Rotation, spectral variability, magnetic geometry and magnetosphere of the Of?p star CPD $-28^\circ$ 2561


ABSTRACT

We report magnetic and spectroscopic observations and modeling of the Of?p star CPD $-28^\circ$ 2561. Using more than 75 new spectra, we have measured the equivalent width variations and examined the dynamic spectra of photospheric and wind-sensitive spectral lines. A period search results in an unambiguous 73.41 d variability period. High resolution spectropolarimetric data analyzed using Least-Squares Deconvolution yield a Zeeman signature detected in the mean Stokes V profile corresponding to phase 0.5 of the spectral ephemersis. Interpreting the 73.41 d period as the stellar rotational period, we have phased the equivalent widths and inferred longitudinal field measurements. The phased magnetic data exhibit a weak sinusoidal variation, with maximum of about 565 G at phase 0.5, and a minimum of about -335 G at phase 0.0, with extrema approximately in phase with the (double-wave) $H\alpha$ equivalent width variation. Modeling of the $H\alpha$ equivalent width variation assuming a quasi-3D magnetospheric model produces a unique solution for the ambiguous couplet of inclination and magnetic obliquity angles: $(i, \beta)$ or $(\beta, i) = (35^\circ, 90^\circ)$. Adapting either geometry, the longitudinal field variation yields a dipole polar intensity $B_d = 2.6 \pm 0.9$ kG, consistent with that obtained from direct modelling of the Stokes V profiles. We derive a wind magnetic confinement parameter $\eta_c \approx 100$, leading to an Alfvén radius $R_A \approx 3 - 5 R_*$, and a Kepler radius $R_K \approx 20 R_\odot$. This supports a physical scenario in which the $H\alpha$ emission and other line variability have their origin in an oblique, co-rotating ‘dynamical magnetosphere’ structure resulting from a magnetically channeled wind. Nevertheless, the details of the formation of spectral lines and their variability within this framework remain generally poorly understood.

Key words: Stars : rotation – Stars: massive – Instrumentation : spectropolarimetry.

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1 INTRODUCTION

The Of?p stars are mid-O-type stars identified by a number of peculiar observational properties. The classification was first introduced by Walborn (1972) according to the presence of C iv 4650 emission with a strength comparable to the neighbouring N m lines. Well-studied Of?p stars are now known to exhibit recurrent, and apparently periodic, spectral variations (in Balmer, He i, C iv and Si iv lines), narrow P Cygni or emission components in the Balmer lines and He i lines, and UV wind lines weaker than those of typical Of supergiants (see Nazé et al. 2010, and references therein).

Only 5 Galactic Of?p stars are known (Walborn et al. 2010): HD 108, HD 148937, HD 191612, NGC 1624-2 and CPD −28° 2561. Four of these stars - HD 108, HD 148937, HD 191612 and NGC 1624-2 - have been studied in detail based on optical spectra, and in some cases UV spectra as well. In recent years, they have been carefully examined for the presence of magnetic fields (Donati et al. 2006; Martins et al. 2010; Wade et al. 2011, 2012a,b) and all have been clearly detected. It therefore appears that the particular spectral peculiarities that define the Of?p class are a consequence of their magnetism. Indeed, Hubrig et al. (2011, 2012) reported 4 FORS+VLT measurements of the longitudinal magnetic field of CPD −28° 2561 at the few hundred G level, several of which correspond to detections at somewhat more than 3σ significance.

Like HD 108, HD 191612, HD 148937 and NGC 1624-2, CPD −28° 2561 is a spectroscopic variable star (e.g. Walborn et al. 2010). The spectrum of CPD −28° 2561 was first described as a peculiar Of type by Walborn (1973) and Garrison, Hillner & Schild (1977). Garrison et al. commented: 'Very peculiar spectrum. Carbon (C iv λ4070) is strong, nitrogen weak. H and He ii lines are broad, while He i lines are sharp.' Walborn classified the spectrum as O6.5p. The outstanding peculiarity was that although He ii λ4686 emission was very strong, appropriate for an Of supergiant, the N m λ4640 emission was incommensurate.

While Walborn (1973) announced HD 191612 as a new, third member of the Of?p class, it is significant that no such association was made for CPD −28° 2561. The reason was that no C iv λ4650 emission was detected in CPD −28° 2561, while a comparable emission strength to that of N m was the primary defining Of?p characteristic. (Of course, it was subsequently discovered that the C iv emission disappearance actually at the minimum phase of HD 191612 (Walborn et al. 2004; Howarth et al. 2007; Wade et al. 2011).) It was not until the intensive OWN survey’s high-resolution monitoring of CPD −28° 2561 (Barbá et al. 2010) revealed extreme variations in λ4686 and Balmer lines, analogous to those of HD 191612, that the association was made (Walborn et al. 2010) and CPD −28° 2561 identified as an Of?p star.

In this paper we perform a first detailed investigation of the combined magnetic and variability properties of CPD −28° 2561 using an extensive spectroscopic and high-resolution spectropolarimetric dataset. In Sect. 2 we discuss the data acquired and the methods of analysis used. In Sect. 3 we examine the physical properties of the star, as well as its projected rotational velocity. In Sect. 4 we examine the spectral characteristics and variability, identifying periodic variability of the Hα and other emission and absorption lines and deriving the rotational period of the star. In Sect. 5 we analyse in detail the magnetic data acquired at the Canada-France-Hawaii Telescope (CFHT). In Sect. 6 we employ the Hα EW variation and CFHT magnetic data to constrain the stellar and magnetic geometry and the surface field strength. In Sect. 7 we derive the magnetospheric properties of CPD −28° 2561. Finally, in Sects. 8 we summarize our results, and explore the implications of our study, particularly regarding the variability and other properties of CPD −28° 2561, the confinement and structure of its stellar wind, and of the properties of the general class of Of?p stars.

2 OBSERVATIONS

2.1 Spectroscopic observations

Spectroscopic observations of CPD −28° 2561 were obtained in the framework of two spectroscopic surveys: the High-resolution spectroscopic monitoring of Southern O and WN-type Stars (The "OWN Survey", Barbá et al. 2010) and the Galactic O Star Spectroscopic Survey ("GOSSS", Maiz Apellániz et al. 2011). In the OWN Survey program, high-resolution and high signal-to-noise spectra are being acquired for a sample of 240 massive southern stars selected from the Galactic O Star Catalogue (GOSC version 1, Maiz-Apellániz et al. 2004). One of the goals is to determine precise radial velocities in order to detect new binaries among these stars for which there is scarce or no indication of multiplicity. An additional goal is to detect possible spectral variations which can be related to multiplicity, the presence of magnetic fields, pulsation, or eruptive behaviour. CPD −28° 2561 was included early in the observed sample, and monitored systematically as large variations in the intensity of the Hα line were detected. The star was observed spectroscopically from three different locations: Las Campanas Observatory (LCO), and La Silla Observatory (LSO), both in Chile, and Complejo Astronómico El Leoncito (CASLEO), in Argentina, during 33 nights between 2006 and 2012.

Seven spectrograms were obtained with an échelle spectrograph attached to the 2.5 m LCO/du Pont telescope. The spectral resolving power is about 46,000. Twelve spectrograms were obtained with the FEROS échelle spectrograph attached to the 2.2 m ESO/MPI telescope. In this case, the spectral resolving power is about 48,000. Additionally, four spectrograms were obtained with the REOSC échelle spectrograph[1] attached to the 2.15 m CASLEO/Jorge Sahade telescope, with a resolving power of about 15,000. Thorium-Argon comparison lamp exposures were obtained before or after the science exposures. LCO and CASLEO échelle

Table 1. Spectrographs used for the acquisition of CPD −28° 2561 spectroscopy and spectropolarimetry.

