

Larger and faster: revised properties and a shorter orbital period for the WASP-57 planetary system from a pro-am collaboration*

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ABSTRACT

Transits in the WASP-57 planetary system have been found to occur half an hour earlier than expected. We present ten transit light curves from amateur telescopes, on which this discovery was based, thirteen transit light curves from professional facilities which confirm and refine this finding, and high-resolution imaging which show no evidence for nearby companions. We use these data to determine a new and precise orbital ephemeris, and measure the physical properties of the system. Our revised orbital period is 4.5 s shorter than found from the discovery data alone, which explains the early occurrence of the transits. We also find both the star and planet to be larger and less massive than previously thought. The measured mass and radius of the planet are now consistent with theoretical models of gas giants containing no heavy-element core, as expected for the sub-solar metallicity of the host star. Two transits were observed simultaneously in four passbands. We use the resulting light curves to measure the planet’s radius as a function of wavelength, finding that our data are sufficient in principle but not in practise to constrain its atmospheric properties. We conclude with a discussion of the current and future status of transmission photometry studies for probing the atmospheres of gas-giant transiting planets.

1 INTRODUCTION

Although the first transiting extrasolar planet (TEP) was only discovered in late 1999 (Henry et al. 2000; Charbonneau et al. 2000), and the second as recently as 2003 (Konacki et al. 2003), the number currently known has already exceeded 1200¹. The great majority of those are small objects observed using the NASA *Kepler* satellite: validation of the planetary nature of these bodies has been greatly helped by their occurrence in systems of multiple planets (see Rowe et al. 2014) but detailed studies are difficult due to their small size and long orbital periods (P_{orb} s).

A significant number (231 as of 2015/07/21) of the known TEPs are hot Jupiters, adopting a definition of mass $M_b > 0.3 M_{\text{Jup}}$ and $P_{\text{orb}} < 10$ d. These are much better suited to characterisation with existing facilities, as their relatively large masses and radii, short orbital periods, and bright host stars make photometric and spectroscopic observations easier and more productive. Perhaps the single most important observable property of a planet is its orbital period: the period distributions of exoplanets provide an insight into the mechanisms governing their formation and evolution (e.g. Mordasini et al. 2009a,b; Benítez-Llambay et al. 2011; Hellier et al. 2012), and a precise value is mandatory for performing follow-up observations.

In this work we study the WASP-57 system, whose planetary nature was discovered by the SuperWASP consortium (Faedi et al. 2013, hereafter F13). WASP-57 contains a star slightly cooler and less massive than the Sun ($T_{\text{eff}} = 5600 \pm 100$ K, $M_A = 0.89 \pm 0.07 M_{\odot}$) orbited by a planet which is the same size but less massive than Jupiter ($M_b = 0.64 \pm 0.06 M_{\text{Jup}}$, $R_b = 1.05 \pm 0.05 R_{\text{Jup}}$). The moderately different properties found by F13 placed the planet at the lower edge of the distribution of gas giant TEPs in the mass–radius diagram, making it a good candidate for hosting a heavy-element core despite the subsolar metallicity of the host star ($[\text{Fe}/\text{H}] = -0.25 \pm 0.10$). The analysis by F13 was based on SuperWASP photometry (Pollacco et al. 2006), radial velocities from CORALIE spectra, plus two complete and one partial transit light curves from the Euler and TRAPPIST telescopes at ESO La Silla, Chile. No further work on this system has been published.

Early in the 2014 observing season a group of amateur astronomers noticed that the transits of WASP-57 were occurring half an hour earlier than expected, which is a significant fraction of the 2.3 hr total transit duration. This was immediately confirmed using a transit of WASP-57 which had been serendipitously observed on 2014/05/18 using the Danish 1.5 m telescope, at La Silla. A second transit observation was scheduled on 2014/06/24 with the Danish telescope, the ESO 2.2 m telescope and GROND imager, and the immediately following transit with the amateur observers in Europe, allowing its early arrival to be reconfirmed. We also possess high-precision light curves of WASP-57 obtained in the 2012 season, before the imprecision of the original orbital ephemeris became apparent. In this work we present all the data we have obtained for WASP-57, produce a revised ephemeris which can be

used for follow-up observations in future, measure the physical properties of the system to high precision, and search for variations of the measured planetary radius with wavelength

2 OBSERVATIONS AND DATA REDUCTION

A total of ten transit light curves were obtained by LB, GC, CL, AM and GM using telescopes of apertures between 180 mm and 300 mm, sited in Italy. Further details of the observational setup and numbers of datapoints are given in Table 1.

Two complete transits of WASP-57 were observed using the 1.54 m Danish Telescope and DFOSC instrument at ESO La Silla, Chile (see Dominik et al. 2010), on the dates 2014/05/18 and 2014/06/24. DFOSC has a plate scale of $0.39'' \text{ pixel}^{-1}$ and a 2048^2 pixel CCD, giving a field of view of $13.7' \times 13.7'$. We windowed the CCD down to 1100×900 and 1045×920 pixels to shorten the dead time between exposures, resulting in images containing WASP-57 and six decent comparison stars. Both transits were obtained through a Bessell R filter. The instrument was defocussed in order to improve the efficiency of the observations, and to combat time-correlated noise (see Southworth et al. 2009). The telescope was autoguided to limit pointing drifts to less than five pixels over each observing sequence. An observing log is given in Table 2 and the light curves are plotted individually in Fig. 1.

The transit on 2014/06/24 observed with DFOSC was also monitored using GROND (Greiner et al. 2008) mounted on the MPG 2.2 m telescope at La Silla, Chile. GROND was used to obtain light curves simultaneously in four passbands, which approximate the SDSS g , r , i and z bands. The small field of view of this instrument ($5.4' \times 5.4'$ at a plate scale of $0.158'' \text{ pixel}^{-1}$) meant that fewer comparison stars were available. The instrument was defocussed and the telescope was autoguided throughout the observing sequence. Poor weather conditions (high wind) forced closure of the telescope before the transit finished, so the light curves have only partial coverage of the transit (see Table 2 and Fig. 1).

We now turn to observations obtained prior to the conception of the current work. We observed WASP-57 on the night of 2012/05/10 using the BUSCA instrument on the 2.2 m telescope at Calar Alto Astronomical Observatory. BUSCA is capable of observing simultaneously in four passbands, for which we chose Gunn u , g , r and z . The motivation for these choices, and a detailed discussion on the use of BUSCA for planetary transit observations, can be found in Southworth et al. (2012). All four CCDs on BUSCA have a plate scale of $0.176'' \text{ pixel}^{-1}$, but were operated with 2×2 binning. Whilst the full field of view of 12×12 arcmin was accessible in the u band, the available field in the g , r and z bands was vignetted into a circle of diameter of approximately $6'$. The instrument was defocussed and the telescope was autoguided throughout the observations (Table 2). The u -band data are of insufficient quality for full modelling but can be used to obtain a time of minimum and to check for a possible variation of measured planetary radius with wavelength.

One transit of WASP-57 was observed on 2012/04/01 with EulerCam, using the same methods as for the EulerCam transit in F13. EulerCam is a CCD imager mounted on the 1.2 m Euler-Swiss telescope, La Silla, with a field of view of $14.7' \times 14.7'$

* Based on data collected by MiNDStEP with the Danish 1.54 m telescope, and data collected with GROND on the MPG 2.2 m telescope, both located at ESO La Silla.

¹ See TEPcat (Transiting Extrasolar Planet Catalogue; Southworth 2011) at: <http://www.astro.keele.ac.uk/jkt/tepcat/>

Table 1. Instrumental setup for the amateur observations. N_{obs} is the number of observations.

Date	Observer	Telescope	CCD	Filter	N_{obs}
2014/05/24	C. Lopresti	180 mm Maksutov-Newton	SBIG ST10XME	red	39
2014/06/10	G. Corfini	200 mm aperture, 800 mm focal length	SBIG STT-1603	clear	87
2014/06/10	C. Lopresti	300 mm aperture, 1500 mm focal length	SBIG ST10XME	red	42
2014/06/10	A. Marchini	300 mm Zen Maksutov-Cassegrain	STL-6303	Cousins R	58
2014/06/27	L. Barbieri	300 mm aperture, 3000 mm focal length	SBIG ST9	clear	40
2014/06/27	G. Corfini	200 mm aperture, 800 mm focal length	SBIG STT-1603	clear	34
2014/06/27	C. Lopresti	180 mm Maksutov-Newton	SBIG ST10XME	red	16
2014/06/27	C. Lopresti	300 mm aperture, 1500 mm focal length	SBIG ST10XME	red	58
2014/06/27	G. Marino	250 mm aperture, 1200 mm focal length	SBIG ST7-XME	clear	64
2015/05/28	C. Lopresti	180 mm Maksutov-Newton	SBIG ST10XME	red	30

Table 2. Log of the observations obtained from professional telescopes. N_{obs} is the number of observations, T_{exp} is the exposure time, T_{dead} is the dead time between exposures, ‘Moon illum.’ is the fractional illumination of the Moon at the midpoint of the transit, and N_{poly} is the order of the polynomial fitted to the out-of-transit data. The aperture radii are target aperture, inner sky and outer sky, respectively. The ‘bb’ filter is a blue-blocking filter.

