

Structure and Dynamics of Candidate O Star Bubbles in N44

Yaël Nazé^{1,5}, You-Hua Chu^{2,6}, Martín A. Guerrero^{2,6}, M. S. Oey³, Robert A. Gruendl², R. Chris Smith⁴

ABSTRACT

Dynamical studies of superbubbles and Wolf-Rayet ring nebulae show discrepancies from the standard, adiabatic model for wind-blown bubbles. We therefore study the physical properties and kinematics of three candidate bubbles blown by single O stars, to evaluate whether these discrepancies are also found in these simpler objects. Our sample candidates are N44 F, N44 J, and N44 M, in the outskirts of the H II complex N44 in the Large Magellanic Cloud. We have obtained ground-based and *HST* emission-line images and high dispersion echelle spectra for these objects. From the H α luminosities and the [O III]/H α ratios of these nebulae, we estimate the spectral types of the ionizing stars to be O7V, O9.5V and O9.5V for N44 F, N44 J, and N44 M, respectively. We find that the observed expansion velocity of 12 km s⁻¹ for N44 F is consistent with the stellar wind luminosity expected from the central ionizing star, as predicted by the standard bubble model. The observed upper limits for the expansion velocities of N44 J and N44 M are also compatible with the expected values, within the uncertainties. We also report the discovery in N44 F of strongly-defined dust columns, similar to those seen in the Eagle Nebula. The photoevaporation of these dense dust features may be kinematically important and may actually govern the evolution of the shell. The inclusion of photoevaporation processes may thus undermine the apparent agreement between the observed bubble dynamics and the simple adiabatic models.

Subject headings: H II regions— ISM: bubbles— ISM: kinematics and dynamics— ISM: individual (N44)— Magellanic Clouds

1. Introduction

Mechanical feedback from massive stars is a fundamental driver of galaxy evolution. The supernovae (SNe) and supersonic stellar winds from these stars generate bubbles and shells in the interstellar medium (ISM), which may dominate

the ISM structure formation in many star-forming galaxies. The standard, adiabatic model for these shells and superbubbles (e.g., Weaver et al. 1977) predicts the production of hot, 10⁶ – 10⁷ K, low-density gas within these shells, which bear the products of SN and massive star nucleosynthesis. Therefore, these bubble structures are thought to be the source of the diffuse, hot ionized component of the ISM. Vigorous star-forming regions may generate supergiant shells that blow out of galaxy disks, thereby dispersing metals and mass far from the parent star formation event. In addition, this mechanical energy contributes to ISM kinematics and is thought to be a primary source of turbulence (e.g., Norman & Ferrara 1996; Goldman 2000).

The effects of mechanical feedback have been clearly demonstrated, for example, in starburst

¹Institut d’Astrophysique et de Géophysique, Allée du 6 Août 17, Bat. B5c, B 4000 Liège (Sart-Tilman), Belgium; naze@astro.ulg.ac.be

²Astronomy Department, University of Illinois, 1002 W. Green Street, Urbana, IL 61801, USA; chu@astro.uiuc.edu, mar@astro.uiuc.edu, gruendl@astro.uiuc.edu

³Lowell Observatory, 1400 West Mars Hill Rd., Flagstaff, AZ 86001, USA; oey@lowell.edu

⁴Cerro Tololo Inter-American Observatory, Casilla 603, La Serena, Chile; csmith@noao.edu

⁵Research Fellow FNRS (Belgium)

⁶Visiting astronomer, Cerro Tololo Inter-American Observatory

galaxies (Johnson et al. 2000; Strickland et al. 2000) and in superbubbles around OB associations (e.g., Saken et al. 1992; Oey 1996). However, several troubling discrepancies remain unsolved. For example, almost all clean examples of bubbles blown by Wolf-Rayet stars and of superbubbles present shells that are too small for their observed parent stellar population (e.g., Treffers & Chu 1982; Drissen et al. 1995). Many superbubbles also show anomalous kinematics and X-ray emission (e.g., Rosado et al. 1982; Chu & Mac Low 1990; Wang & Helfand 1991). To resolve these puzzles and more clearly understand the mechanical feedback process, it is therefore necessary to examine simpler systems.

In active star-forming regions, high concentrations of OB stars collectively produce superbubbles. But in their peripheries, where stars are loosely distributed, discrete bubbles may be produced by individual stars. $H\alpha$ images of H II regions ionized by OB associations in the Large Magellanic Cloud (LMC) have revealed large (50–150 pc) gas shells around the OB associations and numerous small (< 15 pc), ring-like nebulae outside the large shell (see, e.g., Davies, Elliott, & Meaburn 1976). These small nebulae are candidate wind-blown bubbles of single O or B stars (Weis & Duschl 1999), and offer perhaps the simplest test of the standard, adiabatic bubble model.

We have investigated the kinematics and physical nature of three ring-like nebulae in N44 (Henize 1956) in the LMC. N44 is a bright H II complex with a $44 \text{ pc} \times 67 \text{ pc}$ superbubble at the center and several compact H II regions along the shell rim. In the surrounding area, small dense H II regions as well as extended diffuse nebulae are also present at distances up to 160 pc. Figure 1 shows these features and their identifications given by Henize (1956). Three OB associations exist in N44 (Lucke & Hodge 1970): LH 47 in the central shell, LH 48 in the compact H II region N44I, and LH 49 partially embedded in the bright H II region N44D. The three nebulae that we have studied are N44F, N44J, and N44M. N44F is a bright, circular H II region at the northwest rim of the superbubble and is adjacent to a bright filament connected to the main nebula. In contrast, N44J and N44M, appearing respectively on the northern and eastern parts of N44, are fainter and relatively isolated. None of these three nebulae have been studied

since their original identification by Henize (1956).

