

# Neural Ratio Estimation for Simulation-Based Inference

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Computational Methods for Likelihood-Free Bayesian Inference  
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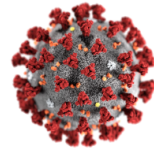
# Simulation-based inference



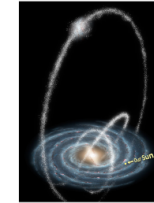
Chemical reactions



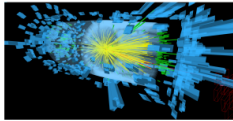
Flames



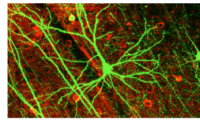
Epidemics



Stellar streams



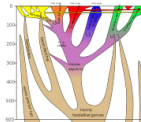
Collider experiments



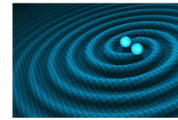
Neurons



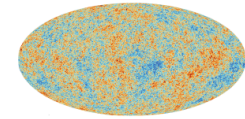
Robotics



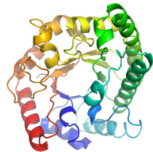
Evolution



Gravitational waves



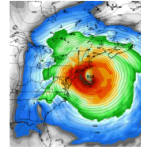
Evolution of the Universe



Protein networks



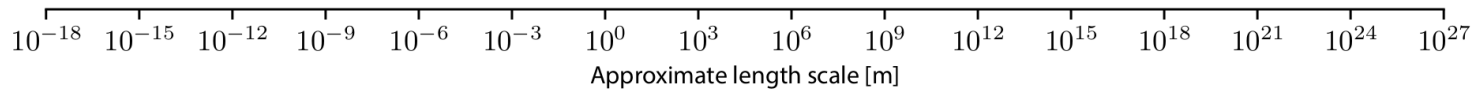
Ecological systems

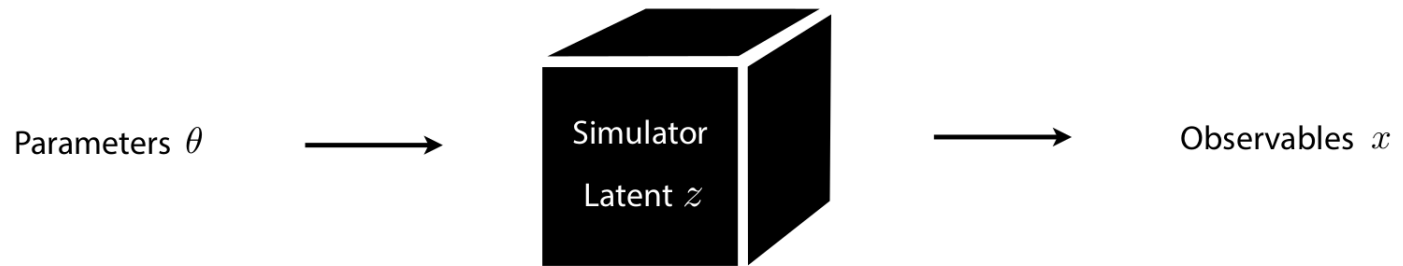


Weather and climate



Gravitational lensing





- Prediction:
- Well-motivated mechanistic, causal model
  - Simulator can generate samples  $x \sim p(x|\theta)$

- Inference:
- Interactions between low-level components lead to challenging inverse problems
  - Likelihood  $p(x|\theta) = \int dz p(x, z|\theta)$  is intractable

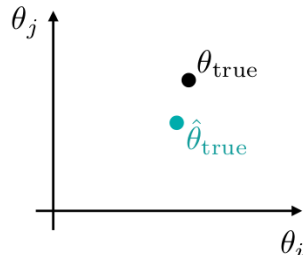
# Problem statement(s)

Start with

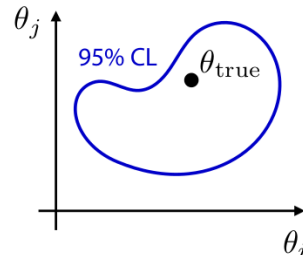
- a simulator that lets you generate  $N$  samples  $x_i \sim p(x_i | \theta_i)$  (for parameters  $\theta_i$  of our choice),
- observed data  $x_{\text{obs}} \sim p(x_{\text{obs}} | \theta_{\text{true}})$ ,
- a prior  $p(\theta)$ .

