

Neural Ratio Estimation For Simulation-based Inference

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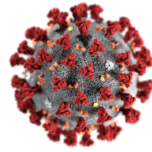
Simulation-based inference



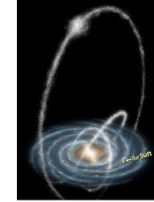
Chemical reactions



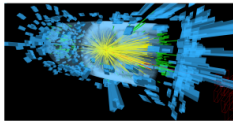
Flames



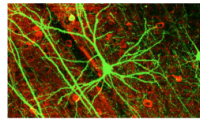
Epidemics



Stellar streams



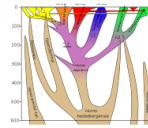
Collider experiments



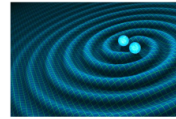
Neurons



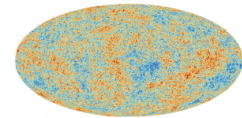
Robotics



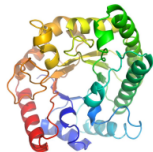
Evolution



Gravitational waves



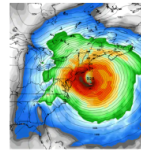
Evolution of the Universe



Protein networks



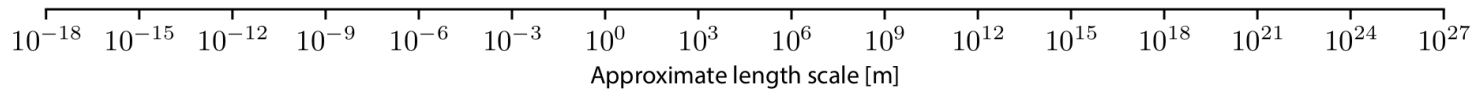
Ecological systems

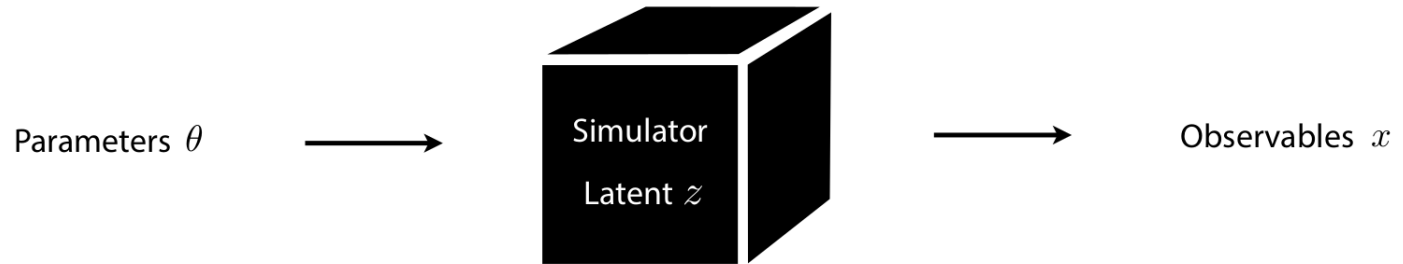


Weather and climate



Gravitational lensing





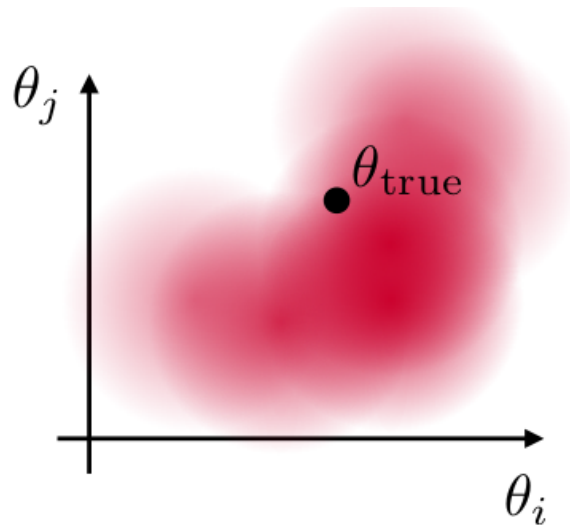
- Prediction:
- Well-motivated mechanistic, causal model
 - Simulator can generate samples $x \sim p(x|\theta)$

- Inference:
- Interactions between low-level components lead to challenging inverse problems
 - Likelihood $p(x|\theta) = \int dz p(x, z|\theta)$ is intractable

Start with

- a simulator that lets you generate N samples $x_i \sim p(x_i|\theta_i)$ (for parameters θ_i of our choice),
- observed data $x_{\text{obs}} \sim p(x_{\text{obs}}|\theta_{\text{true}})$,
- a prior $p(\theta)$.