We have obtained ground-based CCD images and high-dispersion echelle spectra of N44F, N44J, and N44M to study the morphology and kinematics of these nebulae. Additionally, we have used *Hubble Space Telescope* (*HST*) WFPC2 images of N44F in the $H\alpha$ and [S II] lines to analyze its ionization structure. In this paper, we first describe these observations in §2, then present the analysis and results for N44F in §3 and N44J and N44M in §4. We present our conclusions in §5.

2. Observations

The datasets used in this study include: (1) CCD images taken with the Curtis Schmidt and 0.9 m telescopes at Cerro Tololo Inter-American Observatory (CTIO), (2) CCD images taken with the *HST* WFPC2, and (3) high-dispersion echelle spectra taken with the CTIO 4 m telescope. The high-dispersion spectra are useful to diagnose expanding shells that are not morphologically identifiable.

2.1. Ground-based CCD Images

CCD images of N44 were taken with the Curtis Schmidt telescope at CTIO in 1993 December. The detector was a front-illuminated Thomson 1024×1028 CCD camera with $19 \mu\text{m}$ pixels, or $1''.835 \text{ pixel}^{-1}$. Images were obtained with $H\alpha$ ($\lambda_c=6561\text{\AA}$, $\Delta\lambda=26\text{\AA}$), [O III] ($\lambda_c=5010\text{\AA}$, $\Delta\lambda=50\text{\AA}$), and [S II] ($\lambda_c=6720\text{\AA}$, $\Delta\lambda=50\text{\AA}$) filters for exposure times of 1800 s, 2700 s, and 2700 s, respectively. The $H\alpha$ filter includes not only the $H\alpha$ line but also the neighboring [N II] $\lambda\lambda 6548, 6583$ lines. At the LMC's radial velocity, $\sim 300 \text{ km s}^{-1}$, the [N II] $\lambda 6583$ line is red-shifted further away from the red edge of filter, but the [N II] $\lambda 6548$ line is red-shifted further into the $H\alpha$ bandpass. Using the [N II]/ $H\alpha$ intensity ratio measured from our echelle spectra (see Section 2.3), we estimate that the [N II] $\lambda 6548$ line emission contributes 2.4%, 4.3% and 3.3% of the total flux of N44F, N44J and N44M, respectively. This [N II] contamination has been corrected in our $H\alpha$ flux measurements.

Higher-resolution $H\alpha$ images were taken with the 0.9 m telescope at CTIO in 2000 December. The detector was a thinned, AR coated SITE 2048 \times 2048 CCD with $24 \mu\text{m}$ pixels, or

$0''.4 \text{ pixel}^{-1}$. The $\text{H}\alpha$ filter ($\lambda_c=6559\text{\AA}$, $\Delta\lambda=64\text{\AA}$) used transmits $\text{H}\alpha$ and both $[\text{N II}]$ lines, therefore the “ $\text{H}\alpha$ ” images are really $\text{H}\alpha+[\text{N II}]$ images, and thus are only used for morphology analysis. The angular resolution, determined by the FWHM of the stellar images in the field of view, was approximately $1''.3$. To map the entire N44 complex, five fields were observed with exposure times of $2\times 600 \text{ s}$ for each field.

All CCD images were bias-subtracted and flat-fielded in a standard way using IRAF software, and multiple frames were combined to remove cosmic rays and improve the S/N ratio. Figure 1 shows the $\text{H}\alpha$ mosaic of the N44 complex, and Fig. 2 presents close-ups of N44 F, N44 J and N44 M in the $\text{H}\alpha$, $[\text{O III}]$, and $[\text{S II}]$ lines.

2.2. *HST* WFPC2 Images

HST WFPC2 images of N44 F were acquired via the Hubble Observation Problem Report ⁷ (HOPR) observations of the Cycle 6 program 6698. In these observations, made with the *F656N* filter for $2\times 500 \text{ s}$ and the *F673N* filter for $2\times 600 \text{ s}$, N44 F was included in the Wide-Field Camera 3 (WF3). The *F656N* filter, centered at 6563.7\AA with a FWHM of 21.4\AA , includes the $\text{H}\alpha$ line and the neighboring $[\text{N II}]$ lines. Using the $[\text{N II}]/\text{H}\alpha$ intensity ratio measured from our echelle spectra (see Section 2.3), we estimate that the $[\text{N II}] \lambda 6548$ line emission contributes 2.4% of the flux in N44 F. This $[\text{N II}]$ contamination has been corrected in our $\text{H}\alpha$ flux measurements. The *F673N* filter, centered at 6732.2\AA with a FWHM of 47.2\AA includes the nebular lines of $[\text{S II}] \lambda\lambda 6716, 6731$.

The calibrated WFPC2 images were produced by the standard *HST* pipeline. We processed them further with IRAF and STSDAS routines. The images taken with the same filter were combined to remove cosmic rays and to produce a total-exposure map. The combined $[\text{S II}]$ and $\text{H}\alpha$ images were then each corrected for the intensity- and position-dependent charge transfer efficiency as

⁷The original observations of N44 were made with an incorrect orientation, resulting in a loss of 1/3 of the field of view. To compensate the loss, HOPR observations were granted. To facilitate the scheduling of these HOPR observations and to maximize the scientific yield, a different pointing center in N44 suitable for arbitrary orientations was chosen.

explained by Holtzman et al. (1995). Following the recommended procedures for narrowband WFPC2 photometry⁸, we find the rectangular widths of the *F656N* and *F673N* filters to be 28.3\AA and 63.3\AA , respectively, and use these widths to convert the flux densities into fluxes.