<table>
<thead>
<tr>
<th>Spectrograph</th>
<th>Telescope</th>
<th>Spectral Range (Å)</th>
<th>Resolving Power</th>
</tr>
</thead>
<tbody>
<tr>
<td>Echelle</td>
<td>2.5 m LCO/du Pont</td>
<td>3600 – 9200</td>
<td>46,000</td>
</tr>
<tr>
<td>FEROS</td>
<td>2.2 m ESO/MPI</td>
<td>3500 – 9200</td>
<td>48,000</td>
</tr>
<tr>
<td>REOSC</td>
<td>2.15 m CASLEO</td>
<td>3800 – 6000</td>
<td>15,000</td>
</tr>
<tr>
<td>Boller &amp; Chivens</td>
<td>2.5 m LCO/du Pont</td>
<td>3900 – 5500</td>
<td>2,500</td>
</tr>
<tr>
<td>ESPaDOnS</td>
<td>3.6 m CFHT</td>
<td>3650 – 10000</td>
<td>65,000</td>
</tr>
</tbody>
</table>

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[1] Jointly built by REOSC and Liége Observatory, and on long-term loan from the latter.
Table 2. Log of spectroscopic observations showing heliocentric Julian Date, rotational phase according to Eq. (1), and exposure time. The column ‘Spectrograph’ corresponds to the facilities described in Table 1.

<table>
<thead>
<tr>
<th>HJD</th>
<th>Phase</th>
<th>Exp. time (s)</th>
<th>Spectrograph</th>
</tr>
</thead>
<tbody>
<tr>
<td>2453875.483</td>
<td>0.511</td>
<td>600</td>
<td>Echelle</td>
</tr>
<tr>
<td>2454246.515</td>
<td>0.565</td>
<td>1800</td>
<td>FEROS</td>
</tr>
<tr>
<td>2454610.486</td>
<td>0.523</td>
<td>900</td>
<td>B&amp;C</td>
</tr>
<tr>
<td>2454627.467</td>
<td>0.754</td>
<td>1800</td>
<td>FEROS</td>
</tr>
<tr>
<td>2454786.865</td>
<td>0.926</td>
<td>900</td>
<td>Echelle</td>
</tr>
<tr>
<td>2454842.716</td>
<td>0.687</td>
<td>1800</td>
<td>REOSC</td>
</tr>
<tr>
<td>2454846.677</td>
<td>0.741</td>
<td>1800</td>
<td>REOSC</td>
</tr>
<tr>
<td>2454874.609</td>
<td>0.753</td>
<td>2000</td>
<td>REOSC</td>
</tr>
<tr>
<td>2454884.650</td>
<td>0.767</td>
<td>1800</td>
<td>REOSC</td>
</tr>
<tr>
<td>2454954.537</td>
<td>0.210</td>
<td>1800</td>
<td>FEROS</td>
</tr>
<tr>
<td>2454955.500</td>
<td>0.223</td>
<td>1800</td>
<td>FEROS</td>
</tr>
<tr>
<td>2454964.476</td>
<td>0.345</td>
<td>1500</td>
<td>Echelle</td>
</tr>
<tr>
<td>2455341.455</td>
<td>0.481</td>
<td>1200</td>
<td>Echelle</td>
</tr>
<tr>
<td>2455495.586</td>
<td>0.584</td>
<td>900</td>
<td>B&amp;C</td>
</tr>
<tr>
<td>2455604.796</td>
<td>0.068</td>
<td>1800</td>
<td>REOSC</td>
</tr>
<tr>
<td>2455641.586</td>
<td>0.569</td>
<td>1800</td>
<td>REOSC</td>
</tr>
<tr>
<td>2455643.569</td>
<td>0.596</td>
<td>1800</td>
<td>REOSC</td>
</tr>
<tr>
<td>2455646.597</td>
<td>0.637</td>
<td>900</td>
<td>B&amp;C</td>
</tr>
<tr>
<td>2455669.553</td>
<td>0.950</td>
<td>900</td>
<td>B&amp;C</td>
</tr>
<tr>
<td>2455672.532</td>
<td>0.990</td>
<td>1200</td>
<td>Echelle</td>
</tr>
<tr>
<td>2455696.473</td>
<td>0.317</td>
<td>2700</td>
<td>REOSC</td>
</tr>
<tr>
<td>2455697.494</td>
<td>0.331</td>
<td>2400</td>
<td>REOSC</td>
</tr>
<tr>
<td>2455698.518</td>
<td>0.344</td>
<td>2400</td>
<td>REOSC</td>
</tr>
<tr>
<td>2455699.508</td>
<td>0.358</td>
<td>2400</td>
<td>REOSC</td>
</tr>
<tr>
<td>2455716.479</td>
<td>0.589</td>
<td>900</td>
<td>B&amp;C</td>
</tr>
<tr>
<td>2455898.729</td>
<td>0.072</td>
<td>900</td>
<td>B&amp;C</td>
</tr>
<tr>
<td>2455976.599</td>
<td>0.133</td>
<td>1200</td>
<td>Echelle</td>
</tr>
<tr>
<td>2456053.500</td>
<td>0.180</td>
<td>900</td>
<td>B&amp;C</td>
</tr>
<tr>
<td>2456080.469</td>
<td>0.547</td>
<td>1800</td>
<td>Echelle</td>
</tr>
<tr>
<td>2456098.461</td>
<td>0.793</td>
<td>2000</td>
<td>FEROS</td>
</tr>
<tr>
<td>2456340.705</td>
<td>0.092</td>
<td>900</td>
<td>B&amp;C</td>
</tr>
<tr>
<td>2456367.653</td>
<td>0.460</td>
<td>900</td>
<td>B&amp;C</td>
</tr>
<tr>
<td>2456381.587</td>
<td>0.649</td>
<td>900</td>
<td>B&amp;C</td>
</tr>
</tbody>
</table>

2 Spectrograms were reduced using the Echelle package layered in IRAF while FEROS spectrograms were reduced using the standard MIDAS pipeline. All spectrograms were bias subtracted, flat-fielded, echelle order identified, and extracted, and finally wavelength calibrated. Table 1 presents technical details of the different spectrographs utilized.

Under the GOSSS program, we have obtained ten spectrograms of CPD $-28\degree$ 2561 between 2008 and 2013 using the Boller & Chivens spectrograph attached to the 2.5 m LCO/du Pont telescope, with a spectral resolving power of about 2,500. Helium-Neon-Argon comparison lamps were used for wavelength calibration. A dedicated pipeline was developed for the complete reduction of GOSSS observations. Detailed description about the observing procedures and data reduction are described by Sota et al. (2011) and Sota et al. (2014).

The log of spectroscopic observations is reported in Table 1.

2.2 Spectropolarimetric observations

Spectropolarimetric observations of CPD $-28\degree$ 2561 were obtained using the ESPaDOnS spectropolarimeter at the CFHT in 2012 and 2013 within the context of the Magnetism in Massive Stars (MiMeS) Large Program [Wade et al. 2014, Wade et al., in prep.]. Altogether, 44 Stokes V sequences were obtained.