Then,

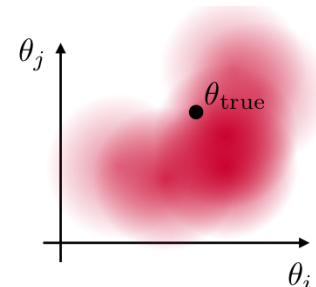
a) estimate  $\theta_{\text{true}}$   
(e.g., MLE)



b) construct  
confidence sets



c) estimate the posterior  
 $p(\theta | x_{\text{obs}})$   
(or sample from it)



# Amortizing Bayes

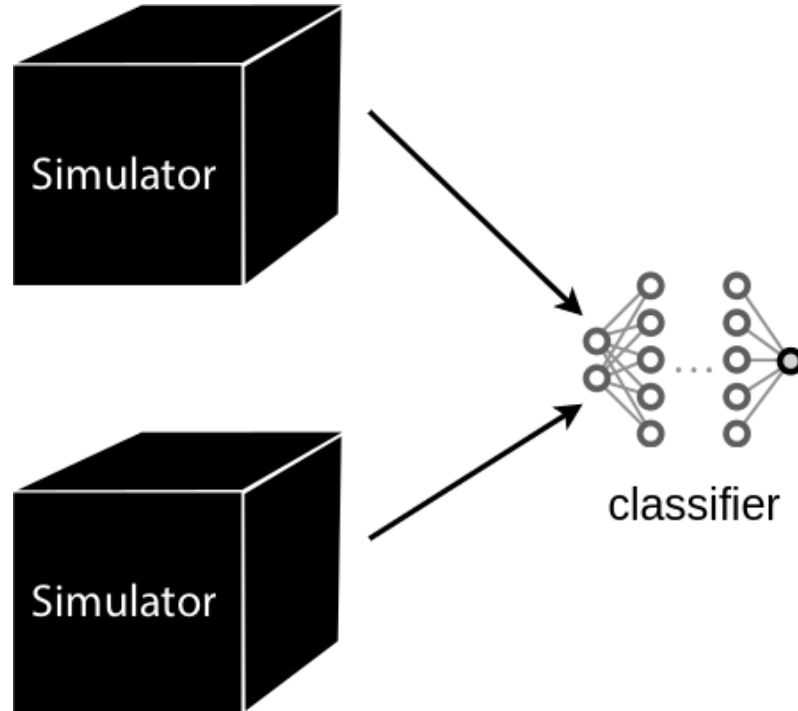
The Bayes rule can be rewritten as

$$p(\theta|x) = \frac{p(x|\theta)p(\theta)}{p(x)} = r(x|\theta)p(\theta) \approx \hat{r}(x|\theta)p(\theta),$$

where  $r(x|\theta) = \frac{p(x|\theta)}{p(x)}$  is the likelihood-to-evidence ratio.