Figure 3 shows the WFPC2 images of N44 F. To show the excitation variations, the $[\text{S II}]/\text{H}\alpha$ ratio map of N44 F is also presented, in addition to the $\text{H}\alpha$ and $[\text{S II}]$ images.

2.3. CTIO 4 m Echelle Spectra

Echelle spectra were obtained with the CTIO 4 m telescope in 2000 December. The spectrograph was used with a 79 line mm^{-1} echelle grating and the long-focus red camera. We observed a single order centered on the $\text{H}\alpha$ line by inserting a post-slit $\text{H}\alpha$ interference filter and replacing the cross-dispersing grating with a flat mirror. A Tek 2048×2048 CCD with $24 \mu\text{m pixel}^{-1}$ was used to record the images. This provided a spectral sampling of $\sim 0.08 \text{\AA pixel}^{-1}$ and a spatial sampling of $0''.26 \text{ pixel}^{-1}$. The wavelength coverage, limited by the $\text{H}\alpha$ filter and the echelle order, was 125\AA ; the spatial coverage, limited by the optics, was $\sim 200''$. The angular resolution, determined by the FWHM of the seeing, was approximately $1''$. A slitwidth of $1''.65$ was used. The resultant spectral resolution, measured from the Th-Ar lamp lines, was about 13 km s^{-1} FWHM.

Each of the three nebulae, N44 F, N44 J, and N44 M, was observed with an east-west oriented slit. The exposure times were 600 s for N44 F, but only 300 s for N44 J and N44 M. The spectral lines detected include nebular $\text{H}\alpha$ and $[\text{N II}]$ lines, and telluric $\text{H}\alpha$ and OH lines (Osterbrock et al. 1996). We have adopted rest wavelengths of 6562.7885\AA , and 6548.080\AA and 6583.454\AA for the $\text{H}\alpha$ and $[\text{N II}] \lambda\lambda 6548, 6583$ nebular lines, respectively (Spyromilio 1995).

3. Physical Structure of N44 F

We will first describe the overall morphology of N44 F, examine its excitation and ionization structure, and derive its rms electron density. Special emphasis is given to the prominent dust “pillars” where star formation might be taking place. Next,

⁸available at http://www.stsci.edu/instruments/wfpc2/Wfpc2_faq/wfpc2_nrw_phot.

we describe the kinematic features detected in the echelle spectra, fit Gaussian components to the line profiles, and use the resultant velocities to identify expanding structures and to determine expansion velocities. Finally, we combine the morphological and kinematic information and present an integrated view of this H II region. Throughout this paper, we will adopt a distance of 50 kpc to the LMC (Feast 1999).

3.1. Morphology and Density

N44F is a bright circular nebula nestled on a bright filament at the northwest rim of the superbubble in N44. In low-resolution H α images, N44F appears to have a double-ring structure with a very bright region in the northwest. The details of these features are resolved in the *HST* WFPC2 images. The outer ring, of radius $\sim 22''$, is surrounded by dusty features. Finger- or pillar-like dust columns, similar to those in the Eagle Nebula (=M16; Hester et al. 1996), are also seen at the periphery, with two especially prominent ones at the west, close to the brightest region of N44F. These dust columns are revealed in detail by the *HST* images, and will be discussed more extensively in Section 3.2. The inner ring, of radius $\sim 6''$, is resolved into diffuse emission without sharp features. The western part of this ring merges into the brightest region of N44F. A group of bright stars are located within the central ring, but they are substantially offset towards the bright emission region at the northwest. Unfortunately, none of the stars in the vicinity of N44F have been studied photometrically or spectroscopically.

The [S II] images of N44F show a very different morphology: only the outer ring is visible and it is mostly composed of sharp filaments. The pillar-like features are also present and stars at the tips of two of them are clearly seen (see Fig. 4). The [S II]/H α ratio in the central region is very low, about 0.1. The outer ring possesses a higher [S II]/H α ratio, with values ranging from 0.2 to 0.4. Enhanced [S II]/H α ratios are often associated with shell structures or high velocity shocks. As we do not find the latter in the nebula (see Section 3.3), the somewhat elevated [S II]/H α ratio is likely to result from the low ionization parameter⁹

⁹The ionizing flux derived further can be used to estimate the ionization parameter $U \propto (Q(H^0)n f^2)^{1/3}$, where

caused by the shell morphology. The values of [S II]/H α are in fact similar to those seen in other shell structures (e.g., Lasker 1977; Hunter 1994).

Using the Curtis Schmidt images, we also examine the [O III] morphology of N44F. Because of the poorer resolution, only the overall morphology can be described. The [O III] image of N44F looks very similar to that in H α , except for the bright H α outer ring, which disappears almost completely in [O III]. We have used the *HST* WFPC2 observations of N44C (Garnett, Galarza, & Chu 2000) to calibrate the observed [O III]/H α ratio derived with the Curtis Schmidt images. The [O III]/H α ratio of N44F is relatively high, about 0.5–1.2 inside the ring, with the highest value near the star(s). The outer ring has the lowest [O III]/H α ratio, about 0.25. The integrated [O III]/H α ratio for the whole nebula is about 0.55. The high ratio near the central region indicates that the ionizing source must be an early-type massive star.