Then, estimate the posterior $p(\theta|x_{\text{obs}})$.



Amortizing Bayes

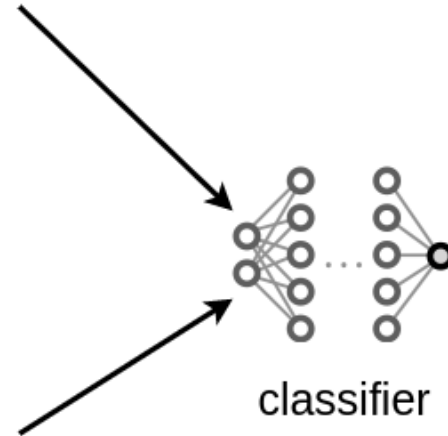
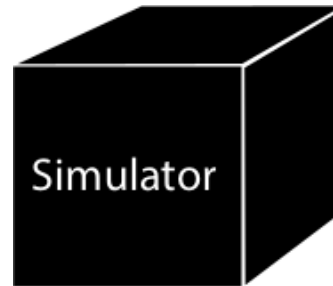
The Bayes rule can be rewritten as

$$p(\theta|x) = \frac{p(x|\theta)p(\theta)}{p(x)} = r(x|\theta)p(\theta) \approx \hat{r}(x|\theta)p(\theta),$$

where $r(x|\theta) = \frac{p(x|\theta)}{p(x)}$ is the likelihood-to-evidence ratio.

The ratio can be learned with machine learning, even when neither the likelihood nor the evidence can be evaluated!

$$x, \theta \sim p(x, \theta)$$



$$x, \theta \sim p(x)p(\theta)$$



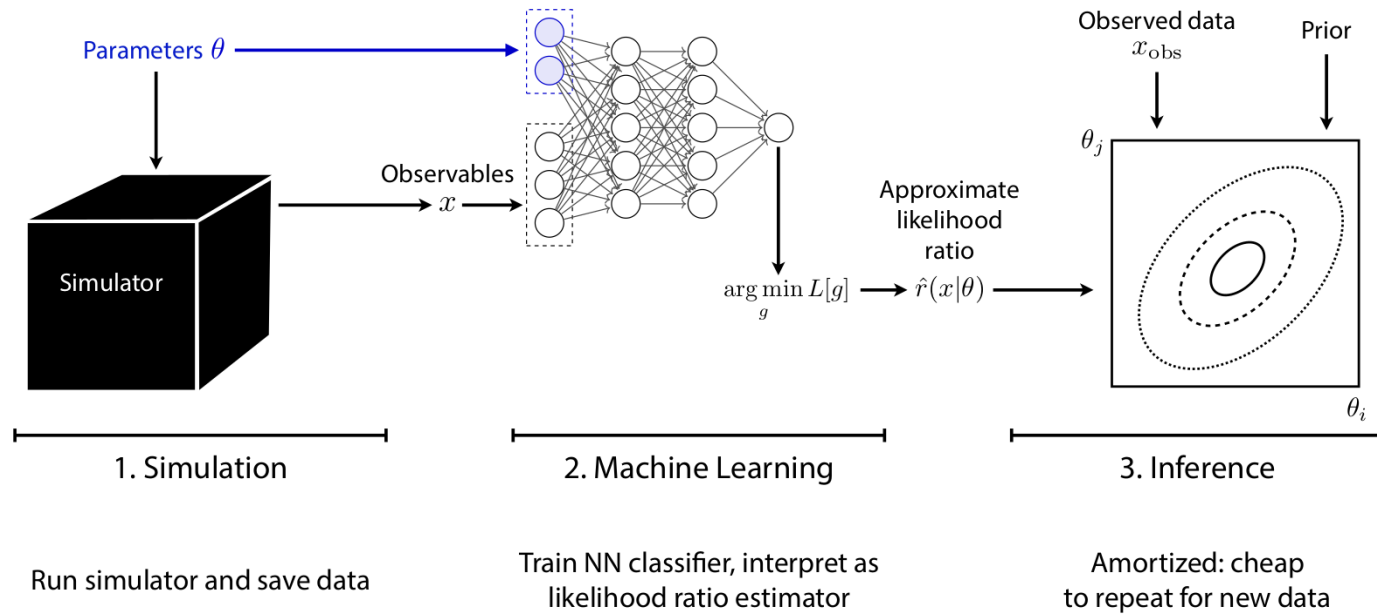
The solution d found after training approximates the optimal classifier