## The likelihood ratio trick

$$x, \theta \sim p(x, \theta)$$



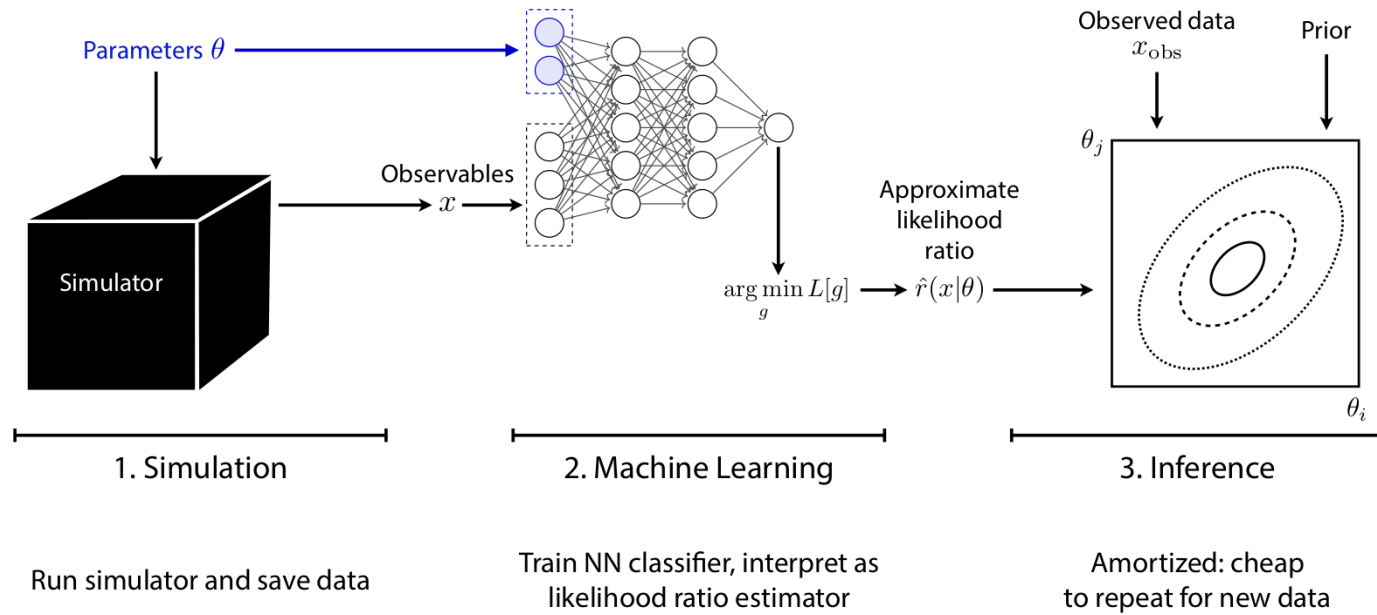
The solution  $d$  found after training approximates the optimal classifier

$$d(x, \theta) \approx d^*(x, \theta) = \frac{p(x, \theta)}{p(x, \theta) + p(x)p(\theta)}.$$

Therefore,

$$r(x|\theta) = \frac{p(x|\theta)}{p(x)} = \frac{p(x, \theta)}{p(x)p(\theta)} \approx \frac{d(x, \theta)}{1 - d(x, \theta)} = \hat{r}(x|\theta).$$