The WFPC2 H α image can be further analyzed to estimate the rms density of N44F. A color excess of $E(B-V)=0.11$ mag for N44 (Oey & Massey 1995) is adopted in the extinction correction. The emission measure (EM) derived from the extinction-corrected surface brightness (Peimbert, Rayo, & Torres-Peimbert 1975) ranges from $7 \times 10^3 \text{ cm}^{-6} \text{ pc}$ in the faintest interior regions of N44F to $4 \times 10^4 \text{ cm}^{-6} \text{ pc}$ at the tips of the ‘‘pillars’’. To derive the density from the EM, we can assume two simple geometries. First, the density can be estimated by assuming the depth to be equal to the diameter of the nebula: for the central region, we find a value ranging from 25 cm^{-3} to 35 cm^{-3} . Second, N44F is limb-brightened and we can therefore estimate the density by assuming a constant density spherical shell with radius R and thickness ΔR . This geometry predicts a unique radial surface brightness profile for any $\Delta R/R$ based on the varying path length through the shell material. Comparing these simple modeled profiles to the observed H α surface brightness profiles near the shell rim, we find a shell thickness $\Delta R/R$ of ~ 0.25 . Using the peak surface brightness and the length of the longest line of sight,

$Q(H^0)$ is the ionizing flux in photons s^{-1} , n the density in cm^{-3} , and f the filling factor. Using $Q(H^0)=1.0 \times 10^{49}$ photons s^{-1} , $n=40 \text{ cm}^{-3}$ and $f=1 - \left(1 - \frac{\Delta R}{R}\right)^3 \sim 0.6$, we find a value of $\sim 10^{-3}$ for the ionization parameter.

$2\sqrt{2R\Delta R - (\Delta R)^2}$, we then derived a density in the shell of 40 cm^{-3} . On average, the diffuse gas of N44 F probably possesses a density of a few tens cm^{-3} .

Using the Curtis Schmidt $\text{H}\alpha$ image and the flux calibration from Kennicutt & Hodge (1986), we measured over the entire N44 F nebula an $\text{H}\alpha$ flux of $3.6 \times 10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$. For a distance of 50 kpc, the $\text{H}\alpha$ luminosity of N44 F is then $1.4 \times 10^{37} \text{ erg s}^{-1}$, after correcting for an extinction of $E(B-V)=0.11 \text{ mag}$. This $\text{H}\alpha$ luminosity requires an ionizing flux of $Q(H^0)=1.0 \times 10^{49} \text{ photons s}^{-1}$. A second evaluation of the extinction correction has been made by comparing radio data and the Curtis Schmidt $\text{H}\alpha$ image: a mean value $E(B-V)$ of 0.28 mag for N44 F, N44 J, and N44 M is found (J. R. Dickel, private communication), and this results in a $Q(H^0)$ of $1.4 \times 10^{49} \text{ photons s}^{-1}$ for N44 F. This value suggests that the total ionizing power of the star(s) within N44 F is equivalent to a single O6.5 V–O7.5 V star (Panagia 1973; Vacca, Garmany, & Shull 1996; Schaerer & de Koter 1997). This is compatible with the rather high $[\text{O III}]/\text{H}\alpha$ ratio: the photoionization models of Stasińska & Schaerer (1997) for CoStar stellar atmosphere models suggest that the observed integrated $[\text{O III}]/\text{H}\alpha$ ratio of N44 F could be produced by the equivalent of an O8 V–O9 V star¹⁰. The comparison to the photoionization models should however be interpreted with caution because the models do not assume a shell structure, which affects the excitation and ionization structure. The dominant spectral type could therefore be somewhat earlier than inferred from the modeled excitation.

3.2. Pillar-like Dust Columns

The dark clouds of N44 F exhibit new examples of the intriguing finger- or pillar-like structures, similar to those reported in the Eagle Nebula (Hester et al. 1996). The most prominent two are embedded in the bright northwest region (see Figs. 3 and 4). Several additional features of this type are present in the south and possibly in the southwest, but they are fainter. In Fig. 4, we present

¹⁰We have compared N44 F, N44 J, & N44 M to the closest model available in Stasińska & Schaerer (1997), i.e. a spherical nebula of density 10 cm^{-3} and metallicity 0.25 Z_{\odot} , ionized by one single star.

the most striking of these structures in the $\text{H}\alpha$ and $[\text{S II}]$ images as well as in the ratio map. The dust columns in N44 F are $5''$ to $10''$ long and $1''$ to $2''$ wide, corresponding to 1 to 2 pc long and 0.25 to 0.5 pc wide at the LMC’s distance. These sizes and morphologies are similar to those seen in 30 Doradus (Scowen et al. 1998) and M16 (Hester et al. 1996). Two of these columns harbor stars at their tips, as indicated by arrows in Fig. 4. Hester et al. (1996) suggested that stars at the tips of pillars could be young stellar objects which were formed in the molecular cloud and finally emerge after the photoevaporation of most of their parent cloud. However, the Hester et al. (1996) model of pre-existing condensations in the molecular cloud does not provide the unique conditions for producing photoionized pillars. Hydrodynamical simulations by Williams, Ward-Thompson, & Whitworth (2001) illustrate that these structures can be formed by a variety of initial conditions in the structure with inhomogeneities in density and/or radiation field, but current observations do not yet permit discrimination among these different models. It is therefore of great interest to identify additional examples of such objects.

Other important characteristics of these dust columns are their ionization structure and their density. All of the dust columns point towards the central ionizing source, i.e., a bright star or a small cluster between the center and the northwest edge of N44 F. As can be seen in Fig. 5, the $[\text{S II}]$ peak is separated from the $\text{H}\alpha$ peak by approximately $0''.1$, i.e., 0.024 pc, with $\text{H}\alpha$ peaking slightly closer to the source of UV radiation which is photoevaporating the cloud. The shape of these profiles is very similar to those observed in the pillars of 30 Doradus, or in ground-based images of M16 (Scowen et al. 1998). Unfortunately, no WFPC2 $[\text{O III}]$ images of N44 F are available, thus we can not undertake a complete analysis using photoionization models or constrain further the properties of the ionizing source. Following the method used in section 3.1, we estimate an rms density larger than 300 cm^{-3} at the tip of one of the pillars. This density is compatible with that determined by Scowen et al. (1998) in 30 Doradus, 700 cm^{-3} . These pillars are thus photoevaporating clouds, whose evaporated matter enhances the density and brightness of the diffuse material seen in the northwest part of the nebula.