Each polarimetric sequence consisted of four individual subexposures taken in different polarimeter configurations. From each set of four subexposures we derive a mean Stokes V spectrum following the procedure of [Donati et al. 1997], ensuring in particular that all spurious signatures are suppressed at first order. Diagnostic null polarization spectra (labeled $N$) are calculated by combining the four subexposures in such a way that polarization cancels out, allowing us to check that no spurious signals are present in the data (see [Donati et al. 1997] for more details on how $N$ is defined). All frames were processed using the Upena pipeline feeding Libre ESPRIT [Donati et al. 1997], a fully automatic reduction package installed at CFHT for optimal extraction of ESPaDOnS spectra. The peak signal-to-noise ratios per 1.8 km s$^{-1}$ velocity bin in the reduced spectra range from about 70 to nearly 266, with a median of 200, depending on the exposure time and on weather conditions.

The log of CFHT observations is presented in Table 3.

We also obtained one observation of CPD $-28\degree$ 2561 with the HARPS instrument in polarimetric model. The observing procedure was fundamentally the same as that described above for ESPaDOnS. Two consecutive observations were acquired with exposure times of 3600 s apiece. The peak SNR of the combined spectrum was 110 per 1 km/s spectral pixel. The mean HJD of the observation is 2455905.798, corresponding to phase 0.168 according to the ephemeris defined in Eq. (1).

3 STELLAR PHYSICAL AND WIND PROPERTIES

We have used the atmosphere code CMFGEN [Hillier & Miller 1998] to determine the stellar parameters of CPD $-28\degree$ 2561. CMFGEN computes non-LTE, spherical models including an outflowing wind and line-blanketing. A complete description of the code is given by Hillier & Miller (1998). We have included the following elements in our models: H, He, C, N, O, Ne, Mg, Si, S, Ar, Ca, Fe and Ni. A solar metallicity [Grevesse et al. 2010] was adopted.

We used the classical helium ionization balance method to constrain the effective temperature. As shown in Sect. 4, spectral variability is observed in He i lines. Part of the variability may be due to emission from the stellar wind (i.e. the magnetosphere inferred later in the paper), contaminating the underlying photospheric features. The determination of $T_{eff}$ is thus difficult. It was performed using spectra #1051146 (corresponding to phase 0.718 of the rotational ephemeris described in Sect. 4) where the He i lines were the strongest, presumably minimizing contamination from the wind and leaving access to the cleanest photospheric profiles. Nevertheless, some wind-sensitive lines are poorly reproduced. Notably He i $\lambda$4686 exhibits a P Cyg-like profile, with a strong redshifted absorption that (unlike most other absorption features in the spectrum) is under-fit by the model. The profile of this line may be suggestive of accretion.

The relative strength of He i and He ii lines indicates a temperature of about 35000 K, with an uncertainty of 2000 K. The surface gravity was determined from the wings of the numerous Balmer lines available in the ESPaDOnS spectrum. The line cores
The evolutionary mass, determined from the position in the HR diagram and interpollation between the tracks of Meynet & Maeder (2005), is $35 \pm 6 \, M_\odot$. The spectroscopic mass, obtained from the surface gravity and the radius, is $61 \pm 33 \, M_\odot$. The mass estimates are consistent within the error bars, which remain large because of the uncertainties on the distance (and thus on the luminosity). We also note that the evolutionary tracks adopted here do not include the effects of a strong, large-scale stellar magnetic field.

The shape of the photospheric lines could be reproduced with different combinations of rotational velocities and macroturbulence. In the extreme cases, a macroturbulent isotropic velocity of 40 km s$^{-1}$ and a negligible rotational velocity, or no macroturbulence and $v \sin i = 80$ km s$^{-1}$, correctly reproduce the shape of the photospheric lines. Using C m and N m lines as indicators (C m $\lambda 4070$ and N m $\lambda 4505-4515$ being the main diagnostics), we estimated the nitrogen and carbon content to be N/H $< 5.0 \times 10^{-5}$ and C/H $= (5.0 \pm 3.0) \times 10^{-5}$. Both elements are underabundant com-
Table 4. Summary of physical, wind and magnetic properties of CPD $-28^\circ$ 2561. Apart from the luminosity $L/L_\odot$ (which was adopted), all parameters are derived in this paper. The mass-loss rate corresponds to the CMFGEN radiatively-driven rate required to reasonably reproduce the H\alpha profile; this value is only indicative, since the spherical symmetry assumed by CMFGEN is clearly broken in the case of CPD $-28^\circ$ 2561. The wind magnetic confinement parameter $\eta_*$, the rotation parameter $W$ and the characteristic spin down time $\tau_{\text{spin}}$ are defined and described in Sect. 7. Values in brackets for $\eta_*, R_A$ and $R_K$ correspond to the best-fit parameters.

<table>
<thead>
<tr>
<th>Parameter</th>
<th>Value</th>
</tr>
</thead>
<tbody>
<tr>
<td>Spectral type</td>
<td>O$^3$p</td>
</tr>
<tr>
<td>$T_{\text{eff}}$ (K)</td>
<td>$35,000 \pm 2000$</td>
</tr>
<tr>
<td>log $g$ (cgs)</td>
<td>$4.0 \pm 0.1$</td>
</tr>
<tr>
<td>$R_*$ ($R_\odot$)</td>
<td>$12.9 \pm 3.0$</td>
</tr>
<tr>
<td>log($L_*/L_\odot$)</td>
<td>$5.35 \pm 0.15$</td>
</tr>
<tr>
<td>$M_{\text{tot}}$ ($M_\odot$)</td>
<td>$35 \pm 6$</td>
</tr>
<tr>
<td>$M_{\text{spec}}$ ($M_\odot$)</td>
<td>$61 \pm 33$</td>
</tr>
<tr>
<td>$v_{\text{sin} i}$ (km s$^{-1}$)</td>
<td>$\leq 80$</td>
</tr>
<tr>
<td>$P_{\text{rot}}$ (d)</td>
<td>$73.41 \pm 0.05$</td>
</tr>
<tr>
<td>log $M_{\text{B-o}}$ (M$_{\odot}$ yr$^{-1}$)</td>
<td>$-6.0$</td>
</tr>
<tr>
<td>$v_\infty$ (km s$^{-1}$)</td>
<td>$2400$</td>
</tr>
<tr>
<td>$B_0$ (G)</td>
<td>$2600 \pm 900$</td>
</tr>
<tr>
<td>$i$ (°)</td>
<td>$35 \pm 3$</td>
</tr>
<tr>
<td>$\beta$ (°)</td>
<td>$90 \pm 4$</td>
</tr>
<tr>
<td>$\eta_*$</td>
<td>$20-900$ (93)</td>
</tr>
<tr>
<td>$R_A$ ($R_\odot$)</td>
<td>$2.1-5.5$ (3.4)</td>
</tr>
<tr>
<td>$R_K$ ($R_\odot$)</td>
<td>$14.2-29.9$ (18.7)</td>
</tr>
<tr>
<td>$\tau_{\text{spin}}$</td>
<td>$0.45$ Myr</td>
</tr>
<tr>
<td>N/H</td>
<td>$&lt;5 \times 10^{-5}$</td>
</tr>
<tr>
<td>C/H</td>
<td>$(5 \pm 3) \times 10^{-5}$</td>
</tr>
</tbody>
</table>

4 SPECTRAL VARIABILITY AND PERIOD

Many lines in the spectrum of CPD $-28^\circ$ 2561 show significant variability.

