. A., & Li, Z. 2000a, *ApJ*, 520, L29
 Chevalier, R. A., & Li, Z. 2000b, *ApJ*, 536, 195
 Cwiok, M., Dominik, W., MaÅek, K., et al. 2007, *Ap&SS*, 309, 531
 Cwiok, M., Dominik, W., Kasprowicz, G., et al. 2008, *GCNC*, 7439
 Gehrels, N., Chincarini, G., & Giommi, P. 2004, *ApJ*, 611, L1005
 Gendre, B., Galli, A., Corsi, A., et al. 2007, *A&A*, 462, 565
 Golenetskii, S., Aptekar, R., Mazets, E., et al. 2008, *GCNC*, 7482
 Gomboc, A., Guidorzi, C., Mundell, C. G. 2006, *NCimB*, 121, 1303
 Granot, J., & Sari, R. 2002, *ApJ*, 568, 820
 Guilloteau, S., Delannoy, J., Downes, D., et al. 1992, *A&A*, 262, 624
 Henden, A. 2008, *GCNC*, 7528
 Holland, S. T., & Racusin, J. L. 2008, *GCNC*, 7496
 Jelinek, M., Prouza, M., Kubanek, P., et al. 2006, *A&A*, 454, L119
 Kamble, A. P., Resmi, L., & Misra, P. 2007, *ApJ*, 664, L5
 Katz, J. I. 1994, *ApJ*, 422, 248
 Klotz, A., Gendre, B., Stratta, G., et al. 2006, *A&A*, 451, L39
 Kobayashi, S. 2000, *ApJ*, 545, 507
 Kumar, P., & Panaiteescu, A. 2008, *MNRAS*, 391, L19
 Li, Z., & Chevalier, R. A. 2001, *ApJ*, 551, 940
 MacFadyan, A. I., & Woosley, S. E. 1999, *ApJ*, 524, 262
 MeÅafos, P. 2006, *Rep. Prog. Phys.*, 69, 2259
 MeÅafos, P., & Rees, M. J. 1994, *MNRAS*, 269, L41
 MeÅafos, P., & Rees, M. J. 1997, *ApJ*, 476, 232
 MeÅafos, P., & Rees, M. J. 1999, *MNRAS*, 306, L39
 Molinari, E., Bondar, S., Karpov, S., et al. 2006, *Nuovo Cimento*, 121, 1525
 Nousek, J., Kouveliotou, C., Grupe, D., et al. 2006, *ApJ*, 642, 389
 Panaiteescu, A., & Kumar P. 2001, *ApJ*, 560, L49
 Panaiteescu, A., & Kumar P. 2002, *ApJ*, 571, 779
 Panaiteescu, A., MeÅafos, P., & Rees, M. J. 1998, *ApJ*, 503, 314
 Pandey, S. B., Sahu, D. K., Resmi, L., et al. 2003, *BASI*, 31, 19
 Price, P. A., Berger, E., Reichart, D. E., et al. 2002, *ApJ*, 572, L51
 Piran, T. 1999, *Phys. Rep.*, 314, 575
 Racusin, J. L., Gehrels, N., Holland, S. T., et al. 2008a, *GCNC*, 7427
 Racusin, J. L., Karpov, S. V., Sokolowski, M., et al. 2008b, *Nature*, 455, 183
 Ramirez-Ruiz, E., & Fenimore, E. E. 2000, *ApJ*, 539, 712
 Rees, M. J., & MeÅafos, P. 1992, *MNRAS*, 258, L41
 Resmi, L., Ishwara, C. H., Castro-Tirado, A. J., et al. 2005, *A&A*, 440, 477
 Sagar, R. 2005, *BASI*, 33, 352
 Sari, R., & Piran, T. 1997a, *ApJ*, 485, 270
 Sari, R., & Piran, T. 1997b, *MNRAS*, 287, 110
 Sari, R., Piran, T., & Narayan, R., 1998, *ApJ*, 497, L17
 Sari, R., Piran, T., & Halpern, J. P. 1999, *ApJ*, 519 L17
 Stamatikos, M., Ukwatta, T. N., Sakamoto, T., et al. 2009
 [[arXiv:0809.2132v1](https://arxiv.org/abs/0809.2132v1)]
 Stanek, K. J., Dai, X., Prieto, J. L., et al. 2007, *ApJ*, L21
 Starling, R. L. C., Van Der Horst, A. J., Rol, E., et al. 2008, *ApJ*, 672, 433
 Tanvir, N. R., Evert, R., Andrew, L., et al. 2009, *ApJ*, submitted
 [[arXiv:0812.1217](https://arxiv.org/abs/0812.1217)]
 Vestrand, W. T., Wozniak, P. R., Wren, J. A., et al. 2005, *Nature*, 435, 178
 Vestrand, W. T., Wren, J. A., Wozniak, P. R., et al. 2006, *Nature*, 442, 172
 Vreeswijk, P. M., Smette, A., Malesani, D., et al. 2008, *GCNC* 7451
 Wang, J. H., Schwamb, M. E., Huang, K. Y., et al. 2008, *ApJ*, 679, L5
 Wijers, R. A. M. J., & Galama, T. J., 1999, *ApJ*, 523, 177
 Wozniak, P. R., Vestrand, W. T., Panaiteescu, A., et al. 2009, *ApJ*, 691, 495
 Yu, Y. W., Wang, X. Y., & Dai, Z. G. 2008 [[arXiv:0806.2010](https://arxiv.org/abs/0806.2010)]
 Zhang, B. 2007, *Chin. J. Astron. Astrophys.*, 7, 1