# Inference



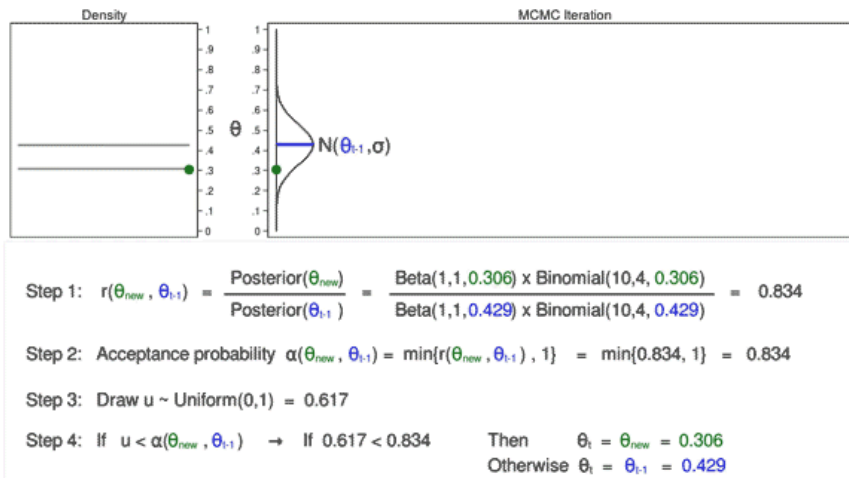


# Likelihood-free MCMC

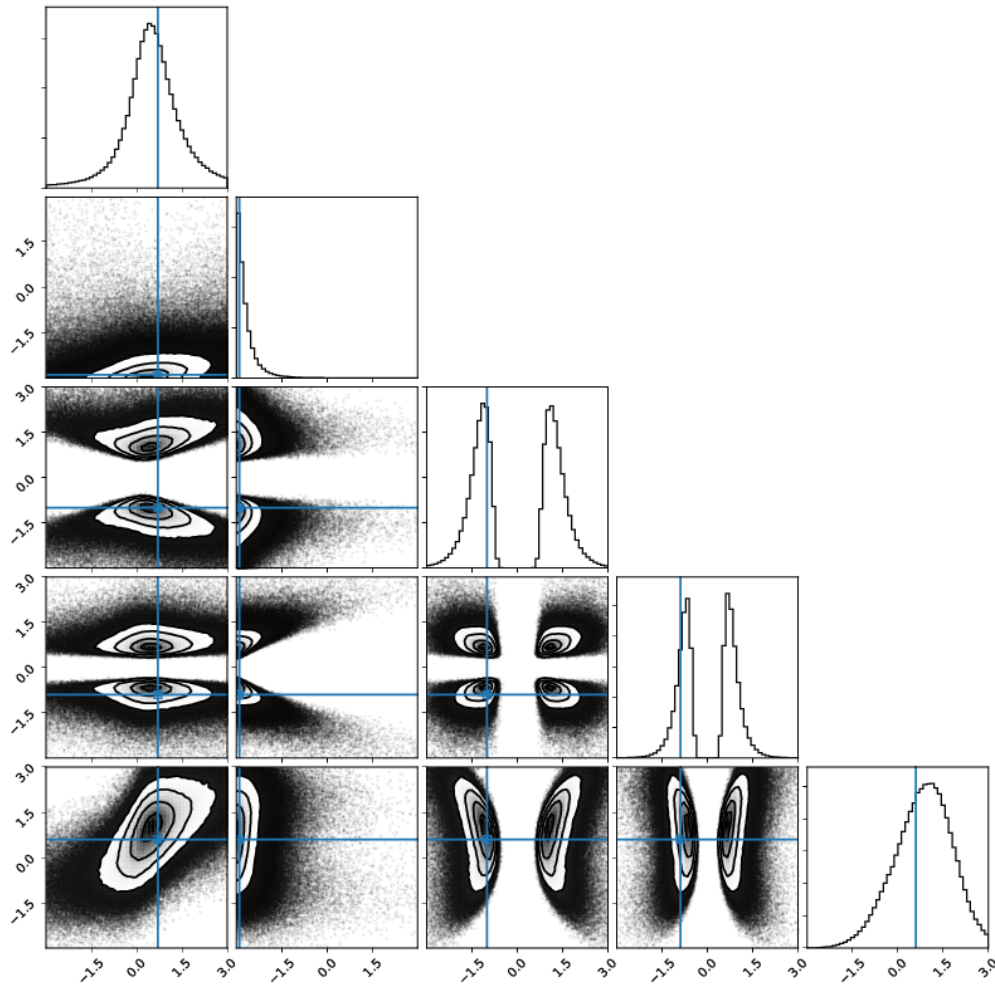
MCMC samplers require the evaluation of the posterior ratios:

$$\begin{aligned} \frac{p(\theta_{\text{new}} | x)}{p(\theta_{t-1} | x)} &= \frac{p(x | \theta_{\text{new}}) p(\theta_{\text{new}}) / p(x)}{p(x | \theta_{t-1}) p(\theta_{t-1}) / p(x)} \\ &= \frac{r(x | \theta_{\text{new}}) p(\theta_{\text{new}})}{r(x | \theta_{t-1}) p(\theta_{t-1})}. \end{aligned}$$

Extensions with HMC is possible since  $\nabla_{\theta} p(x | \theta) = \frac{\nabla_{\theta} r(x | \theta)}{r(x | \theta)}$ .