Instrument	Date of first obs	Start time (UT)	End time (UT)	N_{obs}	T_{exp} (s)	T_{dead} (s)	Filter	Airmass	Moon illum.	Aperture radii (pixels)	N_{poly}	Scatter (mmag)
TRAPPIST	2012/03/15	04:34	08:20	367	20	10	bb	1.83 → 1.12 → 1.13	0.463	11.3 22.7 36.3	1	2.25
TRAPPIST	2012/04/01	04:40	09:36	701	15	7	bb	1.36 → 1.12 → 1.45	0.649	13.5 19.3 30.9	1	3.49
Euler	2012/04/01	05:20	09:35	212	50–180	16	r	1.23 → 1.12 → 1.44	0.649	24	0	0.98
BUSCA	2012/05/10	23:16	02:49	82	120	35	u	1.31 → 1.29 → 1.86	0.665	15 25 45	1	3.46
BUSCA	2012/05/10	23:16	03:37	100	120	35	g	1.31 → 1.29 → 2.46	0.665	17 27 50	1	1.26
BUSCA	2012/05/10	23:16	03:34	99	120	35	r	1.31 → 1.29 → 2.44	0.665	18 28 50	1	0.82
BUSCA	2012/05/10	23:16	03:37	98	120	35	z	1.31 → 1.29 → 2.46	0.665	18 28 50	1	1.58
DFOSC	2014/05/18	01:58	06:51	162	100	9	R	1.32 → 1.12 → 1.70	0.761	12 20 40	1	0.74
DFOSC	2014/06/24	23:33	04:15	155	100	8	R	1.27 → 1.12 → 1.53	0.050	16 26 45	1	0.71
GROND	2014/06/24	23:38	02:34	53	90–110	54	g	1.25 → 1.12 → 1.17	0.050	40 60 90	1	0.89
GROND	2014/06/24	23:38	02:34	53	90–110	54	r	1.25 → 1.12 → 1.17	0.050	40 60 90	1	0.58
GROND	2014/06/24	23:38	02:34	53	90–110	54	i	1.25 → 1.12 → 1.17	0.050	32 55 75	1	0.58
GROND	2014/06/24	23:38	02:34	53	90–110	54	z	1.25 → 1.12 → 1.17	0.050	30 50 80	1	1.02

at $0.23'' \text{ pixel}^{-1}$. We obtained 212 images through a Gunn r filter, without applying a defocus to the instrument. Further details on EulerCam and the data reduction procedure can be found in Lendl et al. (2012).

Two transits of WASP-57 were observed on 2012/03/15 and 2012/04/01 using the 0.6 m TRAPPIST robotic telescope located at La Silla (Gillon et al. 2011; Jehin et al. 2011). The $2\text{k} \times 2\text{k}$ CCD was thermoelectrically cooled and yielded a field of view of $22' \times 22'$ at $0.65'' \text{ pixel}^{-1}$. Images were obtained with a slight defocus and through a blue-blocking filter² that has a transmittance greater than 90% from 500 nm to beyond 1000 nm.

2.1 Data reduction

The data from the amateur telescopes were all reduced using MaxIm DL³. In each case the science images were calibrated using dark and flat-field frames.

The data from DFOSC, GROND and BUSCA were reduced using aperture photometry as implemented in the DEFOT code (Southworth et al. 2009, 2014), which relies on the

IDL⁴/ASTROLIB⁵ implementation of DAOPHOT (Stetson 1987). Master bias and flat-field images were constructed but generally found to have an insignificant effect on the quality of the photometry. Image motion was tracked by cross-correlating individual images with a reference image.

We obtained photometry on the instrumental system using software apertures of a range of sizes, and retained those which gave light curves with the smallest scatter (Table 2). We found that the choice of aperture size affects the scatter but not the shape of the transit in the final light curve. The instrumental magnitudes were then transformed to differential-magnitude light curves, normalised to zero magnitude outside transit using first-order polynomials (Table 2) fitted to the out-of-transit data. The differential magnitudes are relative to a weighted ensemble of typically five (DFOSC) or two to four (GROND) comparison stars. The comparison star weights and polynomial coefficients were simultaneously optimised to minimise the scatter in the out-of-transit data.

Finally, the timestamps for the datapoints were converted to the BJD(TDB) timescale (Eastman et al. 2010). We performed manual time checks for several images obtained with DFOSC

² <http://www.astrodon.com/products/filters/exoplanet/>

³ <http://www.cyanogen.com/maximmain.php>

⁴ The acronym IDL stands for Interactive Data Language and is a trademark of ITT Visual Information Solutions. For further details see: <http://www.exelisvis.com/ProductsServices/IDL.aspx>

⁵ The ASTROLIB subroutine library is distributed by NASA. For further details see: <http://idlastro.gsfc.nasa.gov/>.

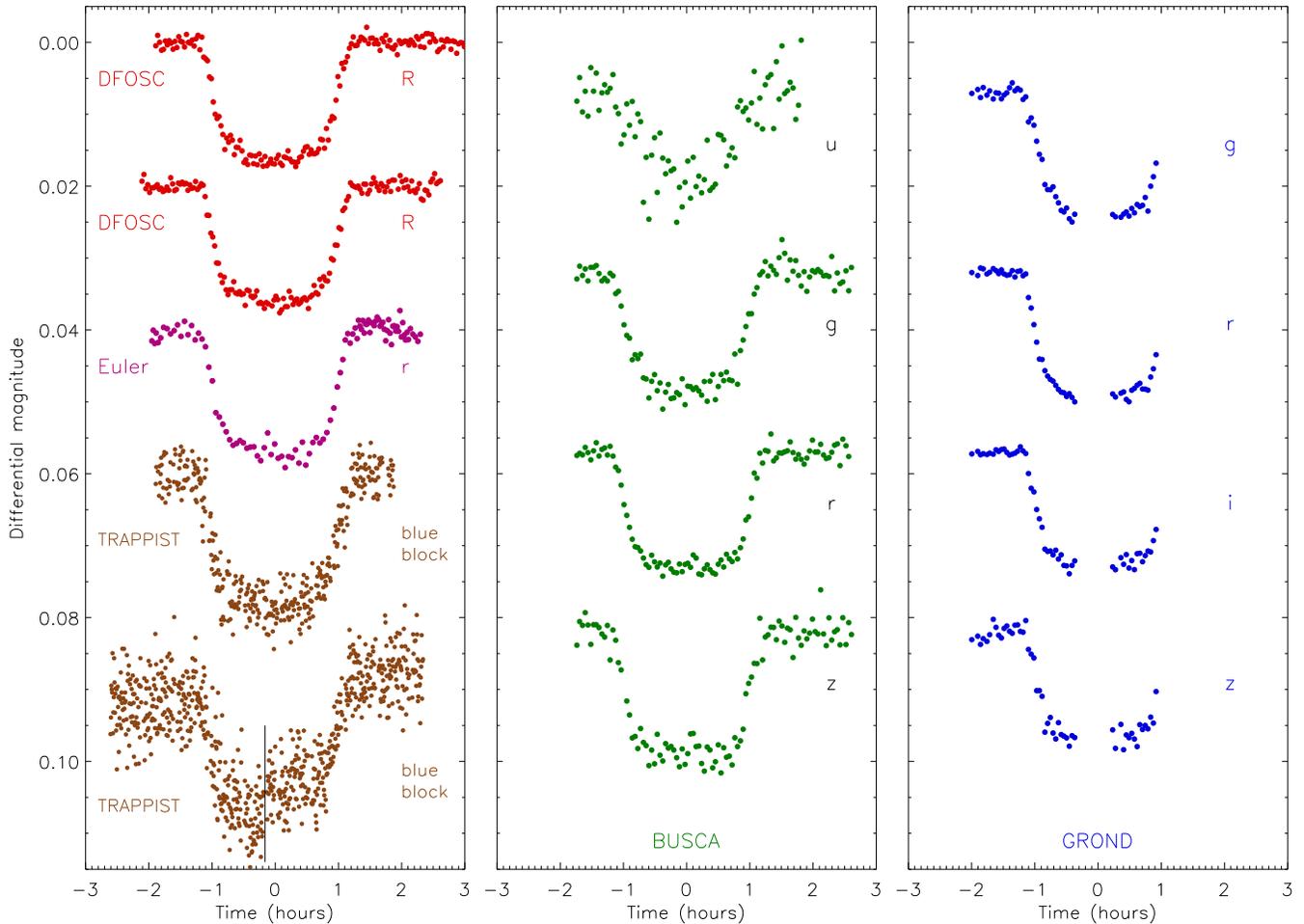


Figure 1. All light curves from professional facilities presented in this work, grouped and colour-coded according to the telescope used. The instrument and filter are labelled individually for each light curve. The second light curve from TRAPPIST has a discontinuity shortly before the midpoint of the transit due to a meridian-flip. This is indicated using a vertical black line.

and verified that the FITS file timestamps are on the UTC system to within a few seconds. In recent work on the WASP-103 system we found that the timestamps from DFOSC and GROND agree to within a few seconds, supporting the reliability of both (Southworth et al. 2015). The reduced data are given in Table 3 and will be lodged with the CDS⁶.