3.3. Kinematic Properties

The kinematic properties of N44F are derived from the high-dispersion echelle observations of the H α and [N II] λ 6583 lines. The [N II] line has a smaller thermal width than the H α line, thus it shows the line splitting better although it is noisier. We have extracted H α and [N II] line profiles and fitted them with Gaussian components. The resultant velocity-position plot (Fig. 6) shows that N44F is an expanding shell with a radius of 22'', coincident with the outer ring described in §3.1. Both lines indicate a shell expansion velocity of ~ 12 km s $^{-1}$. An example of the [N II] and H α profiles are presented in Fig. 7. The approaching (blueshifted) component is fainter, suggesting a lower density. Further east, outside N44F, an additional blue component is present, but it is probably associated with the complex filamentary structure at the outskirts of N44 and is not directly related to N44F.

3.4. Test of the Standard Bubble Model

N44F has a classical ring-nebula morphology, and its kinematics confirms an expanding shell, most likely created by the wind of the early-type massive star(s) within this bubble. We can now test whether the observed expansion dynamics of this bubble are consistent with those expected from the standard, adiabatic bubble model (e.g., Weaver et al. 1977). With the observed radius R in units of pc and expansion velocity V in units of km s $^{-1}$, we can compute the dynamical timescale t_6 in units of 10^6 yr using $t_6 = 0.6R/V$, and further determine the ratio of stellar wind luminosity L_{36} in units of 10^{36} ergs s $^{-1}$ to ambient density n_0 in units of H-atom cm $^{-3}$, using $R = 27 \left(\frac{L_{36}}{n_0}\right)^{1/5} t_6^{3/5}$.

For a radius of 5.3 pc and an expansion velocity of ~ 12 km s $^{-1}$, N44F has a dynamical timescale of 2.7×10^5 yr and a L_{36}/n_0 ratio of 0.016 ergs s $^{-1}$ cm 3 . If we assume that the shell consists of mass swept up from the volume, the ambient density, n_0 , can be estimated by $n_0 = 3n \frac{\Delta R}{R}$, where n is the rms density of the shell, ΔR the shell thickness, and R the shell radius. For $n \sim 40$ cm $^{-3}$ and $\frac{\Delta R}{R} = 0.25$ (see Section 3.1), we derive an ambient density of ~ 30 cm $^{-3}$ and a wind luminosity of $\sim 5 \times 10^{35}$ ergs s $^{-1}$. This wind luminosity is compatible with that expected from the spectral

type determined in Section 3.1, i.e., an O7V star (de Jager et al. 1988; Prinja, Barlow, & Howarth 1990). One uncertainty in these estimates is the shell thickness; however, even if the shell thickness is off by a factor of two, the predicted wind luminosity will still be compatible, within the errors, with an O7V star.

Thus, the dynamics of N44F appear to be consistent with the standard adiabatic model for bubble evolution, within uncertainties generated by the complex interstellar environment. This contrasts with the results from superbubbles, which show evidence of retarded growth (e.g., Oey 1996). To investigate further this discrepancy, it is worth considering an alternative bubble model.

The presence of dust pillars around N44F clearly indicates a highly inhomogeneous ISM. The photoevaporation of the dense clumps may serve as an energy sink and govern the dynamic evolution of this H II region. Since photoevaporation processes take place in the early evolutionary stages of a bubble/superbubble, the energy loss by photoevaporation may explain the smaller-than-expected sizes of previously studied superbubbles and Wolf-Rayet ring nebulae (see §1). Therefore, we will now consider models of bubbles formed in a cloudy medium by McKee, Van Buren, & Lazareff (1984). In their models, only bubbles blown by stars with weak winds will follow the Weaver et al. model, while stars with stronger winds will produce bubbles whose expansion is governed by the dynamics of photoevaporating clouds. They define a dimensionless wind luminosity $L_W^* \equiv L_W / (1.26 \times 10^{36} (S_{49}^2 / n_m)^{1/3})$ and a homogenization radius $R_h(\text{pc}) \equiv 0.009 t^{4/7} \left(\frac{S_{49}}{n_m}\right)^{1/7}$, where L_W is the wind luminosity in erg s $^{-1}$, S_{49} the ionizing flux in photons s $^{-1}$, n_m the mean cloud density in cm $^{-3}$ and t the dynamical timescale in years. Strong winds possess $L_W^* \gtrsim 1$ and blow bubbles of radius $\sim R_h$.

Using the observed parameters of N44F (with $n_m = n_0$) and assuming a stellar wind luminosity for an O7V star, we estimate $L_W^* \sim 1$ and $R_h \sim 4.5$ pc, close to the observed bubble's radius. This seems to suggest that the dynamics of photoevaporation actually governs the bubble's expansion. The agreement of the data with both models could be a simple coincidence, or maybe the bubble has just encountered the cloudy medium and

begun the photoevaporation processes, in which case there may not be a large deviation from the Weaver et al. model. Note, however, that all these conclusions hinge upon the validity of the stellar wind luminosity, which is based on an assumed spectral type of the central star. Future study of the actual stellar content of N44F is needed to place stronger constraints on the applicable bubble models.