$$d(x, \theta) \approx d^*(x, \theta) = \frac{p(x, \theta)}{p(x, \theta) + p(x)p(\theta)}.$$

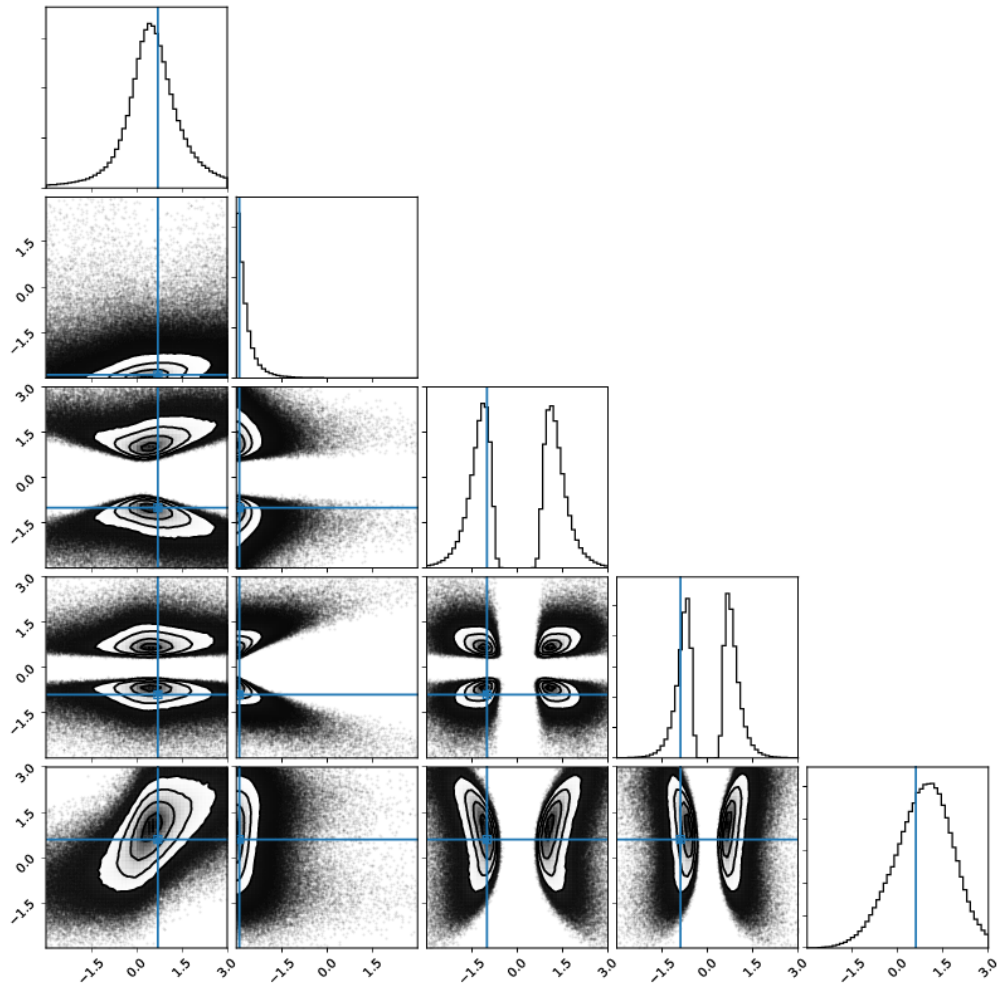
Therefore,

$$r(x|\theta) = \frac{p(x|\theta)}{p(x)} = \frac{p(x, \theta)}{p(x)p(\theta)} \approx \frac{d(x, \theta)}{1 - d(x, \theta)} = \hat{r}(x|\theta).$$

Inference



Diagnostics



How to assess that the approximate posterior is not wrong?