The data from EulerCam were reduced using aperture photometry following the methods given by Lendl et al. (2012). Differential aperture photometry was also used on the TRAPPIST data, using carefully selected extraction apertures and reference stars. For more details on the TRAPPIST data reduction procedures, see e.g. Gillon et al. (2013). The transit on 2012/04/01 was obtained in two sequences separated by a meridian flip, and the two sets of data were reduced independently.

2.2 High-resolution imaging

We obtained several images of WASP-57 with DFOSC in sharp focus, allowing us to check for the presence of faint nearby stars whose light might act to decrease the observed transit depth (Daemgen et al. 2009). The closest stars we found on any image

are much fainter than WASP-57, and are over $45''$ distant, so are too far away to affect our photometry.

We also obtained a high-resolution image of WASP-57 using the Lucky Imager (LI) mounted on the Danish telescope (see Skottfelt et al. 2013, 2015). The LI uses an Andor 512×512 pixel electron-multiplying CCD, with a pixel scale of $0.09'' \text{ pixel}^{-1}$ giving a field of view of $45'' \times 45''$. The data were reduced using a dedicated pipeline and the 2% of images with the smallest point spread function (PSF) were shifted and added to yield a combined image whose PSF is smaller than the seeing limit. A long-pass dichroic was used, resulting in a response function which approximates that of SDSS $i+z$. An overall exposure time of 600 s corresponds to an effective exposure time of 12 s for the best 2% of the images (Fig. 2). The FWHM of the PSF is 4.0×4.2 pixels ($0.36'' \times 0.38''$).

Two faint stars were detected on the LI image, at angular distances of $10.99 \pm 0.05''$ and $21.56 \pm 0.07''$ from WASP-57, and fainter by 8.7 ± 0.7 mag and 8.2 ± 0.4 mag. Neither of these stars is sufficiently bright and close to WASP-57 to affect the analysis presented in the current work. We assessed the limiting magnitude of the LI image by placing a box with sides equal to the FWHM of the star on each pixel on the image. The standard deviation of the counts within each box was calculated, and a 3σ detection threshold was generated. The mean detection threshold at a given radius from

⁶ <http://vizier.u-strasbg.fr/>

Table 3. Sample of the data presented in this work (the first and last datapoints of each light curve). The full dataset will be made available at the CDS.

Instrument	Filter	BJD(TDB)	Diff. mag.	Uncertainty
TRAPPIST	bb	2456001.690740	0.0014100	0.0033979
TRAPPIST	bb	2456001.847510	0.0018100	0.0019776
TRAPPIST	bb	2456018.694550	0.0062000	0.0041227
TRAPPIST	bb	2456018.900580	-0.0042000	0.0041094
Euler	r	2456018.722460	0.0015300	0.0014096
Euler	r	2456018.899430	0.0006000	0.0010572
BUSCA	u	2456058.476778	0.0011721	0.0027719
BUSCA	u	2456058.624676	-0.0073014	0.0050536
BUSCA	g	2456058.476778	0.0008973	0.0010736
BUSCA	g	2456058.657776	-0.0006933	0.0015483
BUSCA	r	2456058.476778	0.0004552	0.0007219
BUSCA	r	2456058.655976	0.0005960	0.0009239
BUSCA	z	2456058.476778	0.0018345	0.0015082
BUSCA	z	2456058.657776	0.0003519	0.0017916
DFOSC	R	2456796.588810	0.0004779	0.0006752
DFOSC	R	2456796.792254	0.0006067	0.0008685
DFOSC	R	2456833.486396	-0.0006900	0.0007120
DFOSC	R	2456833.683083	-0.0007523	0.0007212
GROND	g	2456833.490392	0.0000938	0.0008559
GROND	g	2456833.611846	0.0097990	0.0009372
GROND	r	2456833.490392	0.0000128	0.0005593
GROND	r	2456833.611846	0.0114375	0.0006182
GROND	i	2456833.490392	0.0002314	0.0005610
GROND	i	2456833.611846	0.0107411	0.0009272
GROND	z	2456833.490392	0.0010646	0.0010125
GROND	z	2456833.611846	0.0082877	0.0010501

the target star was then converted to a relative magnitude. Further details of the detection and reduction methods are given in Evans et al. (2015, in preparation). The contrast curve is shown in Fig. 3.

3 TRANSIT TIMING ANALYSIS

The issue which brought WASP-57 to our attention was the offset between the predicted and actual times of transit. We have therefore obtained as many measured times of mid-transit as possible. We first modelled the two DFOSC transits individually using the JKTEBOP code (see below), as these are the two light curves which have full coverage of a transit with a low scatter in the data. We scaled the errorbars for each light curve to yield a reduced χ^2 of $\chi^2_\nu = 1.0$ versus the fitted model. This step is necessary because the uncertainties from the APER algorithm tend to be moderately too small.

We then modelled the light curves from the amateur telescopes with JKTEBOP but fitting for only the time of mid-transit and the out-of-transit brightness of the system. The other photometric parameters were fixed to the best-fitting values from the two DFOSC light curves. The uncertainties in the transit times were multiplied by $\sqrt{\chi^2_\nu}$ to account for the underestimated observational errors in most of the datasets. We performed the same process on the GROND data, except this time we fitted the out-of-transit brightness as a linear function of time rather than just a constant offset from zero differential magnitude.

We then turned to published data. The discovery paper of WASP-57 (F13) contains two light curves observed with TRAP-

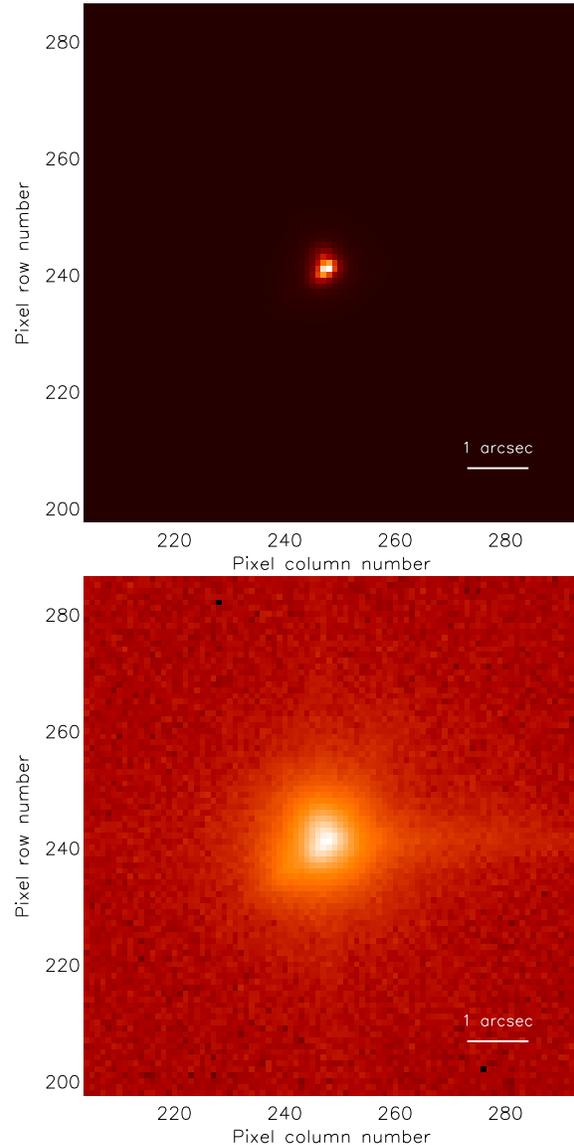


Figure 2. High-resolution Lucky Image of the field around WASP-57. The upper panel has a linear flux scale for context and the lower panel has a logarithmic flux scale to enhance the visibility of any faint stars. Each image covers $8'' \times 8''$ centred on WASP-57. A bar of length $1''$ is superimposed in the bottom-right of each image. The image is a sum of the best 2% of the original images.

PIST and one with the Euler telescope. We fitted these as above, with the photometric parameters fitted for the two light curves with complete transit coverage and fixed for the TRAPPIST light curve which only contains the second half of a transit. We also included one transit time obtained by U. Dittler and lodged on the Exoplanet Transit Database⁷ (Poddaný et al. 2010).