# Diagnostics



*How to assess that the approximate posterior is not wrong?*

## Coverage

- For every  $x, \theta \sim p(x, \theta)$  in a validation set, compute the  $1 - \alpha$  credible interval based on  $\hat{p}(\theta|x) = \hat{r}(x|\theta)p(\theta)$ .
- The fraction of samples for which  $\theta$  is contained within the interval corresponds to the empirical coverage probability.
- If the empirical coverage is larger than the nominal coverage probability  $1 - \alpha$ , then the ratio estimator  $\hat{r}$  passes the diagnostic.

## Convergence towards the nominal value

If the approximation  $\hat{r}$  is correct, then the posterior  $\hat{p}(\theta|\mathcal{X})$  should concentrate around  $\theta^*$  as the number of observations

$$\mathcal{X} = \{x_1, \dots, x_n\},$$

for  $x_i \sim p(x|\theta^*)$ , increases.

## ROC AUC score

The ratio estimator  $\hat{r}(x|\theta)$  is only exact when samples  $x$  from the reweighted marginal model  $p(x)\hat{r}(x|\theta)$  cannot be distinguished from samples  $x$  from a specific likelihood  $p(x|\theta)$ .

Therefore, the predictive ROC AUC performance of a classifier should be close to **0.5** if the ratio is correct.

# Constraining dark matter with stellar streams

**Palomar 5 (Pal5) stream**  
Pal5 was discovered in 2001 as the first thin stream formed from a globular cluster. Its current orbit takes it far over the galactic center.

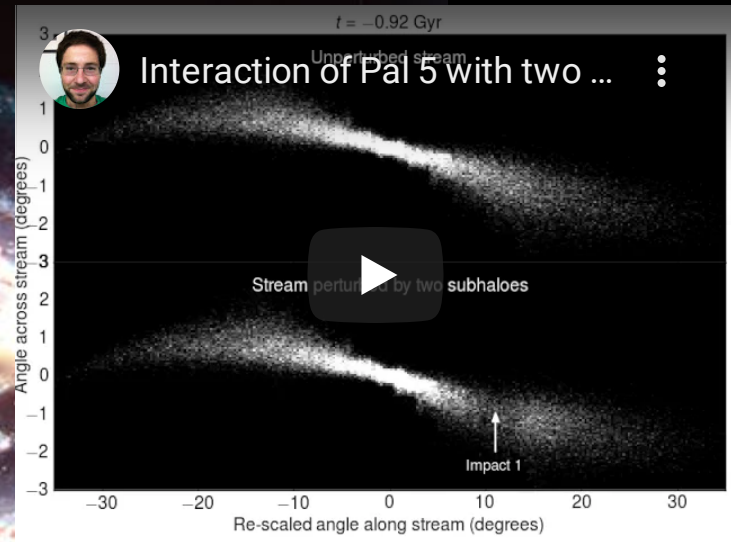
**Globular clusters**  
These hives typically hold 100,000 stars or fewer and give rise to long, thin streams.