Coverage

- For every $x, \theta \sim p(x, \theta)$ in a validation set, compute the $1 - \alpha$ credible interval based on $\hat{p}(\theta|x) = \hat{r}(x|\theta)p(\theta)$.
- The fraction of samples for which θ is contained within the interval corresponds to the empirical coverage probability.
- If the empirical coverage is larger than the nominal coverage probability $1 - \alpha$, then the ratio estimator \hat{r} passes the diagnostic.

Convergence towards the nominal value θ^*

If the approximation \hat{r} is correct, then the posterior

$$\begin{aligned}\hat{p}(\theta|\mathcal{X}) &= \frac{p(\theta)p(\mathcal{X}|\theta)}{p(\mathcal{X})} = p(\theta) \left[\int p(\theta') \prod_{x \in \mathcal{X}} \frac{p(x|\theta')}{p(x|\theta)} d\theta' \right]^{-1} \\ &\approx p(\theta) \left[\int p(\theta') \prod_{x \in \mathcal{X}} \frac{\hat{r}(x|\theta')}{\hat{r}(x|\theta)} d\theta' \right]^{-1}\end{aligned}$$

should concentrate around θ^* as the number of observations

$$\mathcal{X} = \{x_1, \dots, x_n\},$$

for $x_i \sim p(x|\theta^*)$, increases.

ROC AUC score

The ratio estimator $\hat{r}(x|\theta)$ is only exact when samples x from the reweighted marginal model $p(x)\hat{r}(x|\theta)$ cannot be distinguished from samples x from a specific likelihood $p(x|\theta)$.

Therefore, the predictive ROC AUC performance of a classifier should be close to **0.5** if the ratio is correct.

Constraining dark matter with stellar streams

Palomar 5 (Pal5) stream
Pal5 was discovered in 2001 as the first thin stream formed from a globular cluster. Its current orbit takes it far over the galactic center.

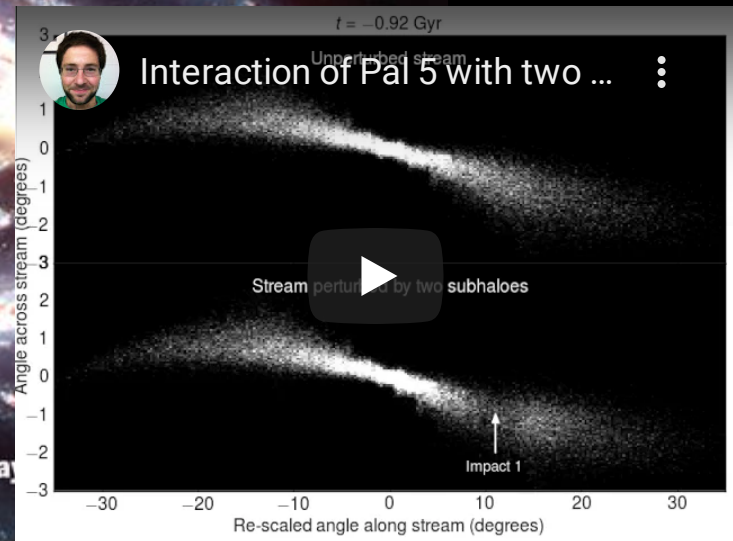
Globular clusters
These hives typically hold 100,000 stars or fewer and give rise to long, thin streams.