The SuperWASP data which triggered the discovery of the planetary nature comprise approximately 30 000 datapoints obtained during the 2008, 2009 and 2010 observing seasons. These data were obtained and separated into individual seasons, then fit-

⁷ The Exoplanet Transit Database (ETD) can be found at: <http://var2.astro.cz/ETD/credit.php>; see also TRESKA at: <http://var2.astro.cz/EN/tresca/index.php>.

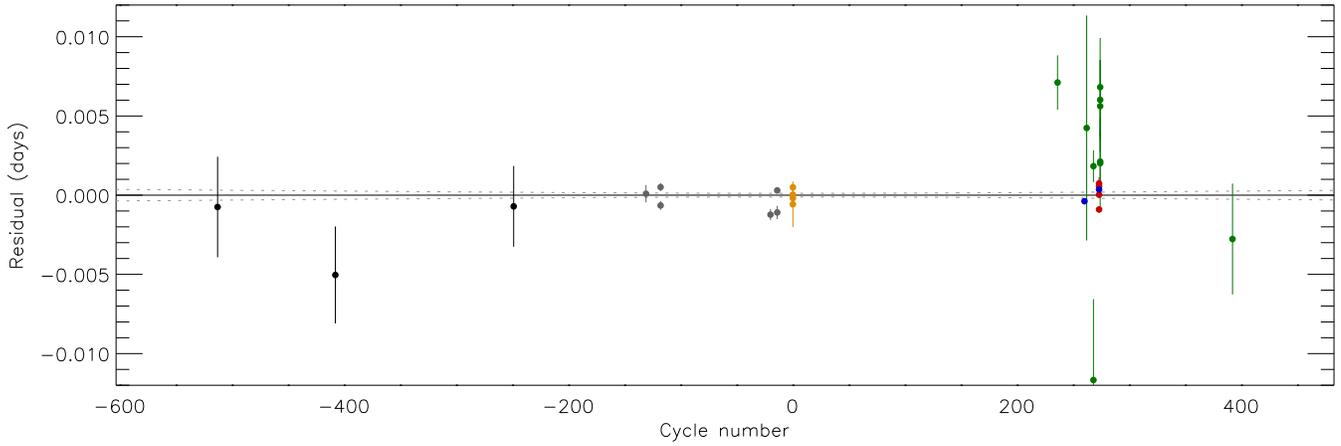


Figure 4. Plot of the residuals of the timings of mid-transit for WASP-57 versus a linear ephemeris (see Table 4). The points are colour-coded according to their source: black for the WASP data, green for the amateur timings in the current work, blue for DFOSC, red for GROND, off-yellow for BUSCA, and grey for the TRAPPIST and Euler telescopes. The dotted lines show the 1σ uncertainty in the ephemeris as a function of cycle number.

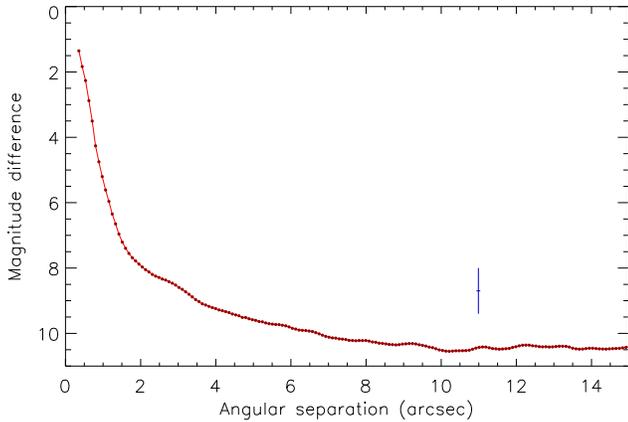


Figure 3. Contrast curve giving the limiting magnitude of the LI observation as a function of angular distance from WASP-57 (dark red circles connected by a red line). The closest detected star is shown as a blue data-point.

ted with JKTEBOP in the same way as for the data obtained using amateur telescopes. The resulting season-averaged times of minimum are consistent with the linear ephemeris found below, but are of low precision. WASP-57 A, at $V = 13.04$, is comparatively faint for the SuperWASP telescopes so suffers from a large scatter in its light curve. We included these times of minimum light in the following analysis, but note that they do not have a significant effect on the results.

All times of mid-transit were then fitted with a straight line versus cycle number to determine a new linear orbital ephemeris. Table 4 gives all transit times plus their residual versus the fitted ephemeris. We chose the reference epoch to be that for our BUSCA observations, in order to limit the covariance between the reference time of minimum and the orbital period. The resulting ephemeris is

$$T_0 = \text{BJD(TDB)} 2\,456\,058.54910(16) + 2.83891856(81) \times E$$

where E gives the cycle count versus the reference epoch and the bracketed quantities indicate the uncertainty in the final digit of the preceding number. This orbital period is 4.5 s (24σ) smaller than the value of 2.838971 (2) d found by F13, explaining why we found the transits of WASP-57 to occur earlier than predicted. There are

Table 4. Times of minimum light and their residuals versus the ephemeris derived in this work. All but one of the timings were derived in the current work, from the source data given in the final column.

Time of min. (BJD/TDB)	Error (d)	Cycle no.	Residual (d)	Source
2454602.18313	0.00317	-513.0	-0.00075	This work (WASP 2008)
2454900.26529	0.00306	-408.0	-0.00504	This work (WASP 2009)
2455351.65767	0.00255	-249.0	-0.00071	This work (WASP 2010)
2455686.65086	0.00055	-131.0	0.00009	This work (TRAPPIST)
2455723.55722	0.00025	-118.0	0.00051	This work (Euler)
2455723.55606	0.00028	-118.0	-0.00065	This work (TRAPPIST)
2456001.76950	0.00035	-20.0	-0.00123	This work (TRAPPIST)
2456018.80454	0.00017	-14.0	0.00030	This work (Euler)
2456018.80315	0.00042	-14.0	-0.00109	This work (TRAPPIST)
2456058.54852	0.00143	0.0	-0.00058	This work (BUSCA u)
2456058.54891	0.00039	0.0	-0.00019	This work (BUSCA g)
2456058.54959	0.00023	0.0	0.00049	This work (BUSCA r)
2456058.54913	0.00042	0.0	0.00003	This work (BUSCA z)
2456728.54099	0.00171	236.0	0.00711	Dittler (TRESKA)
2456796.66754	0.00019	260.0	-0.00038	This work (DFOSC)
2456802.35000	0.00710	262.0	0.00424	This work (Lopresti)
2456819.38110	0.00100	268.0	0.00183	This work (Corfini)
2456819.36760	0.00510	268.0	-0.01167	This work (Marchini)
2456833.57422	0.00017	273.0	0.00036	This work (DFOSC)
2456833.57442	0.00034	273.0	0.00056	This work (GROND g)
2456833.57296	0.00021	273.0	-0.00090	This work (GROND r)
2456833.57389	0.00022	273.0	0.00003	This work (GROND i)
2456833.57460	0.00039	273.0	0.00074	This work (GROND z)
2456836.41480	0.00290	274.0	0.00202	This work (LBarbieri)
2456836.41880	0.00250	274.0	0.00602	This work (Corfini)
2456836.41490	0.00190	274.0	0.00212	This work (Lopresti)
2456836.41960	0.00310	274.0	0.00682	This work (Lopresti)
2456836.41840	0.00290	274.0	0.00562	This work (Marino)
2457171.40240	0.00350	392.0	-0.00277	This work (Lopresti)

several plausible reasons for such a discrepancy to occur, but we are not in a position to choose between them.

The χ^2_ν of the fit is 1.99, and we interpret this as an indication that the uncertainty estimates for the timings are too small. Fig. 4 shows the residuals of the times of mid-transit versus the ephemeris given above. There is no sign of long-term transit timing variations.

Table 5. Parameters of the fit to the light curves of WASP-57 from the JKTEBOP analysis (top). The final parameters are given in bold and the parameters found by F13 are given below this. Quantities without quoted uncertainties were not given by F13 but have been calculated from other parameters which were.