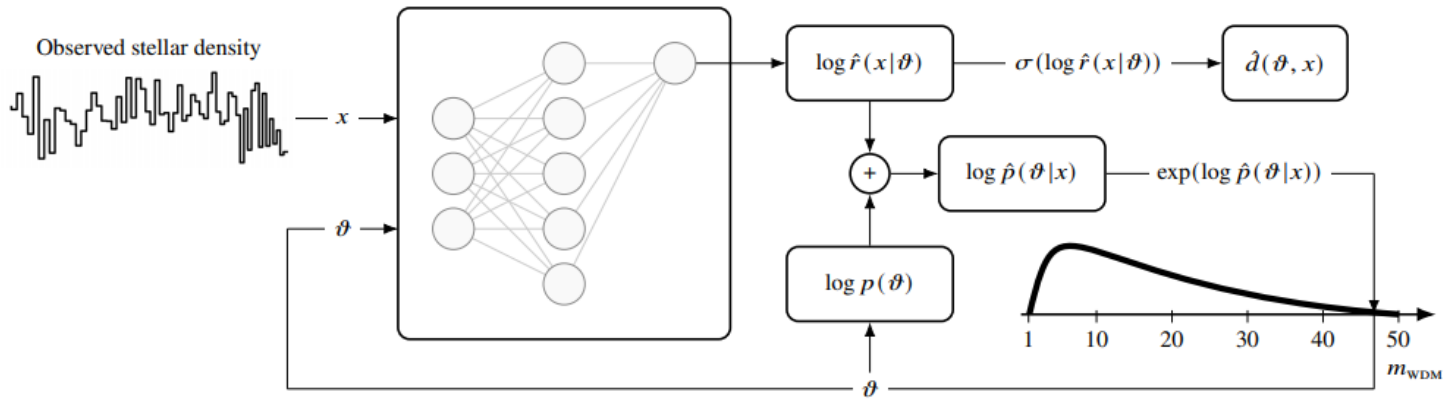
Gap

Sun

**GD1 stream**  
Discovered in 2006, GD1 is the longest known thin stream, stretching across more than half the northern sky. It contains a gap that could be the scar of a dark matter collision 500 million years ago.

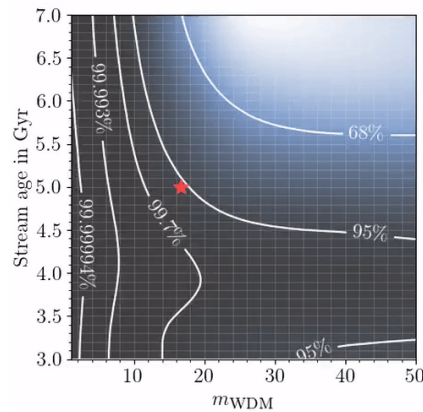
Milky Way



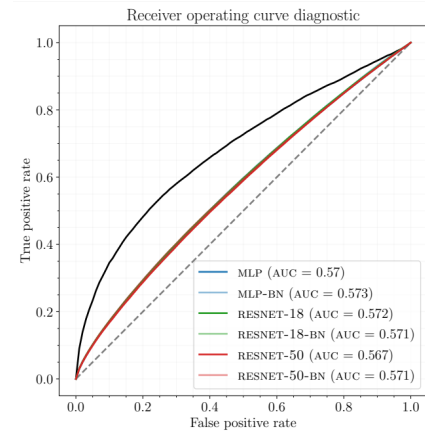


Architecture	68% CR	95% CR
$\hat{f}(x \vartheta)$ with $\vartheta \pm (m_{\text{WDM}})$		
MLP	0.685 $\pm$ 0.004	0.954 $\pm$ 0.002
MLP-BN	0.687 $\pm$ 0.006	0.951 $\pm$ 0.002
RESNET-18	0.667 $\pm$ 0.004	0.943 $\pm$ 0.002
RESNET-18-BN	0.672 $\pm$ 0.004	0.945 $\pm$ 0.001
RESNET-50	0.671 $\pm$ 0.005	0.947 $\pm$ 0.003
RESNET-50-BN	0.678 $\pm$ 0.004	0.949 $\pm$ 0.004
$\hat{f}(x \vartheta)$ with $\vartheta \pm (m_{\text{WDM}}, t_{\text{age}})$		
MLP	0.685 $\pm$ 0.005	0.953 $\pm$ 0.002
MLP-BN	0.685 $\pm$ 0.004	0.952 $\pm$ 0.003
RESNET-18	0.666 $\pm$ 0.005	0.945 $\pm$ 0.002
RESNET-18-BN	0.671 $\pm$ 0.003	0.945 $\pm$ 0.003
RESNET-50	0.674 $\pm$ 0.006	0.944 $\pm$ 0.002
RESNET-50-BN	0.677 $\pm$ 0.004	0.947 $\pm$ 0.003