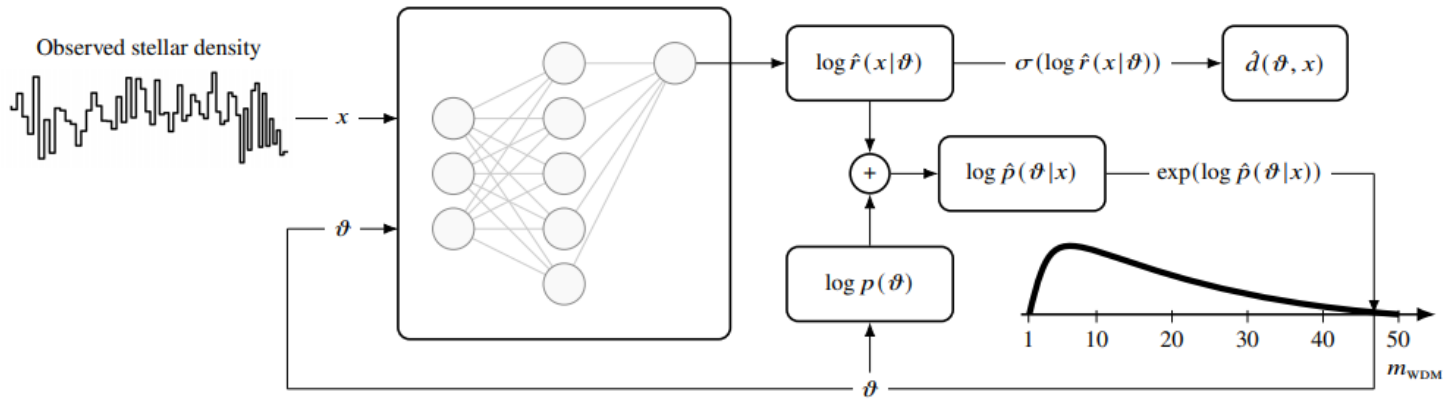
Gap

Sun

Milky Way

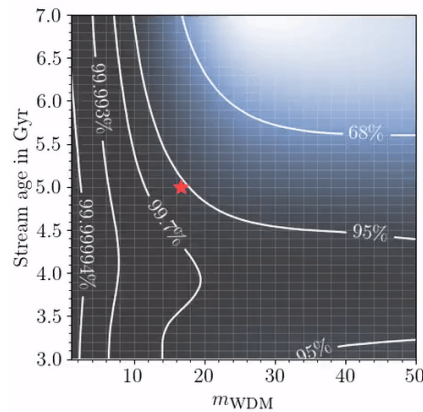
GD1 stream
Discovered in 2006, GD1 is the longest known thin stream, stretching across more than half the northern sky. It contains a gap that could be the scar of a dark matter collision 500 million years ago.



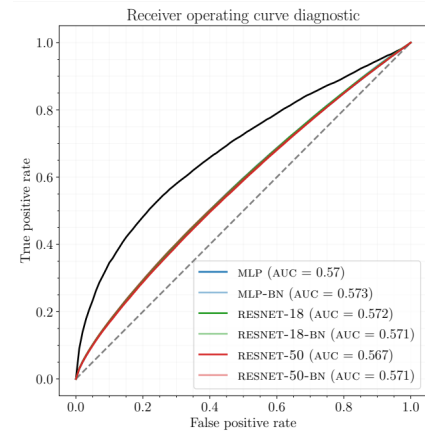


Architecture	68% CR	95% CR
$\hat{f}(x \vartheta)$ with $\vartheta \pm (m_{\text{WDM}})$		
MLP	0.685 \pm 0.004	0.954 \pm 0.002
MLP-BN	0.687 \pm 0.006	0.951 \pm 0.002
RESNET-18	0.667 \pm 0.004	0.943 \pm 0.002
RESNET-18-BN	0.672 \pm 0.004	0.945 \pm 0.001
RESNET-50	0.671 \pm 0.005	0.947 \pm 0.003
RESNET-50-BN	0.678 \pm 0.004	0.949 \pm 0.004
$\hat{f}(x \vartheta)$ with $\vartheta \pm (m_{\text{WDM}}, t_{\text{age}})$		
MLP	0.685 \pm 0.005	0.953 \pm 0.002
MLP-BN	0.685 \pm 0.004	0.952 \pm 0.003
RESNET-18	0.666 \pm 0.005	0.945 \pm 0.002
RESNET-18-BN	0.671 \pm 0.003	0.945 \pm 0.003
RESNET-50	0.674 \pm 0.006	0.944 \pm 0.002
RESNET-50-BN	0.677 \pm 0.004	0.947 \pm 0.003