Following a similar analysis as applied by the MiMeS collaboration to other magnetic O-type stars (e.g. Grunhut et al. 2012, Wade et al. 2012), we first characterise the line variability using the equivalent width (EW). Before measuring the EW, each spectral line was locally re-normalised using the surrounding continuum, and the EWs were computed by numerically integrating over the line profile. The $1\sigma$ uncertainties were calculated by propagating...
the individual pixel uncertainties in quadrature. Precomputed pixel
error bars were only available for the ESPaDOnS and HARPS spec-
tra, which required us to assign a single uncertainty to each pixel
for the other spectra that was inferred from the RMS scatter of the
Stokes I flux in the continuum regions around each spectral line.

A period search was performed on the EW measurements
from all spectra using the Lomb-Scargle technique (Press et al.
1992) and an extension of this technique to higher harmonics
(Schwarzenberg-Czerny 1996). In this approach, a harmonic func-
tion with free parameters \(a_1, a_2, a_0\) and \(P\) (the latter being the pe-
riod) corresponding to the form \(a_1 \sin(\omega t) + a_2 \cos(\omega t) + a_0\), where
\(\omega = 2\pi/P\), is fit to the data. Adding the first harmonic leads
to the introduction of additional terms of the form \(a_3 \sin(2\omega t) +
a_4 \cos(2\omega t)\). The resulting periodograms are defined as the power
\[\sum_{j=1}^{4} a_j^2\] versus period \(P\).

The periodograms from the strongly variable lines (such as
H\(\alpha\), H\(\beta\), He i \(\lambda\)5876 and He ii \(\lambda\)4686) show significant power at
\(\sim 36.7\) d. When phased with this period, the EW variations show
clear sinusoidal variations. However, the sinusoidal variations of
the He i \(\lambda\)5876 measurements show considerably more scatter than
the other lines. Furthermore, a comparison of spectra obtained at
different epochs, but at similar phases according to the \(\sim 36.7\) d pe-
riod, show important differences. These include a systematic veloc-
ity offset, a difference in peak emission flux, and opposite skew, of
the emission line profiles. This effect is particularly striking for
H\(\alpha\) and He ii \(\lambda\)4686 observations.

The phased EW measurements are illustrated in Fig. 3.

![Figure 3](image-url)

**Figure 3.** Periodogram obtained from the H\(\alpha\) EW (solid black) and He ii \(\lambda\)4686 EW (solid red) measurements. A clear and unique signal at 36.7 d is detected.

The detailed explanation is no different then the implementa-
tion used in fsrch:

We therefore carried out a new period search including the
contributions from the first harmonic - by including additional har-
monics we change not only the relative power in each peak, but
also the precise location of the peaks. The resulting periodograms
showed two clear peaks, one consistent with the previous peri-

- **Figure 4.** Comparison of H\(\alpha\) and He ii \(\lambda\)4686 profiles obtained at similar phases, but at different epochs. **Left -** Assuming a period of 36.7 d, profile shapes at similar phases do not agree. **Right -** Assuming a period of 73.4 d, profile shapes at similar phases agree well. As discussed in the text, we identify the \(P = 73.41 \pm 0.05\) as the rotation period based on the better agreement of the profiles at all phases.

- **Figure 5.** Normalized intensity profiles at 36.7 and 73.41 d.

odogram at \(\sim 36.7\) d and a new peak at 73.41\(\pm 0.05\) d, twice the
previously identified period. When phased with this longer period,
the EW variations continue to phase coherently. However, this
new phasing achieves a much better agreement between the line
profile shapes from spectra corresponding to similar phases, but
obtained at different epochs. We therefore conclude that a period
of 73.41 d provides a much better solution to the phase variabil-
ity of CPD \(\sim 28^\circ\) 2561, and we adopt this period as the stellar rota-
tion period within the context of a magnetically-confined wind
and the oblique rotator model (Stibbs 1950; Babel & Montmerle
1997; Wade et al. 2011). Adopting maximum H\(\alpha\) emission (min-
imum EW) as the reference date we derive the following ephemesis:

\[\text{H\(\alpha\)max} = 2454645.49(05) + 73.41(05) \cdot E,\]

where the uncertainties (1\(\sigma\) limits in the last digits are indicated
in brackets). The uncertainties represent the formal 1\(\sigma\) uncertainty
computed from the \(\chi^2\) statistic corresponding to second-order har-
monic fits carried out on the EW variations for periods near 73 d.

The phased EW measurements are illustrated in Fig. 4. The
measurements obtained from the various spectral datasets agree
reasonably well. Examination of Fig 5 shows that most lines with
significant variability exhibit double-wave variations. Hα shows the most significant variability with a peak-to-peak EW variation of ~5 Å. Maximum emission occurs at phase 0 for most lines, while another emission peak occurs one-half of a cycle later. The value of the EW measurements of Hα and He λ4686 at maximum emission are similar at both emission maxima, while a higher emission level occurs at phase 0.5 for some other lines (this is most evident in the EW curves of Hβ and He λ4587), although there is considerable scatter at these phases. Unlike HD 148937, which showed most significant variability with a peak-to-peak EW variation of 0.2 mag for λ4686 and 0.04 mag for λ4647 and C λ5876, CPD −28′ 2561 shows no significant variability in these lines. We also include EWs measured from the DIB at 5797 Å to illustrate the lack of variability of a reference line that is not formed in the environment of the star.

Hipparcos and ASAS data are also available for this star, and we used these data to attempt to detect photometric variability and confirm the derived period. Only the best quality data were kept: for Hipparcos, this means keeping only data with flag 0; for ASAS data, two filterings were used - either keeping only data with grade A (and discarding four strongly discrepant points with V > 10.1 mag) or keeping data only at 3 σ from the mean (both filtering yielded the same results). Errors amount typically to 0.03 mag for Hλ (Hipparcos), 0.2 mag for Bλ and Vλ (Tycho), and 0.04 mag for V (ASAS). A χ² test for constancy was performed on each dataset. Only Hλ data were found to be significantly variable with significance level SL < 1%. However, a Fourier period search on those data reveals only white noise, without any significant peaks (or in particular a peak at the expected period of 73.41d). Furthermore, that period does not yield a significant coherent phased variation. Fourier period searches on the Tycho data yield similar conclusions. Period searches (Fourier (Gosset et al. 2001), PDM (Stellingwerf 1978), and entropy (Cincotta, Mendez & Nunez 1995)) on the ASAS data result in the detection of marginal signals with periods of 63.69 ± 0.13d and 73.43±0.16d. The amplitude is very small, ~ 0.0046 mag at most. When phased with these periods, the photometry only shows marginal variations, even after binning the phased data (Fig. 1). We therefore conclude that any astrophysical variability of the photometric signal is below the typical uncertainty of a few tens of mmag, and would require high-precision photometry to be securely detected.

4.1 Line profile variations

4.1.1 Line profiles of CPD −28′ 2561 at high resolution

As discussed above, it is now established that CPD −28′ 2561 is the only O?p star known so far that has double emission-line maxima, corresponding to a geometry that presents both magnetic poles during the rotational cycle. The magnetic O stars HD 47129 (Grunhut & Wade 2013) and HD 57682 (Grunhut et al. 2012) do likewise. In the former case, the complex spectrum of this SB2 system precludes any identification of O?p characteristics. In the latter case, its wind density is too low to produce the defining He λ2561 and N λ5876 emission lines, while it does display the characteristic, variable Balmer emission profiles (first noted in a single Hα observation by Walborn 1980). Thus, HD 57682 does provide points of comparison for CPD −28′ 2561.