Source	$r_A + r_b$	k	i ($^\circ$)	r_A	r_b
Euler (2011/06/10)	0.1198 ± 0.0084	0.1139 ± 0.0027	87.04 ± 1.10	0.1075 ± 0.0073	0.01225 ± 0.00109
TRAPPIST (2011/06/10)	0.1065 ± 0.0065	0.1087 ± 0.0027	89.94 ± 1.36	0.0961 ± 0.0055	0.01044 ± 0.00073
Euler (2012/04/01)	0.1256 ± 0.0066	0.1190 ± 0.0029	86.56 ± 0.73	0.1123 ± 0.0057	0.01335 ± 0.00099
TRAPPIST (2012/03/15)	0.1436 ± 0.0095	0.1263 ± 0.0026	85.28 ± 0.84	0.1275 ± 0.0082	0.01610 ± 0.00136
TRAPPIST (2012/04/01)	0.1405 ± 0.0153	0.1179 ± 0.0042	85.16 ± 1.20	0.1257 ± 0.0140	0.01482 ± 0.00194
BUSCA g (2012/05/10)	0.1370 ± 0.0131	0.1188 ± 0.0056	85.53 ± 1.12	0.1224 ± 0.0113	0.01455 ± 0.00170
BUSCA r (2012/05/10)	0.1340 ± 0.0057	0.1186 ± 0.0015	85.70 ± 0.52	0.1198 ± 0.0050	0.01421 ± 0.00074
BUSCA z (2012/05/10)	0.1328 ± 0.0098	0.1230 ± 0.0018	85.69 ± 0.84	0.1182 ± 0.0086	0.01454 ± 0.00116
DFOSC (2014/05/18)	0.1309 ± 0.0057	0.1173 ± 0.0018	86.04 ± 0.52	0.1171 ± 0.0049	0.01374 ± 0.00077
DFOSC (2014/06/24)	0.1306 ± 0.0066	0.1166 ± 0.0017	86.05 ± 0.63	0.1170 ± 0.0057	0.01364 ± 0.00086
Final results	0.1278 ± 0.0033	0.1182 ± 0.0013	86.05 ± 0.28	0.1143 ± 0.0029	0.01331 ± 0.00051
F13	0.1122	0.1127 ± 0.0006	$88.0^{+0.1}_{-0.2}$	0.1008	0.01135

4 LIGHT CURVE ANALYSIS

Eight of our light curves cover a full transit at high photometric precision. The Euler and one of the two TRAPPIST light curves from F13 also satisfy this criterion. Each of these ten datasets was modelled separately using the JKTEBOP⁸ code (Southworth et al. 2004) and the *Homogeneous Studies* methodology (Southworth 2012, and references therein). We did not subject those light curves with only partial coverage of a transit to this analysis, because the parameters derived from partial light curves are highly uncertain – so have little effect on the final results – and are often unreliable (e.g. Gibson et al. 2009).

The JKTEBOP model is based on the fractional radii of the star and the planet (r_A and r_b), which are the ratios between the true radii and the semimajor axis ($r_{A,b} = \frac{R_{A,b}}{a}$). The parameters of the fit to each light curve were the sum and ratio of the fractional radii ($r_A + r_b$ and $k = \frac{r_b}{r_A}$), the orbital inclination (i), limb darkening coefficients, and the time of mid-transit. We assumed an orbital eccentricity of zero, based on the finding by F13 that the Lucy & Sweeney (1971) test yielded a 100% probability that the orbit was circular. We fixed the orbital period to the value found in Section 3. We also fitted for the coefficients of a first-order polynomial relating differential magnitude and time (Southworth et al. 2014), in order to allow for any errors in flux normalisation which change with time or airmass. The TRAPPIST light curve obtained on the night of 2012/04/01 was split into two sequences by a meridian flip. This was accounted for by modelling both sequences together but specifying a separate polynomial (of order 1) for each sequence.

Limb darkening (LD) was incorporated using each of five laws (see Southworth 2008), with the linear coefficients either fixed at theoretically predicted values⁹ or included as fitted parameters. We did not calculate fits for both LD coefficients in the four two-coefficient laws as they are very strongly correlated (Carter et al. 2008). The nonlinear coefficients were instead perturbed by ± 0.1 on a flat distribution during the error analysis simulations, to account for the uncertainty in theoretical LD coefficients.

⁸ JKTEBOP is written in FORTRAN77 and the source code is available at <http://www.astro.keele.ac.uk/jkt/codes/jktebop.html>

⁹ Theoretical LD coefficients were obtained by bilinear interpolation in T_{eff} and $\log g$ using the JKTL D code available from: <http://www.astro.keele.ac.uk/jkt/codes/jktdld.html>

Error estimates for the fitted parameters were obtained in four steps. Steps 1 and 2 were residual-permutation and Monte Carlo simulations (Southworth 2008), and the larger of the two alternatives was retained for each fitted parameter. For step 3 we ran solutions using the five different LD laws, and increased the errorbar for each parameter to account for any disagreement between these five solutions. For step 4 we calculated the weighted mean of each photometric parameter using the values found separately from each light curve. This final step is a powerful external check on the reliability and mutual agreement between different datasets, as any discrepancies are obvious and quantifiable.

For all ten light curves we found that it was possible to fit for one of the two LD coefficients: reasonable values for the coefficients were obtained as well as a slightly smaller χ^2_ν compared to fits with both LD coefficients fixed. We therefore adopt these results, which are summarised in Table 5. Detailed tables of results for each light curve are available in the online-only Appendix. The best fits are plotted in Fig. 5.

We find that the results from the different light curves are not in perfect agreement, with a χ^2_ν of 2.0 for $r_A + r_b$ and r_A , 2.8 for r_b and 3.7 for k versus the weighted mean of their values. This is due primarily to the TRAPPIST light curve from F13, which has a very small r_A and high i compared to the other datasets (Table 5). A degeneracy between these parameters is common (e.g. Carter et al. 2008 and Southworth 2008) and arises because these two values together specify the observed transit duration, a quantity which is well-determined by high-quality light curves. If we adopt instead the results from fitting this light curve with both LD coefficients fixed, the agreement becomes much better: k has $\chi^2_\nu = 2.5$, the other four parameters in Table 5 all have $\chi^2_\nu < 0.9$, and all photometric parameters change by less than their 1σ errorbars. We have, however, chosen not to take this step for two reasons. Firstly, theoretical LD coefficients are not perfect – if they were then different sources would give exactly the same values – and none are available calculated specifically for the TRAPPIST $I+z$ filter. Secondly, taking an alternative approach for a discrepant dataset raises the possibility of causing an underestimate of the true uncertainties in the measured quantities. We have therefore retained the discrepant values when calculating the weighted means of the photometric parameters, and have inflated the errorbars on the weighted means by $\sqrt{\chi^2_\nu}$ in order to account for the discrepancy.

The photometric parameters found by F13 differ significantly from our results, by 4.7σ for r_A and 3.8σ for r_b (using our errorbars

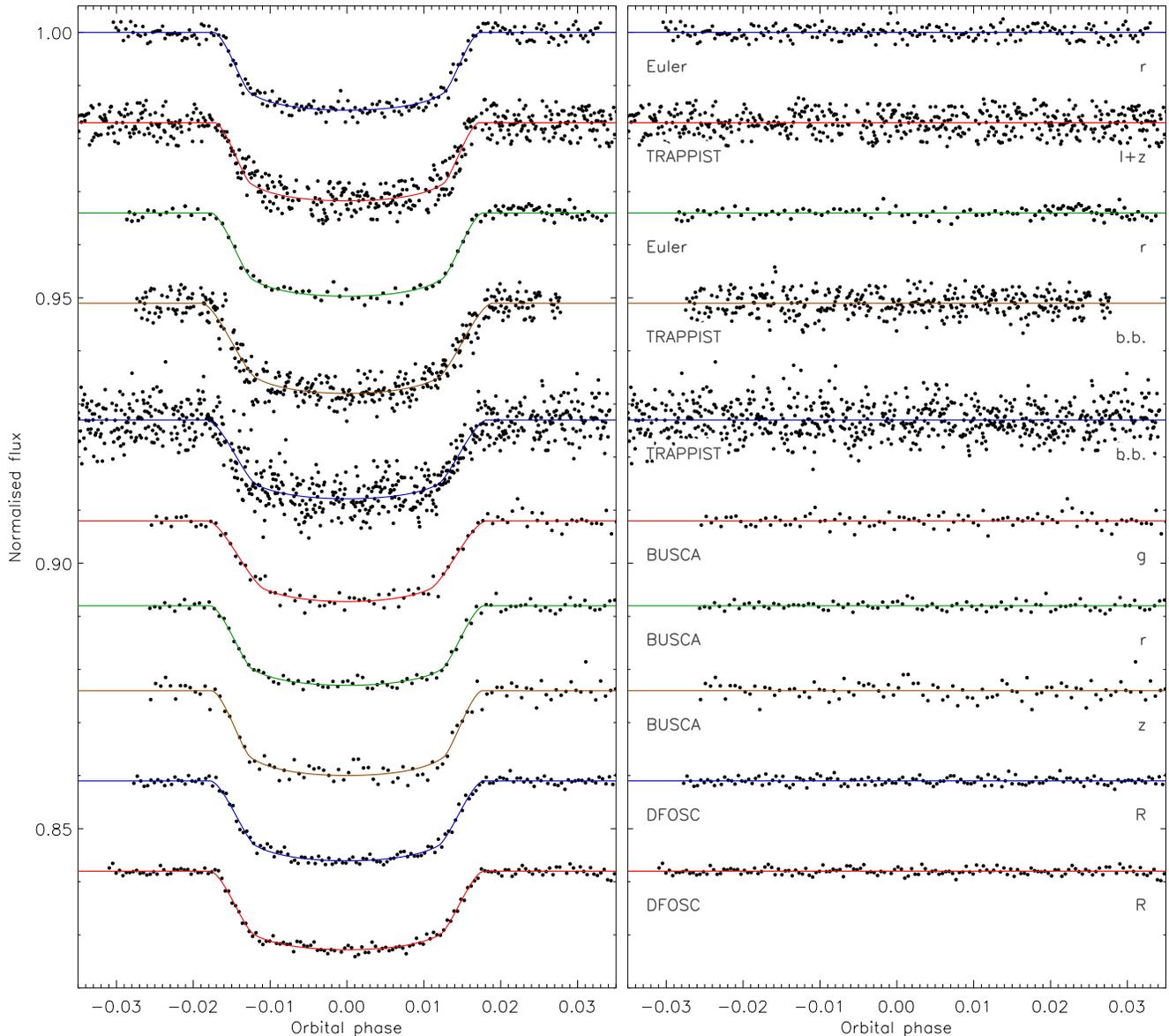


Figure 5. Phased light curves of WASP-57 compared to the JKTEBOP best fits (left) and the residuals of the fits (right). Labels give the source and passband for each dataset. The polynomial baseline functions have been removed from the data before plotting. Only light curves with full coverage of a transit were included in this analysis.