4. Physical Structure of N44 J and N44 M

4.1. N44 J

Situated in the northernmost part of the N44 complex, N44 J is a slightly elongated, ring-like nebula with a size of $37'' \times 32''$ (or $9.0 \text{ pc} \times 7.8 \text{ pc}$). The 0.9 m $H\alpha$ image shows a bright rim around dusty features at the northwest boundary. A bright central star is visible in the $H\alpha$ and [S II] images (see Fig. 2). In $H\alpha$ and [S II], the nebula exhibits modest limb-brightening, with the edge-to-center brightness ratio < 2 , while its central part displays a rather uniform surface brightness. The [S II]/ $H\alpha$ ratio is very similar to that of N44 F, ranging from 0.25 to 0.4. In contrast, the [O III] image of N44 J shows emission only near the central star. The [O III]/ $H\alpha$ ratio is 0.4 near the star, and only 0.05–0.08 elsewhere, significantly lower than the corresponding value in N44 F. The integrated [O III]/ $H\alpha$ ratio of N44 J is about 0.12.

The observed $H\alpha$ flux of N44 J is $5.7 \times 10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$. This requires an ionizing flux $Q(H^0)$ of $1.6 \times 10^{48} \text{ photons s}^{-1}$ when adopting $E(B-V)=0.11 \text{ mag}$, or $2.3 \times 10^{48} \text{ photons s}^{-1}$ for $E(B-V)=0.28 \text{ mag}$. These values are expected for an O9V–B0V type ionizing star (Panagia 1973; Vacca, Garmany, & Shull 1996; Schaerer & de Koter 1997), which is fully consistent with that suggested by the [O III]/ $H\alpha$ ratio of the nebula (Stasińska & Schaerer 1997). Assuming a spherical geometry, we derive an rms electron density of $\sim 25 \text{ cm}^{-3}$, slightly less dense than N44 F. The ionized gas mass of N44 J is then $\sim 200 M_{\odot}$.

The echelle spectrum of N44 J shows a kinematically quiescent H II region. The centroids of the $H\alpha$ and [N II] lines are centered at 292 km s^{-1} , with an rms variation of 3 km s^{-1} (see Figs. 6 and 7). No line splitting is seen. The observed FWHMs are $27 \pm 1 \text{ km s}^{-1}$ for $H\alpha$ and $16 \pm 2 \text{ km s}^{-1}$ for [N II]. These observed line widths im-

ply a turbulent FWHM of about 10 km s^{-1} for $H\alpha$ and 8 km s^{-1} for [N II] if we take into account the instrumental profile and the thermal broadening at 10^4 K . The turbulence can thus explain the small variations of the centroid velocity. The difference between the turbulence widths is probably caused by a difference in the emission regions of the two lines. If N44 J is actually expanding, the expansion velocity should be very small, $< 8 \text{ km s}^{-1}$. For such small expansion velocity, no strong compression is expected, hence no sharp filaments should be seen.

Finally, we may compare the observations with the standard, adiabatic bubble model. As the density in the ring-like nebula is a few tens cm^{-3} , we will assume in our calculations a density of 10 cm^{-3} in the ISM. For an expected wind luminosity of $\sim 3 \times 10^{34} \text{ erg s}^{-1}$ for an O9.5V star (de Jager et al. 1988; Prinja, Barlow, & Howarth 1990) and with the observed $17''.5$ radius of N44 J, the Weaver et al. (1977) model predicts an expansion velocity of about 8 km s^{-1} for N44 J. This expected expansion velocity is just at the limit implied by the velocity widths of the $H\alpha$ and [N II] lines. Raising the ambient density by a factor of 2 can lower the expected expansion velocity to 6.5 km s^{-1} , which is more consistent with the observed limit. We conclude that the observed bubble dynamics of N44 J does not greatly contradict the model predictions.

4.2. N44 M

On the southeast part of N44, another small ring-like nebula is present: N44 M. It is the faintest nebula of our sample. It shows a ring with a radius of $\sim 22''$ (or 5.3 pc), but its north boundary is not well defined (see Fig. 2). The sharp southwest boundary suggests the presence of dust clouds. Several bright stars appear in its interior. N44 M possesses an excitation structure very similar to N44 J. The [S II]/ $H\alpha$ ratio of N44 M is comparable to that in N44 F, but the [O III]/ $H\alpha$ ratio of N44 M is much lower, ranging from 0.07 to 0.16 in the nebula and rising to 0.4 near the stars. The integrated [O III]/ $H\alpha$ ratio of N44 M is about 0.18.

With an observed $H\alpha$ flux of $5.3 \times 10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$ and the two evaluations of the extinction, the required ionizing flux $Q(H^0)$ is $1.5\text{--}2.1 \times 10^{48} \text{ photons s}^{-1}$. This suggests an O9–B0V type ionizing star (Panagia 1973; Vacca, Garmany, & Shull

1996; Schaerer & de Koter 1997), which is again compatible with the $[\text{O III}]/\text{H}\alpha$ ratio (Stasińska & Schaerer 1997). If we assume a spherical geometry, we find that the rms density is about $\sim 17 \text{ cm}^{-3}$ and the estimated mass of the nebula $\sim 275 M_{\odot}$.

The echelle spectrum of N44 M shows only one component, centered at 298 km s^{-1} with an rms variation of 1 km s^{-1} , and a FWHM of $28 \pm 1 \text{ km s}^{-1}$ in the $\text{H}\alpha$ line and $20 \pm 4 \text{ km s}^{-1}$ in the $[\text{N II}]$ line (see Figs. 6 and 7). The turbulent FWHM are then 12 km s^{-1} for $\text{H}\alpha$ and 14 km s^{-1} for $[\text{N II}]$. With an observed radius of $\sim 22''$, a density of 10 cm^{-3} for the ISM and a O9.5V ionizing star (see Section 4.1.), the standard bubble model predicts an expansion velocity of $\sim 7 \text{ km s}^{-1}$ for N44 M, which appears consistent with the upper limit of the echelle observations.