An early indication that the double-wave period of ~73 d is the correct one was provided by the strikingly opposite slopes of the He λ4686 emission peaks at phases 0 and 0.5 (Fig. 7). In addition, there is a velocity offset between the two maximum phases. Both of these effects are also seen in the Balmer emission lines to a lesser degree. Significantly, both effects are likewise seen at Hα in HD 57682. These fine morphological details undoubtedly code significant information about the complex circumstellar structures producing them.

The He λ4686 and Balmer profiles at the intermediate phases (i.e., at or near 0.25 and 0.75) are also noteworthy but display diversity between the two objects. In CPD −28′ 2561, they all have weak emission near the blueward edges of absorption features at both intermediate phases. As already noted, HD 57682 has no He λ4686 emission, but remarkably, the Hα profiles have the weak emission components shifted in opposite directions at the two intermediate phases. Again, these subtle similarities and differences are important physical clues that need to be modeled and understood.

Turning to the systematic variations of He λ4686 during the CPD −28′ 2561 cycle, some further interesting effects are evident. The weak emission toward the shortward sides of the intermediate-phase profiles gradually strengthens and shifts toward the line centre beginning about 0.3 cycle before each maximum (which again, have opposite skews and a small velocity difference). In contrast, however, the disappearance of the maximum emission lines is remarkably abrupt, occurring within about 0.1 cycle. This behavior may be suggestive of a sharp occultation of the line-emitting region.

The behaviour of the dilution-sensitive line He λ5876 is entirely distinct from those of the features just described, which are either similar or opposite between phases 0.5 apart. That of λ5876 is egregiously asymmetrical (Fig. 1): at phase 0.5 it displays a weak absorption flanked by equal emission wings, whereas at phase 0.0 it has a relatively weak but well marked P Cygni profile. At phases 0.25 and 0.75 it is a strong, symmetrical absorption line.

On the other hand, He λ5411 is a moderate absorption line with strong wings at all phases, which at first glance appears to present velocity shifts opposite to those of the λ4686 and Balmer emission lines at the two maxima. However, on closer inspection the effect is seen to be due to wing emissions in the same sense as those of the other lines, although possibly with a larger amplitude. In contrast, C λλ5801, 5812 are weak absorption features with no velocity shifts, albeit with a weak P Cygni tendency at phase 0.0, i.e., in the sense of λ5876. No doubt this apparent chaos will be transformed into valuable diagnostics as our understanding of the intricate phenomenology advances.

Finally, the relative behaviours of He λ54471 and He λ4542 (Fig. 1) are of considerable interest, since their ratio is the primary spectral-type criterion, although classification is a heuristic exercise here because of the effects of variable circumstellar emission on the line strengths; the minimum spectrum is expected to be more closely related to the actual stellar parameters. Indeed, both lines are seen to be affected by wing emission at both maxima, albeit the He λ line much more strongly. In fact, the λ4542 line has the unique characteristic that its profile at phase ~ 0.5 is very similar to that exhibited at the quadrature phases 0.25 and 0.75. Hence it appears that this line is only affected by one of the emission maxima.

It should be noted that the He λ effect is weaker than in other O?p spectra, whereas the He λ one is unprecedented, weaker than but analogous to that in λ5411 described above. In terms of the line depths, the maximum spectral type is O7, but the He λ line has a larger equivalent width so that measurements would yield O6–O6.5. On the other hand, the minimum spectral type is a well defined O8, which again may be presumed to correspond to the stellar photosphere.
Figure 5. Rotationally phased EW measurements for selected lines in the spectrum of CPD $−28^\circ$ 2561. Least-squares sinusoidal fits are shown for those lines with statistically significant EW variability. In the bottom row we also show measurements of the $\lambda5797$ DIB as an example of an intrinsically non variable feature. In these figures, the ESPaDOnS measurements are black squares, the HARPS measurement is the cyan upwards facing triangle, the FEROS measurements are the blue circles, the LCO/echelle measurements are the green diamonds, the CAS measurements are the pink downward facing triangles, and the B&C measures are red crosses.

### 4.1.2 Dynamic spectra

The phased line profile variations are displayed as dynamic spectra in Fig. 6. These images were constructed by subtracting the profile that presented the least overall emission in Hα (corresponding to phase $\sim$0.75 - obtained on HJD 2454627.467). The colour scheme was chosen to maximise the dynamic range of the variability to highlight changes relative to the minimum profile.

The dynamic spectra of the most strongly variable lines (such as the Balmer lines (Hα, Hβ, Hγ), He i $\lambda5876$ and He ii $\lambda4686$) exhibit very similar characteristics. In each case the profiles indicate the presence of structures that vary in intensity, from absorption to strong emission, twice per cycle. The two emission features reach similar peak intensities (as already inferred from their EW variations). The emission features of the Balmer lines are slightly asymmetric in velocity about their central peaks and are found to be broad relative to the width of the spectral line (the emission features in Hα have a FWHM of $\sim$280 km s$^{-1}$). The emission feature that reaches maximum emission at phase 0 appears to be offset from the mean velocity by about $-30$ km s$^{-1}$; the emission feature at phase 0.5 is offset by $+30$ km s$^{-1}$.

While the line profile variability in He i $\lambda5876$ and He ii $\lambda4686$ appears similar to the Balmer lines, there are some outstanding
Figure 6. Phased variations of selected spectral lines in the high resolution spectra, shown as dynamic spectra. Plotted is the difference between the observed profiles and the profile obtained of HJD 2454627.467 (dashed-red, top panel), which occurred at phase ∼0.75, corresponding to minimum observed emission in Hα.
The two emission features appear considerably more asymmetric. In fact, the emission feature occurring at phase 0.5 appears to be a blend of two distinct emission peaks in both lines. The higher intensity peak is centred around -60 km s\(^{-1}\) for the He\(\text{II}\) line (or -35 km s\(^{-1}\) for the He\(\text{I}\) line), while the less intense peak is centred about 30 km s\(^{-1}\) for the He\(\text{II}\) line (or 50 km s\(^{-1}\) for the He\(\text{I}\) line). While the central velocities of these features differ between these two lines, their separation is similar. The emission features that appear at phases 0 and 0.5 reach similar peak intensity for the Balmer lines, whereas the intensity of the emission peak at phase 0 is about 10% stronger than the peak occurring at phase 0.5 for the He\(\text{II}\) line and 20% stronger for the He\(\text{I}\) line.

The lines that display weaker variability exhibit a character of variability that is similar to those described for the previous lines. The He\(\text{I}\) \(\lambda 6678\) line shows weaker emission than the previously discussed spectral lines, but there is still evidence for an emission feature occurring twice per cycle. This feature occurs at a velocity relative to the mean velocity that is similar to the other He lines. Furthermore, the relative intensity of the redward emission feature appears considerably stronger than the blueward feature in the He\(\text{I}\) lines that display less variability (the maximum emission intensity is about 60 percent stronger than the blueward emission feature of the \(\lambda 6678\) line). Some of the weakly-variable lines also show evidence of enhanced absorption (relative to the minimum profile) in a narrow region around the line core. The enhanced absorption reaches a maximum relative absorption in the core at about phase 0.5 (which corresponds to the phase of maximum emission at the blue edge of this line).

The He\(\text{II}\) \(\lambda 4542\) line exhibits only weak variability. However, it stands out from essentially all other lines in the spectrum of CPD –28° 2561 due to its apparent single-wave variation, a phenomenon reflected in the line profiles examined in Sect. 4.1.1. In contrast to the double-wave variation of the other lines illustrated in Figs. 5 and 6 which show EW maxima at phases 0.0 and 0.5, the \(\lambda 4542\) line appears to exhibit an EW maximum at phase 0.0, but a minimum at phase 0.5.