to calculate the σ values as error estimates were not provided by F13 for these two quantities). This is due to the dependence of the F13 solution on only three transit light curves (two complete and one only partially covering a transit), all modelled simultaneously, of which one was the TRAPPIST dataset we find to be discrepant. The value of k found by F13 (0.1127 ± 0.0006) is 3.8σ smaller than ours, and is evidence that the error estimates quoted by F13 are too small (see Southworth 2012 and references therein for other examples). An alternative explanation is the presence of starspots, which is plausible for a star of this temperature. However, no traces of spot occultations are seen in our light curves and no rotational modulation is seen in the long-term WASP light curves.

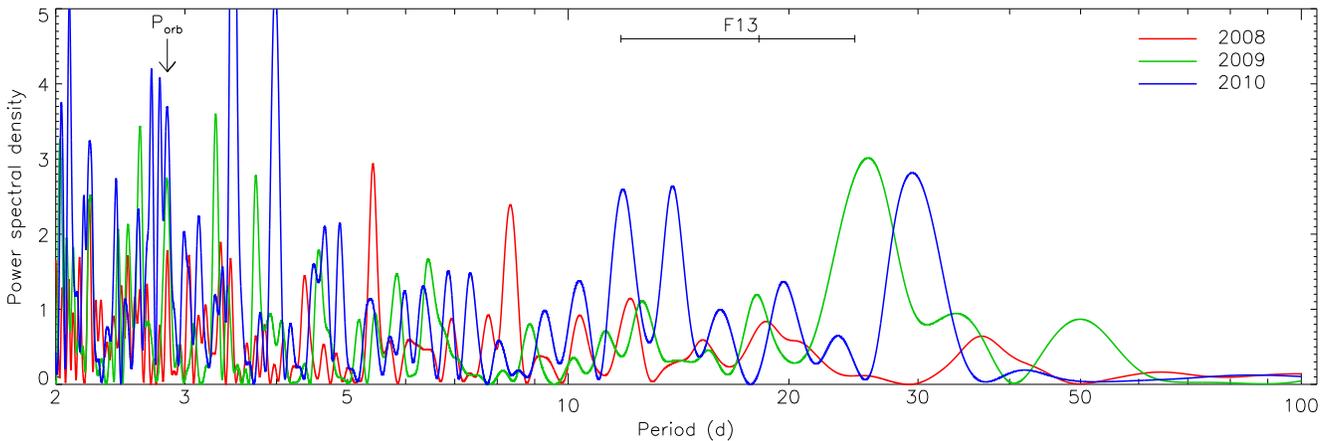
5 PHYSICAL PROPERTIES

We measured the physical properties of the WASP-57 system using the results from Section 4, five grids of predictions from theoretical models of stellar evolution (Claret 2004; Demarque et al. 2004; Pietrinferni et al. 2004; Vandenberg et al. 2006; Dotter et al. 2008), and the host-star spectroscopic properties. Theoretical models provide an additional constraint on the stellar properties, needed because the system properties cannot be obtained from only measured quantities. The spectroscopic properties were obtained by F13 and comprise effective temperature ($T_{\text{eff}} = 5600 \pm 100$ K), metallicity ($[\text{Fe}/\text{H}] = -0.25 \pm 0.10$) and velocity amplitude ($K_A = 100 \pm 7$ m s $^{-1}$). We used the physical constants tabulated by Southworth (2011).

We first estimated the velocity amplitude of the *planet*, K_b , which was used along with the measured r_A , r_b , i and K_A to de-

Table 6. Derived physical properties of WASP-57. The values found by F13 are given for comparison.

Quantity	Symbol	Unit	This work	F13
Stellar mass	M_A	M_\odot	$0.886 \pm 0.061 \pm 0.028$	0.954 ± 0.028
Stellar radius	R_A	R_\odot	$0.927 \pm 0.031 \pm 0.010$	$0.836^{+0.07}_{-0.16}$
Stellar surface gravity	$\log g_A$	cgs	$4.452 \pm 0.024 \pm 0.005$	$4.574^{+0.009}_{-0.012}$
Stellar density	ρ_A	ρ_\odot	1.113 ± 0.085	$1.638^{+0.044}_{-0.063}$
Planet mass	M_b	M_{Jup}	$0.644 \pm 0.060 \pm 0.014$	$0.672^{+0.049}_{-0.046}$
Planet radius	R_b	R_{Jup}	$1.050 \pm 0.052 \pm 0.011$	$0.916^{+0.017}_{-0.014}$
Planet surface gravity	g_b	m s^{-2}	14.5 ± 1.5	$18.3^{+2.9}_{-1.3}$
Planet density	ρ_b	ρ_{Jup}	$0.521 \pm 0.072 \pm 0.006$	$0.873^{+0.076}_{-0.071}$
Equilibrium temperature	T'_{eq}	K	1338 ± 29	1251^{+21}_{-22}
Safronov number	Θ		$0.0522 \pm 0.0045 \pm 0.0006$	
Orbital semimajor axis	a	au	$0.03769 \pm 0.00088 \pm 0.00040$	0.0386 ± 0.0004


Figure 6. Scargle periodograms of the SuperWASP data from the 2008, 2009 and 2010 seasons. The orbital period of the system is shown with a downward-pointing arrow. The stellar rotational period inferred by F13 from its projected rotational velocity is shown by the horizontal error bar.

termine the physical properties of the system (Southworth 2009). The estimate of K_b was then iterated to find the best match between the measured r_A and the calculated $\frac{R_A}{a}$, and the observed T_{eff} and that predicted by a theoretical model for the obtained stellar mass, radius and $[\text{Fe}/\text{H}]$. This was done for a grid of ages from the zero-age main sequence to beyond the terminal-age main sequence for the star, in 0.01 Gyr increments, and the overall best K_b was adopted. The statistical errors in the input quantities were propagated to the output quantities by a perturbation approach.

We ran the above analysis for each of the five sets of theoretical model predictions, yielding five different estimates of each output quantity. These were transformed into a single final result for each parameter by taking the unweighted mean of the five estimates and their statistical errors, plus an accompanying systematic error which gives the largest difference between the mean and individual values. The final results of this process are a set of physical properties for the WASP-57 system, each with a statistical error and a systematic error. The stellar density, planetary surface gravity and planetary equilibrium temperatures can be calculated without resorting to theoretical predictions (Seager & Mallén-Ornelas 2003; Southworth et al. 2007; Southworth 2010), so do not have an associated systematic error.

Bodenheimer et al. (2003, their tables 1 and 2) provided predicted radii for planets of mass $0.69 M_{\text{Jup}}$ and $T'_{\text{eq}} = 1000$ K and 1500 K. Their radii are $1.01\text{--}1.10 R_{\text{Jup}}$ with, and $1.03\text{--}1.13 R_{\text{Jup}}$

without, a core or additional kinetic heating of the planetary interior. Both are in very good agreement with the radius of $1.05 \pm 0.05 R_{\text{Jup}}$ we find for WASP-57 b.

Table 6 contains our measurements of the physical properties of the WASP-57 system. Compared to F13, we find a less massive but larger star. As planetary properties are measured relative to those of their parent star, the planet is similarly affected. The measured planetary density is 3.5σ lower, at $0.521 \pm 0.072 \rho_{\text{Jup}}$ compared to the value of $0.873^{+0.076}_{-0.071} \rho_{\text{Jup}}$ found by F13. Our results are based on a much more extensive set of photometric data so are to be preferred to previous measurements, even though the errorbars have not changed by much. A significant advance in our understanding of the WASP-57 system could be achieved by obtaining further spectroscopy of the host star, from which more precise values for T_{eff} , $[\text{Fe}/\text{H}]$ and K_A could be measured. This is of particular interest because of its metal-poor nature, whose effect on the incidence of different types of planets is currently under discussion (Buchhave et al. 2014; Wang & Fischer 2015).