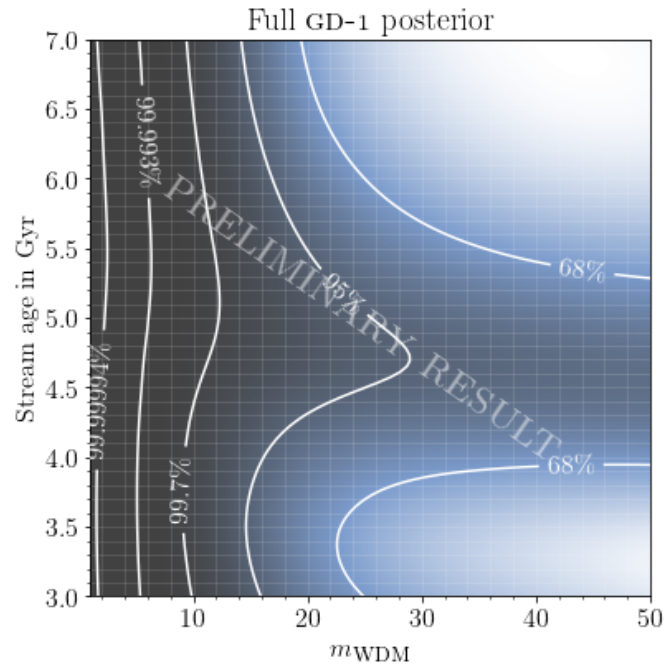
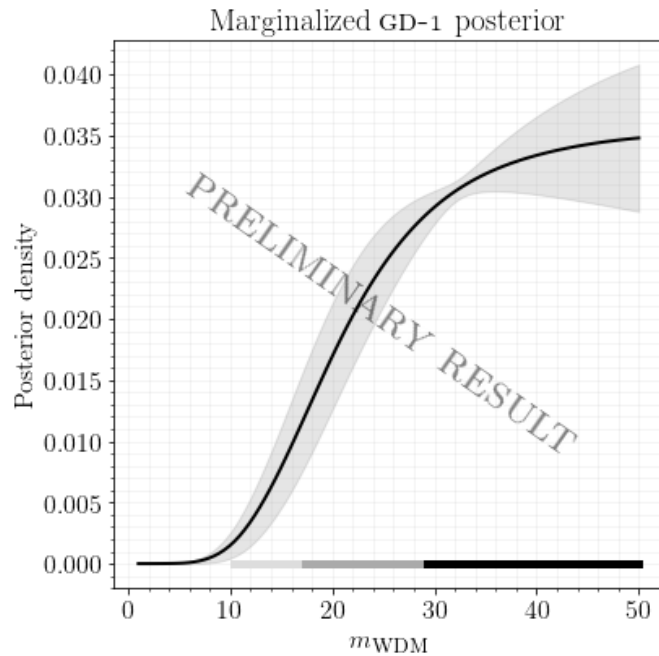
Coverage



Convergence to θ^*



ROC AUC score



Preliminary results for GD-1 suggest a **preference for CDM over WDM**.

In summary

- Much of modern science is based on simulators making precise predictions, but in which inference is challenging.
- Machine learning enables powerful inference methods.
- They work in problems from the smallest to the largest scales.
- Further advances in machine learning will translate into scientific progress.

Thanks!



References

- Hermans, J., Banik, N., Weniger, C., Bertone, G., & Louppe, G. (2020). Towards constraining warm dark matter with stellar streams through neural simulation-based inference. arXiv preprint arXiv:2011.14923.
- Cranmer, K., Brehmer, J., & Louppe, G. (2020). The frontier of simulation-based inference. *Proceedings of the National Academy of Sciences*, 117(48), 30055-30062.
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The end.