The kinematic characteristics of N44 J and N44 M are very similar, despite the somewhat different morphology. Neither H II region contains shells expanding at velocities greater than $\sim 8 \text{ km s}^{-1}$. In the future, additional observations may enable a better comparison of these nebulae with the standard model. Precise photometry and/or stellar spectroscopy would give accurate informations about the stellar content of these nebulae, while nebular spectra with higher resolution should more strongly constrain the kinematics of N44 J and N44 M. Thus, improved data is needed to more stringently test the standard model for N44 J and N44 M.

5. Conclusion

Massive stars inject a large amount of energy into the ISM through their ionizing radiation, fast stellar winds, and SNe. The stellar winds, which dominate the early mechanical feedback, are able to sweep the ambient medium into an expanding shell and create a cavity around the parent stars. H II regions ionized by massive stars are therefore thought to evolve from compact H II regions to ring-like nebulae.

We have studied three candidate wind-blown bubbles in the N44 complex: N44 F, N44 J, and N44 M. They all have diameters of $\sim 10 \text{ pc}$. Their ionizing fluxes and $[\text{O III}]/\text{H}\alpha$ ratios suggest that their ionizing stars have spectral types of roughly O7V, O9.5V and O9.5V, respectively. N44 F has

an expansion velocity of 12 km s^{-1} , while the other two have expansion velocities $< 8 \text{ km s}^{-1}$. The stellar wind luminosity implied by the ionizing stars and the observed bubble dynamics appear consistent with that expected in the standard, adiabatic bubble model within the limits of observational uncertainties. Similar results were found by Oey & Massey (1994) for two bubbles around single O star in M33. This is in sharp contrast to the anomalous kinematics seen in superbubbles and Wolf-Rayet bubbles. We caution, however, that our observed expansion velocities may not provide the most stringent constraints when comparing with models, because the expected expansion velocities of N44 J and N44 M are comparable to or smaller than the convolution of the isothermal sound velocity ($\sim 10 \text{ km s}^{-1}$) and the instrumental FWHM (13 km s^{-1}). Deeper, higher-resolution spectroscopic observations of the $[\text{N II}]$ or $[\text{O III}]$ lines are needed to pin down the expansion velocity for a more definitive test of models.

We also report the discovery of several dust pillars in N44 F, which are similar to those found in the Eagle Nebula (Hester et al. 1996) and 30 Doradus (Scowen et al. 1998). These structures are indicative of the evaporation of molecular clouds by the ionizing source. The examples in N44 F are especially well-suited to investigating the process and role of photoevaporation in star formation and shell evolution, because of their location within such a well-defined, wind-blown bubble. We found that photoevaporation, as evidenced by the pillar structure, may be dynamically important for the shell evolution in N44 F, slowing down the expanding shell and inhibiting its growth. These photoevaporation processes may thus undermine the seeming agreement between the observed bubble dynamics and the simple adiabatic bubble model.

We thank Charles Danforth for assisting in the echelle observations and Rosie Chen for help in the *HST* WFPC2 image reduction. This work has been supported by the NASA grant STScGO-6698.01-95A. Y.N. acknowledges support from the F.N.R.S., contracts P4/05 and P5/36 ‘Pôle d’attraction Interuniversitaire’ (SSTC-Belgium) and the PRODEX XMM-OM and Integral Projects. R.C.S. acknowledges the support of the Dean B. McLaughlin fellowship at the University of Michigan and NSF grant AST 95-30747.

REFERENCES

- Chu, Y.-H., & Mac Low, M.-M. 1990, *ApJ*, 365, 510
- Davies, R. D., Elliott, K. H., & Meaburn, J. 1976, *MemRAS*, 81, 89
- de Jager, C., Nieuwenhuijzen, H., & van der Hucht, K. A. 1988, *A&AS*, 72, 259
- Drissen, L., Moffat, A. F. J., Walborn, N. R., & Shara, M. R. 1995, *AJ*, 110, 2235
- Feast, M. 1999, in “New Views of the Magellanic Clouds”, *IAU Symposium 190*, ed. Y.-H. Chu, N. Suntzeff, J. Hesser, & D. Bohlender, 542
- Garnett, D. R., Galarza, V. C., & Chu, Y.-H. 2000, *ApJ*, 545, 251
- Goldman, I. 2000, *ApJ*, 541, 701
- Henize, K. G. 1956, *ApJS*, 2, 315
- Hester J. J., et al. 1996, *AJ*, 111, 2349
- Holtzman, J. A., et al. 1995, *PASP*, 107, 156
- Hunter, D. A. 1994, *AJ*, 107, 565
- Johnson, K. E., Leitherer, C., Vacca, W. D., & Conti, P. S. 2000, *AJ*, 120, 1273
- Kennicutt, R. C., Jr., & Hodge, P. W. 1986, *ApJ*, 306, 130
- Lasker, B. M. 1977, *ApJ*, 212, 390
- Lucke, P. B., & Hodge, P. W. 1970, *AJ*, 75, 171
- McKee, C. F., Van Buren, D., & Lazareff, B. 1984, *ApJ*, 278, L115
- Norman, C. A., & Ferrara, A. 1996, *ApJ*, 467, 280
- Oey, M. S., & Massey, P. 1994, *ApJ*, 425, 635
- Oey, M. S., & Massey, P. 1995, *ApJ*, 452, 216
- Oey, M. S. 1996, *ApJ*, 467, 666
- Osterbrock, D. E., Fulbright, J. P., Martel, A. R., Keane, M. J., Trager, S. C., & Basri, G. 1996, *PASP*, 108, 277
- Panagia, N. 1973, *AJ*, 78, 929
- Peimbert, M., Rayo, J. F., & Torres-Peimbert, S. 1975, *Rev. Mex. Astron. Astrof.*, 1, 289
- Prinja, R. K., Barlow, M. J., & Howarth, I. D. 1990, *ApJ*, 361, 607
- Rosado, M., Georgelin, Y. M., Georgelin, Y. P., Laval, A., & Monnet, G. 1982, *A&A*, 115, 61
- Saken, J. M., Shull, J. M., Garmany, C. D., Nichols-Bohlin, J., & Fesen, R. A. 1992, *ApJ*, 397, 537
- Schaerer, D., & de Koter, A. 1997, *A&A*, 322, 598
- Scowen, P. A., et al. 1998, *AJ*, 116, 163
- Spyromilio, J. 1995, *MNRAS*, 277, L59
- Stasińska, G., & Schaerer, D. 1997, *A&A*, 322, 615
- Strickland, D. K., Heckman, T. M., Weaver, K. A., & Dahlem, M. 2000, *AJ*, 120, 2965
- Treffers, R. R. & Chu, Y.-H. 1982, *ApJ*, 254, 569
- Vacca, W. D., Garmany, C. D., & Shull, J. M. 1996, *ApJ*, 460, 914
- Wang, Q. D., & Helfand, D. J. 1991, *ApJ*, 373, 497
- Weaver, R., McCray, R., Castor, J., Shapiro, P., & Moore, R. 1977, *ApJ*, 218, 377
- Weis, K., & Duschl, W. J. 1999, in “New Views of the Magellanic Clouds”, *IAU Symposium 190*, ed. Y.-H. Chu, N. Suntzeff, J. Hesser, & D. Bohlender, 134
- Williams, R. J. R., Ward-Thompson, D., & Whitworth A. P. 2001, *MNRAS*, 327, 788