The age of the system is unconstrained in our analysis above, as often occurs when the host star is significantly less massive than $1 M_\odot$. F13 inferred age estimates of $\gtrsim 2$ Gyr for WASP-57 A from its photospheric lithium abundance, and $\sim 1.9^{+2.4}_{-1.2}$ Gyr from gyrochronological arguments and its rotation period derived from its radius and projected rotational velocity. The T_{eff} of the star is within the regime where starspots are common so we have checked

Table 7. Values of r_b and R_b for each of the light curves. The errorbars exclude all common sources of uncertainty in r_b and R_b , so should only be used to compare different values of $r_b(\lambda)$. The final column gives the size of the errorbar on R_b in atmospheric scale heights.

Instrument	Passband	λ_{cen} (nm)	FWHM (nm)	r_b	R_b (R_{Jup})	σ (H)
Euler	Gunn <i>r</i>	660.0	100.0	0.01362 ± 0.00004	0.867 ± 0.038	8.2
BUSCA	Gunn <i>u</i>	350.0	68.0	0.01099 ± 0.00048	1.061 ± 0.008	1.8
BUSCA	Gunn <i>g</i>	495.5	99.5	0.01345 ± 0.00011	1.063 ± 0.005	1.2
BUSCA	Gunn <i>r</i>	663.0	105.0	0.01348 ± 0.00007	1.100 ± 0.011	2.3
BUSCA	Gunn <i>z</i>	910.0	90.0	0.01395 ± 0.00013	1.057 ± 0.017	3.6
DFOSC	Bessell <i>R</i>	648.9	164.7	0.01343 ± 0.00003	1.060 ± 0.009	2.0
GROND	Gunn <i>g</i>	477.0	137.9	0.01340 ± 0.00021	1.033 ± 0.010	2.2
GROND	Gunn <i>r</i>	623.1	138.2	0.01344 ± 0.00011	1.036 ± 0.017	3.7
GROND	Gunn <i>i</i>	762.5	153.5	0.01310 ± 0.00013	1.059 ± 0.003	0.6
GROND	Gunn <i>z</i>	913.4	137.0	0.01314 ± 0.00022	1.074 ± 0.004	0.8

if it is possible to precisely determine its rotation period from spot-induced modulation. A Lomb-Scargle periodogram was calculated for each of the three seasons of SuperWASP data and can be seen in Fig. 6. There are no strong peaks in the period interval of interest (5–30 d), and no moderately strong peaks present at the same period in all three seasons. We conclude that the rotational modulation of the star is below the level of detection with the current data.

5.1 Comparison with theoretical models of giant planets

F13 found that the measured mass and density of WASP-57 b implied the presence of a heavy-element core of mass roughly $50 M_{\oplus}$, via a comparison to the theoretical predictions of Fortney et al. (2007). This rather large core mass is surprising given the significantly sub-solar metal abundance of the host star ($[\text{Fe}/\text{H}] = -0.25 \pm 0.10$). As we have found a significantly lower density for the planet (smaller by 40% or 3.5σ) it is germane to reconsider this conclusion. We have therefore compared our new mass and radius measurements with predictions based on three batches of theoretical models.

Baraffe et al. (2008, their table 4) find planetary radii of $0.97\text{--}1.06 R_{\text{Jup}}$ for planets of mass $0.5\text{--}1.0 M_{\text{Jup}}$ and age $0.5\text{--}5$ Gyr, again in accord with our results. These values are for a heavy-element fraction of $Z = 0.02$, and larger fractions result in progressively smaller radii and thus poorer agreement with our radius measurement.

Finally, the properties of WASP-57 b match the predictions of Fortney et al. (2007, their fig. 6) for a $25 M_{\oplus}$ heavy-element core. The difference in radius between models with and without this core are only $0.05 R_{\text{Jup}}$ for a $1 M_{\text{Jup}}$ planet and $0.18 R_{\text{Jup}}$ for a $0.3 M_{\text{Jup}}$ planet, so are of a comparable size to the uncertainty in the radius of WASP-57 b. Our measured properties for this planet therefore do not provide significant support for a high metallicity or the presence of a heavy-element core.

6 VARIATION OF RADIUS WITH WAVELENGTH

Two of our datasets include observations in four passbands simultaneously (*ugrz* for BUSCA and *griz* for GROND), whereas we have *r*- or *R*-band photometry from four different sources (DFOSC, EulerCam, BUSCA and GROND), so it is relevant to search for possible changes in the measured radius of the planet as

a function of wavelength. Such analyses are the photometric equivalent of transmission spectroscopy (Seager & Sasselov 2000) and have been pioneered at optical wavelengths by Sing et al. (2011a), de Mooij et al. (2012) and Southworth et al. (2012).

Changes in the radius measured from planetary transits, as a function of wavelength, are predicted to occur due to opacity variations which affect the height at which the atmosphere transmits light coming from the parent star in the direction of the observer. At blue wavelengths a greater atmospheric opacity due to Rayleigh and Mie scattering leads to a higher maximum depth at which starlight is transmitted, causing an increase in the measured radius of the planet (e.g. Pont et al. 2008; Nikolov et al. 2015). Enhanced opacity also leads to signatures of sodium and potassium at optical wavelengths (Fortney et al. 2008), although only narrow absorption cores have been detected so far (Nikolov et al. 2015).

The fundamental observable in this work is the transit depth, represented in our notation by the ratio of the radii k or the fractional planetary radius r_b . The parameter directly comparable to theoretical predictions is the true planetary radius, R_b . The parameter r_b is correlated with other photometric parameters (see e.g. Southworth 2008), and its transformation into R_b requires other parameters which have uncertainty but are common to all photometric passbands (the error budgets calculated in the previous section show that these are r_A , T_{eff} and $[\text{Fe}/\text{H}]$).

We removed these two effects in order to determine R_b values with relative errorbars. We did this by refitting the light curves with all parameters fixed except k , T_0 , the linear LD coefficient for the quadratic LD law, and the coefficients of the polynomials between differential magnitude and time. We then transformed the resulting r_b values into R_b using a fixed orbital semimajor axis, a . This yielded a set of R_b values and errorbars which are directly comparable to each other. The uncertainties in r_b were measured using 1000 Monte Carlo simulations each.

The data included in this analysis were the two Euler light curves (modelled simultaneously), the two DFOSC light curves (modelled simultaneously), the BUSCA *ugrz* and the GROND *griz* data. We did not include the TRAPPIST light curves because the very wide passbands (*I+z* or blue-blocking filters) yield minimal spectral resolution. Although the GROND data only partially cover a transit, which precluded their use in Section 4, they give reliable results here because r_A and i were fixed during the fitting process.

Fig. 7 shows the resulting values of R_b as a function of the central wavelength of the passbands used. The FWHMs of the passbands are shown for reference using horizontal lines. The originat-

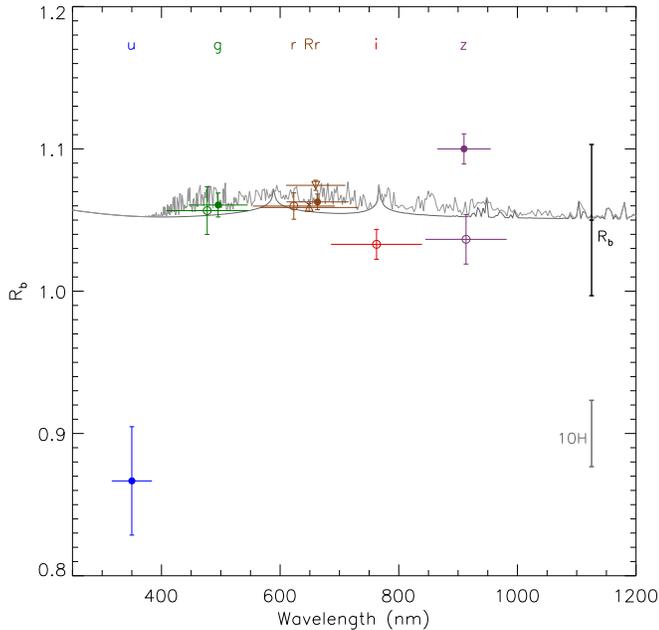


Figure 7. Measured planetary radius (R_b) as a function of the central wavelength of the passbands used for the different light curves. The datapoints show the R_b measured from each light curve. The vertical errorbars show the relative uncertainty in R_b (i.e. neglecting the common sources of error) and the horizontal errorbars indicate the FWHM of the passband. The datapoints are colour-coded according to passband, and the passbands are labelled at the top of the figure. The symbol types are filled circles (BUSCA), open circles (GROND), upward-pointing arrow (DFOSC) and downward-pointing arrow (Euler). On the right of the plot we show the value of R_b measured in Section 4 and the size of ten atmospheric pressure scale heights ($10H$). The grey lines through the empirical datapoints show theoretical predictions for a transmission spectrum of a gas-giant planet of solar chemical composition from Madhusudhan (priv. comm.). The darker-grey line includes features due to Na and K whereas the lighter grey line also includes TiO opacity.

ing r_b values and passband characteristics are collected in Table 7. Two conclusions are immediately apparent from this figure. Firstly, the planetary radius in the u -band is very small and very uncertain. Secondly, the $grRiz$ results are consistent with no variation of R_b with wavelength.