5 DIAGNOSIS OF THE MAGNETIC FIELD

Least-Squares Deconvolution (LSD, Donati et al. 1997) was applied to all CFHT observations using the LSD code of Kochukhov, Makaganiuk & Piskunov (2010). In their detection of
the magnetic field of HD 191612. Donati et al. (2006) developed and applied an LSD line mask containing 12 lines. Similar masks were successfully employed by Wade et al. (2011, 2012) in their analyses of HD 191612 and HD 148937. Given the similarity between the spectra of CPD −28° 2561 and HD 191612, we began with this line list and adjusted the predicted line depths to best match the depths observed in the spectrum of CPD −28° 2561 at phases ∼ 0.25 and 0.75, when the emission is lowest. This involved adjustment of the line depths by typically ∼ 20% relative to their depths in the spectrum of HD 191612. We then used this line list to extract mean circular polarization (LSD Stokes V), mean polarization check (LSD I) and mean unpolarized (LSD Stokes I) profiles from all collected spectra. All LSD profiles were produced on an 1800 km s⁻¹ spectral grid with a velocity bin of 20 km s⁻¹, using a regularization parameter of 0.2 (for more information regarding LSD regularisation, see Kochukhov, Makaganiuk & Piskunov 2010).

Using the 9σ signal detection criteria described by Donati et al. (1997), we evaluated the significance of the signal in both the Stokes V and N LSD profiles in the velocity range [-150, 250] km s⁻¹, consistent with the observed span of the Stokes I profile. No significant signal was detected in any of the individual V (or N) profiles. We also computed the longitudinal magnetic field from each profile set using Eq. (1) of Wade et al. (2000a). To improve our sensitivity, we coadded the LSD profiles of spectra acquired within ±2 nights (resulting in 9 averages of 2-8 spectra; see Table 3). Given the observed variability period (Sect. 5), ±2 nights corresponds to approximately 0.05 cycles - a timespan during which variability should be limited; indeed, this was verified empirically. From these profiles we obtain one marginal detection of signal (false alarm probability FAP < 10⁻³) in the V profiles (for the coadded profile corresponding to spectral IDs 1604363-1604387), and best longitudinal field error bars of ∼ 170 G. The most significant measurement of the longitudinal field from the coadded profiles is −362 ± 186 G (2.0σ). The individual and coadded ESPaDOnS spectra and the corresponding longitudinal fields and detection probabilities are indicated in Table 3. The longitudinal field measured from the HARPSpol spectrum was 280 ± 460 G.

We conclude that we fail to detect a magnetic field in individual and co-added Stokes V spectra of CPD −28° 2561. To proceed further, we note the similarity of the spectrum of CPD −28° 2561 to HD 191612 (and in particular its O'p classification), and its periodic variability, combined with the reported detection of a strong magnetic field by Hubrig et al. (2011, 2012), strongly suggest that CPD −28° 2561 is an oblique magnetic rotator. This has been convincingly demonstrated for HD 191612 (Wade et al. 2011), and is consistent with the behaviour of other magnetic O-type stars (e.g. Wade et al. 2012b; Grunhut et al. 2012; Wade et al. 2012a).

We therefore proceeded to bin the LSD profiles according to rotational phase as computed via Eq. (1) in order to increase the signal-to-noise ratio. We note that this phase binning implicitly assumes an oblique magnetic rotator, i.e. that the magnetic field variation proceeds according to the same period as the spectra (and other variations). We have binned the coadded LSD profiles from phases 0.30–0.619 and 0.981–0.042. From these profiles (illustrated in Fig. 8) we obtain a definite detection of signal in the Stokes V profile (at phase ∼ 0.5) and no detection (although with twice poorer SNR) at phase ∼ 0.0, with no signal detected in the null profiles. Notwithstanding the lack of formal detection at phase 0, the Stokes V LSD profile appears to exhibit a weak signature with polarity opposite to that at phase 0.5. The phase 0.0 and 0.5 LSD profiles yield longitudinal fields of ⟨Bz⟩ = +335 ± 200 G and ⟨Bz⟩ = −290 ± 95 G respectively, and corresponding null profile fields of ⟨N⟩₀ = −71 ± 199 G and ⟨N⟩₀ = −2 ± 95 G. We have also searched the average spectrum at phase ∼ 0.5 for Zeeman signatures in individual line profiles. Weak signatures, compatible with the LSD Stokes V profiles, are visible in the He i λ5876 and C iv λ5801 lines. We also note that coadding LSD profiles phased using one-half the adopted period (i.e. 36.7 d) yields no detection of the magnetic field.

From this analysis, which implicitly assumes that CPD −28° 2561 is an oblique magnetic rotator, we conclude that an organized magnetic field is detected in the photosphere of CPD −28° 2561, with a longitudinal field with a characteristic strength of several hundred G, that likely changes sign.

6 STELLAR AND MAGNETIC GEOMETRY

With the inferred rotational period and derived radius (from Table 1), CPD −28° 2561 should have an equatorial rotational velocity...
in the range $v_c \approx 7 - 11 \text{ km s}^{-1}$. Since the upper limit on $v \sin i$ obtained in Sect. 3 (see Table 4) is much larger than this value due to the significant turbulent broadening of the line profile, we are not able to constrain the inclination of the stellar rotational axis via comparison of $v_c$ and $v \sin i$.

Instead, we derive both the rotational and magnetic geometry of CPD $-28^\circ$ 2561 from considering the rotationally modulated Hα line stemming from the star’s hypothetical circumstellar “dynamical magnetosphere” (Sundqvist et al. 2012). Below the Alfvén radius $R_A$, at which the magnetic and wind energy densities are equal, the magnetic field is strong enough to channel the radiatively driven wind outflow of CPD $-28^\circ$ 2561 along closed field lines (ud-Doula & Owocki 2002). The trapped wind plasma, channeled along field lines from opposite magnetic hemispheres then collides at the magnetic equator, and is pulled back to the star’s surface by gravity. This in turn leads to a statistically overdense region centered around the magnetic equator, which is also characterized by infalling material of quite low velocities (as compared to non-magnetic O star wind velocities). If the magnetic and rotational axes of the star are mutually inclined, an observer at earth will view this dynamical magnetosphere from different perspectives, which leads to rotationally modulated line profiles (Sundqvist et al. 2013, Grunhut et al. 2013, ud-Doula et al. 2013, Petit et al. 2013). Below we use this variability to derive constraints on the magnetic geometry of CPD $-28^\circ$ 2561.

We follow the procedure developed by Sundqvist et al. (2012) (see also Grunhut et al. 2012 and ud-Doula et al. 2013) and use 100 snapshots of a 2-D radiation magnetohydrodynamic (MHD) wind simulation of a magnetic O-star, that we patch together in azimuth to form a 3-D “orange slice” model. The MHD simulation employed here is that computed for HD 191612 (see Sundqvist et al. 2012) as a proxy for CPD $-28^\circ$ 2561. This is a reasonable approach because both HD 191612 and CPD $-28^\circ$ 2561 are slow rotators (in the sense that rotation is dynamically insignificant in determining the wind/magnetosphere properties), and their magnetic field strengths/confinement parameters are formally identical (see Sect. 7). Hence the geometry of the plasma confinement should be similar in both cases (ud-Doula & Owocki 2002).