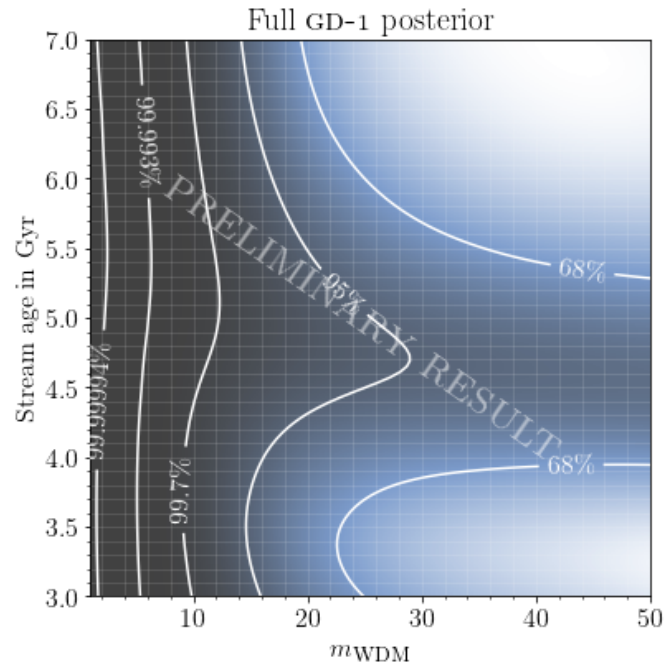
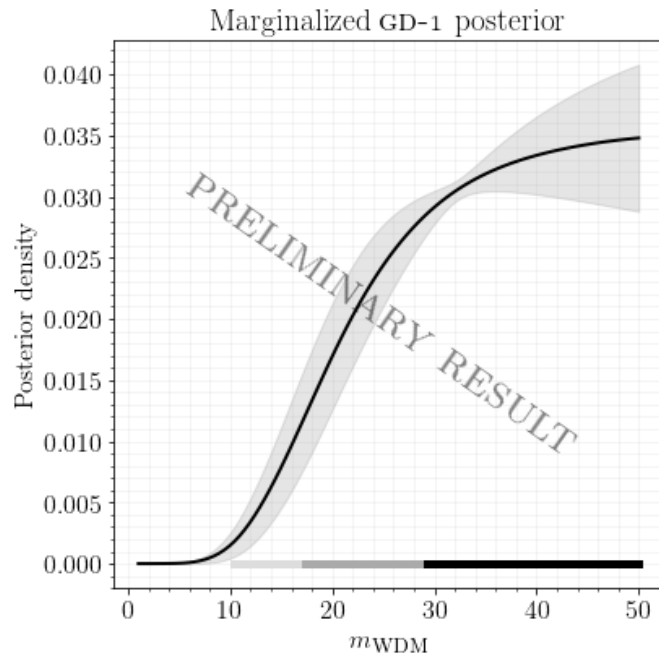
Coverage



Convergence to  $\theta^*$



ROC AUC score



Preliminary results for GD-1 suggest a **preference for CDM over WDM**.



# In summary

- Much of modern science is based on simulators making precise predictions, but in which inference is challenging.
- Machine learning enables powerful inference methods, such as ratio estimation based on neural networks.
- Amortized estimators are well suited for diagnosing the quality of the resulting posteriors.

# Thanks!



# References

- Hermans, J., Banik, N., Weniger, C., Bertone, G., & Louppe, G. (2020). Towards constraining warm dark matter with stellar streams through neural simulation-based inference. arXiv preprint arXiv:2011.14923.
- Hermans, J., Begy, V., & Louppe, G. (2019). Likelihood-free MCMC with Approximate Likelihood Ratios. arXiv preprint arXiv:1903.04057.
- Cranmer, K., Pavez, J., & Louppe, G. (2015). Approximating likelihood ratios with calibrated discriminative classifiers. arXiv preprint arXiv:1506.02169.

The end.