The u -band light curve shows a small transit depth and a high scatter (see Fig. 7), which causes the anomalous r_b measurement in this band. The discrepancy relative to the overall r_b value obtained in Section 4 is highly significant at 4.8σ , and corresponds to approximately $39H$. H is the pressure scale height, and in the case of WASP-57 b is 334 ± 35 km ($0.0047 \pm 0.0005 R_{Jup}$). A variation in R_b of the size of $39H$ is difficult to explain, and is not believable unless confirmed by additional data of much higher quality. The two z -band light curves also show a clear disagreement of size 3.2σ .

Fig. 7 also shows two theoretical transmission spectra calculated under different assumptions by Madhusudhan (priv. comm.) using the atmosphere code of Madhusudhan & Seager (2009). These predictions are for a gas-giant planet of radius $1.25 R_{Jup}$ and surface gravity 25 m s^{-2} so have been scaled to match the smaller radius and lower gravity of WASP-57 b. Our finding of a small u -band r_b is not consistent with these theoretical predictions.

A close-up of the main part of Fig. 7 is shown in Fig. 8. Passband-averaged values for the transmission spectra are shown

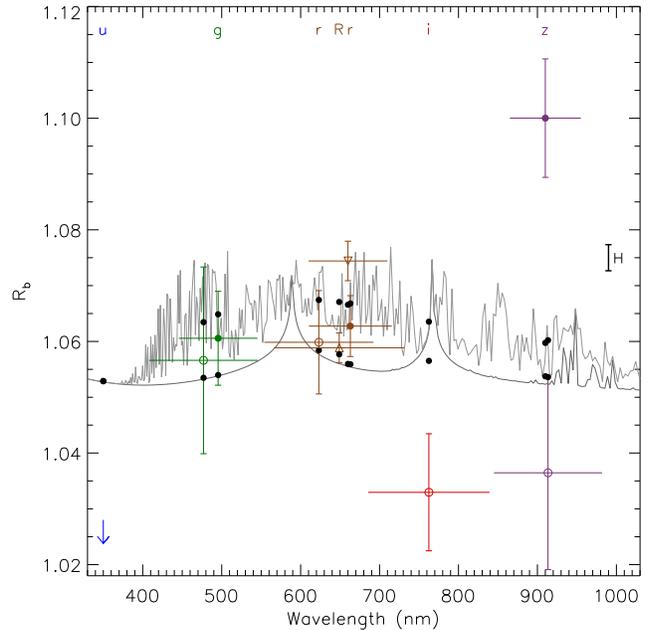


Figure 8. Close-up of the main part of Fig. 7 showing the R_b measurements and theoretical transmission spectra. The u -band result is off the plot. Its central wavelength is indicated with a downward-pointing arrow. The size of one atmospheric scale height is indicated to the right of the plot. The black circles are the values of passband averages of the two transmission spectra, and are shown at the central wavelengths of the relevant passbands.

with black filled circles, and differ by up to $2H$. The Rayleigh scattering slope could be significantly greater than this: Sing et al. (2011b) found that Rayleigh scattering caused the measured radius of HD 189733 b to be larger by $5H$ at 400 nm than at 900 nm.

The relative uncertainties in our measured radii for WASP-57 b are below $1H$ for two, and below $2H$ for five, of the ten light curves (see Table 7). We are therefore sensitive to radius variations at the level of $1H$, which is smaller than both the difference between the two theoretical transmission spectra and the size of the Rayleigh scattering slope detected for HD 189733 b. Our data are therefore sensitive, in principle, to the atmospheric properties of WASP-57 b. However, in practise, our measurements are insufficient for studying the atmosphere of this planet due to the anomalous result for the u -band and the scatter of the radius measurements in the i - and z -bands. The situation could be improved by obtaining data in narrower passbands (i.e. higher spectral resolution), and with repeated observations over the full optical wavelength range. Particular attention should be paid to the u -band, which is an important discriminant between the two transmission spectra and also enhances sensitivity to the Rayleigh scattering slope.

7 SUMMARY AND CONCLUSIONS

WASP-57 b is a relatively low-mass hot Jupiter orbiting a cool star. Amateur astronomers first noticed that its transits were occurring earlier than predicted, a finding subsequently confirmed by observations from professional facilities. We have presented ten transit light curves from amateur astronomers, plus 13 obtained using professional telescopes of which seven predate the discovery of inaccuracy in the orbital ephemeris of the system. We have determined a revised orbital ephemeris which differs by 24σ from the orbital

period in the discovery paper, and can be used to predict transits to a precision of less than 1 minute until the year 2170. We also obtained high-resolution Lucky Imaging observations, which show no evidence for nearby companions whose flux might have contaminated our light curves.

We have used these and previously published data to redetermine the physical properties of the WASP-57 system, finding that both the planet and its host star are larger and less massive than previously thought. A comparison of our new results for WASP-57 b to theoretical predictions for the properties of gaseous planets reveals a good agreement with models lacking a core or additional heat sources. This disagrees with the core mass of $50 M_{\oplus}$ postulated by F13, but is in accord with expectations for a planet which formed around a star of significantly sub-solar metal abundance.

We observed two of the transits of WASP-57 using two 2.2 m telescopes equipped with simultaneous multi-band imaging instruments: GROND (*griz* passbands) and BUSCA (*ugrz* passbands). These data are well-suited to investigating the possible wavelength-dependence of the planet’s measured radius due to effects such as Rayleigh and Mie scattering, and atomic and molecular absorption. Whilst the radii in the *g* and *r/R* bands are in generally good agreement, the *i*-band measurement is slightly smaller than expected and the two *z*-band measurements are discrepant by 3σ . The *u*-band radius is crucial for measuring the Rayleigh scattering slope, as well as separating the pM and pL classes proposed by Fortney et al. (2008). Our measurement is 5σ below theoretical predictions, and the size of the discrepancy is inexplicable using current theoretical transmission spectra. This result is almost certainly spurious, and can plausibly be blamed on the strong absorption by Earth’s atmosphere at blue-optical wavelengths plus the faintness of the host star in this passband.

7.1 Future opportunities for transmission photometry

Successful detections of radius variations in optical transmission photometry have recently been announced for the TEPs GJ 3470 b (Nascimbeni et al. 2013; Biddle et al. 2014), Qatar-2 b (Mancini et al. 2014) and WASP-103 b (Southworth et al. 2015). Only one of these studies presented data obtained shortward of the Balmer jump, which is an important but observationally difficult wavelength interval (see Fig. 8).

Transit light curves in the *u* and *U* bands have previously been presented for several TEPs, and have shown planetary radii either consistent with other optical passbands (WASP-12, Copperwheat et al. 2013; TrES-3, Turner et al. 2013; WASP-17, Bento et al. 2014; WASP-39 and WASP-43, Ricci et al. 2015; XO-2, Zellem et al. 2015) or somewhat larger than other optical passbands (HAT-P-5, Southworth et al. 2012, J. Dittman 2012, priv. comm.; GJ 3470, Nascimbeni et al. 2013). A universal feature of these studies is the reliance on either a single *u*-band transit light curve, which yields large uncertainties on the measured planetary radius, or the use of data not obtained simultaneously in multiple passbands, so the results are hostage to temporal changes such as induced by magnetic activity in the host stars. Most transmission photometry studies also suffer from the use of wide passbands, which are insensitive to spectral features other than broad continuum slopes (see Nikolov et al. 2013).

Whilst suffering from a lower spectral resolution, transmission photometry has several advantages over transmission spectroscopy. These include being able to observe over a wide wavelength interval without being subject to second-order contamination, the ability to use comparison stars more distant from the planet host star, and

the option to use telescope defocussing techniques to avoid systematic noise (but see Burton et al. 2015, for a counter-example). Smaller telescopes can be used, making it easier in particular to observe multiple transits and thus demonstrate the repeatability of the experiment (Bean et al. 2013; Gibson 2014).

We therefore advocate studies based on observations of multiple transits, obtained simultaneously through many intermediate or narrow passbands. These passbands should be well-defined by interference filters, thus avoiding compromises such as the variable red edge of the *z* filter due to its reliance on the quantum efficiency curve of the CCD used (e.g. Fukugita et al. 1996) or the red leak in some *u* and *U* filters which is capable of causing spurious results for optically-blue objects (e.g. Guhathakurta et al. 1998). With a sufficient number of passbands, it should be possible to achieve low-resolution spectroscopy of the atmospheres of extrasolar planets through the full optical wavelength range using the transmission-photometry approach.

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