The tests by ud-Doula et al. (2013) show this patching technique results in a reasonably good representation of the full 3-D magnetosphere. To compute synthetic Hα spectra, we then solve the formal solution of radiative transfer in a 3-D cylindrical system for an observer viewing from angle $\alpha$ with respect to the magnetic axis. For magnetic obliquity $\beta$ and observer inclination $i$ we then have

$$\cos \alpha = \sin \beta \cos \Phi \sin i + \cos \beta \cos i,$$

which gives the observer’s viewing angle as function of rotational phase $\Phi$, thus mapping out the rotational phase variation for a given couple of $\beta$ and $i$. We solve the formal solution only in the wind, assuming Hα occupation numbers and a temperature structure given by a 1-D NLTE model atmosphere calculation, and using an input photospheric Hα line profile as lower boundary condition (see Sundqvist et al. 2012 for more details). The infalling material in our selected 100 snapshots falls preferably toward one pole. This is most presumably due to a subtle numerical issue in the MHD simulations, where over a time this north-south asymmetry is canceled out (see ud-Doula & Owocki 2002). To account for the fact that in these non-rotating stars there should be no preference between the north and south magnetic pole, we here average computed line profiles from the same angle $\alpha$ as counted from the north and south magnetic poles, respectively.

The absolute level of Hα emission should further, in principle, provide constraints on the rate by which the magnetosphere is fed by radiatively driven wind material (in analogy with how Hα emission from non-magnetic O-stars provides constraints on the stellar mass loss). However, as discussed in detail by Grunhut et al. (2012), adjusting the underlying model so that the level of Hα emission is reproduced in the high state results in variability between the high and low states that is too small to reproduce the observations. Since here we are mainly interested in obtaining constraints on the magnetic geometry from the variability itself, we thus compute synthetic Hα equivalent width curves that have somewhat too

![Image](357x642 to 455x662)
strong absolute emission and then simply shift them down so that the absolute level of emission at the extrema are fit. This then results in equivalent width curves that reproduces well the variability, providing good constraints on the geometry. However, because of this shifting we are not able to simultaneously derive constraints on the mass feeding rate; this issue will be discussed in detail in a forthcoming paper.

The right panel of Fig. [B] shows the \( \chi^2 \) landscape from fitting the observed rotational phase variation for given sets of \( \beta \) and \( i \). Because of the strong emission dependence on \( \alpha \), the error bars of the best fit \( i = 90 \, \text{deg} \) and \( \beta = 35 \, \text{deg} \) are quite small. We note, however, that simply switching the obliquity and inclination angles gives equal results, i.e. there is a second "best-fit" model at \( i = 35 \, \text{deg} \) and \( \beta = 90 \, \text{deg} \) (that is not shown in the figure). The left panel then finally compares the best model with the observed variability as function of rotational phase.

6.1 Photometric and broadband polarization variability

A second constraint on the geometry is potentially derived from photometric and broadband (linear) polarimetric variability. We use the Monte-Carlo radiative transfer (RT) code developed by RHDT for simulating light scattering in circumstellar envelopes, first applied in this context to the Of?p star HD 191612 by [Wade et al. 2011]. In this code, photon packets are launched from a central star and allowed to propagate through an arbitrary distribution of circumstellar matter (described by a Cartesian density grid), until they are scattered by free electrons. Upon scattering, a ray is peeled off from the packet toward a virtual observer, who records the packet’s Stokes parameters appropriately attenuated by any intervening material (see [Yusef-Zadeh, Morris & White 1984] for a discussion of this peel-off technique). A new propagation direction is then chosen based on the dipole phase function ([Chandrasekhar 1960]), and the packet’s Stokes parameters are updated to reflect the linear polarization introduced by the scattering process. The propagation is then resumed until, after possible further scatterings, the packet eventually escapes from the system or is reabsorbed by the star.

We employ the circumstellar density model developed for HD 191612, and compute lightcurves for the two geometries derived from the orange-slice MHD modelling. In Fig. [C] we compare the observed ASAS V-band photometric variation with the predictions of the Monte Carlo RT code. Unfortunately, the available photometry is not sufficiently precise to allow us to obtain meaningful constraints on the geometry. Nevertheless, the predictions demonstrate that the existing photometry is at the threshold of being able to detect the predicted variation. High-resolution photometry and broadband polarimetry would provide an additional constraint on the geometry, and in particular the polarimetric variation would enable the determination of the individual values of the angles \( i \) and \( \beta \).

6.2 Longitudinal magnetic field variation

To infer the strength of the magnetic field, and to test its compatibility with the derived geometry, we model the longitudinal field variation of CPD \(-28^\circ \) 2561 as a function of rotational phase. We used the coadded profiles discussed in Sect. 5 and reported in Table \( \text{[2]} \) phased according to Eq. (1). This phase variation \( (B_i)(\phi) \) of the longitudinal field is illustrated in Fig. [D] (upper frame). A fit by Least-Squares of a cosine curve of the form \( (B_i)(\phi) = B_0 + B_1 \cos(2\pi(\phi - \phi_0)) \) to the data yields a reduced \( \chi^2 \) of 0.84, with parameters \( B_0 = +115 \pm 55 \, \text{G}, \ B_1 = 450 \pm 100 \, \text{G} \) and \( \phi_0 = 0.18 \pm 0.15 \).

Therefore, according to Least Squares, the variation of the field is detected at 4.5\( \sigma \) confidence. The reduced \( \chi^2 \) of the data relative to the straight line \( (B_i)(\phi) = 0 \) (the hypothesis of a null field) is 2.3. These results indicate that the variation is significant at about 97\% confidence, and that the null hypothesis can be rejected as an acceptable representation of the data at similar confidence. The longitudinal field measured from the null profiles (lower frame of Fig. [D]) yields similar reduced \( \chi^2 \) for both the cosine and straight-line fits, and indicates no significant variation.

Adopting \( i = 35^\circ \) and \( \beta = 90^\circ \) from the modelling of the Hz EW, we have fit the phase-binned longitudinal field measurements with a synthetic longitudinal field variation with fixed phase of maximum (phase 0.5) and variable polar field strength \( B_d \). The best-fit model (according to the \( \chi^2 \) statistic) is characterised by \( B_d = 2.6 \, \text{kG} \). The extrema of the best-fit dipole model are slightly offset from the best-fit sinusoid in mean longitudinal field strength (by about 100 \( \text{G} \)) and in phase (by about 0.07 cycles). These offsets are well within the uncertainties of the observed variation. Taking them into account, we estimate an uncertainty on the derived dipole strength of \( \pm 900 \, \text{G} \).

If we fix only the inclination and allow both \( \beta \) and \( B_d \) to vary, a direct fit to the longitudinal field variation yields best-fit values of \( \beta = 78.4^{\circ} \) and \( B_d = 2.7 \pm 1.1 \, \text{KG} \). These values are in good agreement with those derived from the fit with fixed geometry derived above.

In Fig. [E] we also show the longitudinal field measurements of Hubrig et al. (2011, 2013) obtained from ‘all’ lines. Those measurements, which have formal errors that are substantially more precise than our own, are in good agreement with both the best-fit synthetic dipole and the best sinusoidal fit. This is noted that 3 of the measurements of Hubrig et al. were acquired at essentially the same phase, which corresponds to crossover (i.e. \( \langle B_i \rangle \approx 0 \)).

6.3 Modeling the Stokes V profiles

We also modelled the magnetic field geometry using the LSD Stokes V profiles. We compared the mean LSD profiles for phases 0.0, 0.5 and 0.75 to a grid of synthetic Stokes V profiles using the method of Petit & Wade (2012). For this modelling we used LSD profiles extracted using a metallic line mask, as described by [Wade et al. 2012a].