This 2-column preprint was prepared with the AAS L^AT_EX macros v5.0.

Fig. 1.— CTIO 0.9 m $H\alpha$ ($+[\text{N II}]$) CCD image of N44. The nebular components defined by Henize (1956) and the OB associations cataloged by Lucke & Hodge (1970) are marked.

Fig. 2.— Ground-based images of N44F (left), N44J (middle), and N44M (right). The first row contains $H\alpha$ ($+[\text{N II}]$) images taken with the CTIO 0.9 m telescope. The second, third, and fourth rows are Curtis Schmidt images taken in $H\alpha$ ($+[\text{N II}]$), $[\text{O III}]$, and $[\text{S II}]$ filters, respectively. The field-of-view (FOV) of each image is $92'' \times 92''$. North is up and east to the left.

Fig. 3.— *HST* WFPC2 $H\alpha$ ($+[\text{N II}]$) image (left), $[\text{S II}]$ image (middle) and $[\text{S II}]/H\alpha$ ratio map (right) of N44F. The rectangular boxes show the position of the close-up images presented in Fig. 4. The bars below the images show the greyscales for surface brightness (in units of $\text{erg cm}^{-2} \text{s}^{-1} \text{arcsec}^{-2}$) or line ratio. North is up and east to the left.

Fig. 4.— Close-up images of dust columns seen in N44F through $H\alpha$, $[\text{S II}]$, and $[\text{S II}]/H\alpha$ (for exact position, see Fig. 3). Left: dust columns situated at the west rim of N44F (FOV= $18'' \times 18''$). Right: dust columns situated at the south rim of N44F (FOV= $14'' \times 14''$). The arrows point to the stars at the tips of the pillar-like dust columns.

Fig. 5.— Left: $H\alpha$ image of the “pillars” situated at the west part of N44F. The rectangular box and the arrow show the position and direction of the cut presented at the right. Right: Ionization structure along the rectangular aperture. $[\text{S II}]$ and $H\alpha$ intensities are normalized to the peak, while the $[\text{S II}]/H\alpha$ ratio is the actual value. Note that the $H\alpha$ and $[\text{S II}]$ peaks are separated by one pixel, i.e. $0''.1$.

Fig. 6.— Top panel: Echellogram of the $[\text{N II}] \lambda 6583$ line for N44F. Bottom panels: Radial velocity-position plots of the echelle slit positions in N44F, N44J, and N44M for both $H\alpha$ and $[\text{N II}]$ lines. The radial velocities are heliocentric. The $H\alpha$ velocity components are plotted in filled symbols, and the $[\text{N II}]$ in open ones. The position axis is defined with the center of the nebula at the origin ($0''$) and increasing towards west. The dot-dashed line in each plot indicates the range covered by the nebula. Arrows show the positions of the profiles presented in Fig. 7.

Fig. 7.— $H\alpha$ and $[\text{N II}]$ profiles at the positions indicated in Fig. 6.

This figure "f1.jpg" is available in "jpg" format from:

<http://arXiv.org/ps/astro-ph/0208430v1>

This figure "f2.jpg" is available in "jpg" format from:

<http://arXiv.org/ps/astro-ph/0208430v1>

This figure "f3.jpg" is available in "jpg" format from:

<http://arXiv.org/ps/astro-ph/0208430v1>

This figure "f4.jpg" is available in "jpg" format from:

<http://arXiv.org/ps/astro-ph/0208430v1>

This figure "f5.jpg" is available in "jpg" format from:

<http://arXiv.org/ps/astro-ph/0208430v1>

This figure "f6.jpg" is available in "jpg" format from:

<http://arXiv.org/ps/astro-ph/0208430v1>

This figure "f7.jpg" is available in "jpg" format from:

<http://arXiv.org/ps/astro-ph/0208430v1>