The emergent intensity at each point on the stellar surface is calculated using the weak-field approximation for a Milne-Eddington atmosphere model. In this model, the source function is linear in optical depth such that \( S(\tau) = 5 S_0[1 + \beta \tau \kappa] \). We use \( \beta = 1.5 \), Voigt-shaped line profiles with a damping constant \( \alpha = 10^{-3} \) and a thermal speed \( v_{\text{th}} = 5 \, \text{km} \, \text{s}^{-1} \). The line-to-continuum opacity ratio \( \kappa \) is chosen to fit the intensity LSD profile. The synthetic flux profiles are then obtained by numerically integrating the emergent intensities over the projected stellar disk. The projected rotational velocity is set to \( v \sin i = 9 \, \text{km} \, \text{s}^{-1} \). We applied isotropic Gaussian macroturbulence\(^4\) compatible with that determined in Sect. 3.

We assume a simple centred dipolar field, parametrized by the dipole field strength \( B_{d_0} \), the rotation axis inclination \( i \) with respect to the line of sight, the positive magnetic axis obliquity \( \beta \) and the rotational phase \( \varphi \). Assuming that only \( \varphi \) may change between different observations of the star, the goodness-of-fit of a given rotation-
Figure 10. Predicted light curve and broadband polarization variations of CPD $-28^\circ$ 2561. Variations are shown for two models: $(i = 35^\circ, \beta = 90^\circ)$, and $(i = 90^\circ, \beta = 35^\circ)$. Whereas both models produce identical light curves, their polarization variations are clearly distinct.

Figure 11. Top panel: Longitudinal field versus phase. Dashed line - best-fit sinusoid. Solid line - Dipole model (fixed phase and geometry, polar strength $B_d = 2.6$ kG). Diamonds represent ESPaDOnS measurements from the current study. The cross is the HARPSpol measurement. The triangles are measurements from Hubrig et al. (2011, 2013). Bottom panel: Null field versus phase. All measurements are phased according to Eq. (1).

Table 5. Odds ratios derived from the analysis of Stokes $V$ profiles.

<table>
<thead>
<tr>
<th>Phase</th>
<th>$\log(M_0/M_1) V$</th>
<th>$\log(M_0/M_1) N$</th>
</tr>
</thead>
<tbody>
<tr>
<td>0.00</td>
<td>-0.91</td>
<td>0.09</td>
</tr>
<tr>
<td>0.50</td>
<td>-7.43</td>
<td>0.24</td>
</tr>
<tr>
<td>0.75</td>
<td>-0.34</td>
<td>0.06</td>
</tr>
<tr>
<td>Combined</td>
<td>-9.06</td>
<td>0.25</td>
</tr>
</tbody>
</table>

Figure 12 shows the posterior probability density function for each model parameters. The 68.3%, 95.4%, 99.0%, and 99.7% confidence regions, tabulated in Table 6, are illustrated in dark to pale shades, respectively. At 95.4% confidence, the polar strength of the dipole magnetic field of CPD $-28^\circ$ 2561 is found to be in the range $1.9 \leq B_d \leq 4.5$ kG. This is in good agreement with the dipole strengths derived from the longitudinal field variation.

7 MAGNETOSPHERE

As presented by ud-Doula & Owocki (2002), the global competition between the magnetic field and stellar wind can be char-
Figure 12. Magnetic field polar strength (upper frame) and geometry ($i, \beta$) (lower frame) constraints derived from of modelling of Stokes $V$ profiles.

Table 6. Credible regions derived from the Bayesian analysis of Stokes $V$ profiles.

<table>
<thead>
<tr>
<th>Credible region</th>
<th>Range in gauss</th>
</tr>
</thead>
<tbody>
<tr>
<td>$V$</td>
<td></td>
</tr>
<tr>
<td>99.7</td>
<td>1754 - 5000</td>
</tr>
<tr>
<td>99.0</td>
<td>1841 - 4947</td>
</tr>
<tr>
<td>94.5</td>
<td>1891 - 4509</td>
</tr>
<tr>
<td>68.3</td>
<td>2137 - 3307</td>
</tr>
</tbody>
</table>

$N$

<table>
<thead>
<tr>
<th>Credible region</th>
<th>Range</th>
</tr>
</thead>
<tbody>
<tr>
<td>99.7</td>
<td>0 - 3835</td>
</tr>
<tr>
<td>99.0</td>
<td>0 - 2767</td>
</tr>
<tr>
<td>95.4</td>
<td>0 - 1508</td>
</tr>
<tr>
<td>68.3</td>
<td>0 - 425</td>
</tr>
</tbody>
</table>

significant stellar rotation is pulled by gravity back onto the star with a dynamical (free-fall) time-scale.

To estimate the Alfvén radius of the magnetosphere of CPD $-28^\circ 2561$, we use the stellar parameters given in Table 4.

The stellar parameter with the largest uncertainty is the wind momentum. Fig. 13 therefore illustrates the variation of $R_A$ with one order of magnitude variation in mass-loss rate (corresponding to a generous estimate of the uncertainty). One can see how this uncertainty is mitigated by the 1/4 power dependence of the wind momentum in the definition of the Alfvén radius. The grey shaded areas represent intervals of stellar radius uncertainty ($\pm 3 R_\odot$) and of dipole field strength (2.1, 2.6, 3.3 kG, reflecting the 68.3% Bayesian credible region) meant to minimize and maximize the Alfvén radius.

We therefore expect the Alfvén radius to be of the order of 3 stellar radii, and certainly no more than 5 stellar radii.

In the presence of significant stellar rotation, centrifugal forces can support any trapped material above the Kepler co-rotation radius $R_K \equiv (GM/\omega^2)^{1/3}$. This requires that the magnetic confinement extend beyond this Kepler radius, in which case material can accumulate to form a centrifugal magnetosphere (e.g. Townsend, Owocki & Groote 2005).

In the case of CPD $-28^\circ 2561$, the slow rotation puts the Kepler radius much farther out than the Alfvén radius ($\sim 18 R_\odot$) and no long-term accumulation of wind plasma is expected, as illustrated by the red curve in Figure 13 (the red shaded area represents $\pm 3 R_\odot$ and a range of mass from 30 to 60 $M_\odot$). In such a dynamical magnetosphere configuration, transient suspension of circumstellar material results in a statistical global over-density in the closed loops. For O-type stars with sufficient mass-loss rates, the resulting dynamical magnetosphere can therefore exhibit strong emission in Balmer recombination lines (Sundqvist et al. 2012; Petit et al. 2013). This conclusion supports the MHD modelling employed to determine the stellar geometry in Sect. 6. We note that due to infalling wind material (which may be reflected in the P Cyg-like profile of He II $\lambda 4686$ at some phases), the global mass-loss rate of a star exhibiting such a dynamical magnetosphere is significantly reduced. According to the scaling relations of Ud-Doula, Owocki & Townsend (2008) (their eqn. 10 and 23), for an estimated $r_A \approx 3 R_{\text{star}}$, $M/M_{\text{B,0}} \approx 0.2$, i.e. the global mass-loss rate is reduced by approximately a factor of 5 due to wind plasma falling back upon the star.

The characteristic magnetic braking time can be estimated using equation 25 of Ud-Doula, Owocki & Townsend (2009). Using the nominal wind parameters in Table 4 and $k \sim 0.1$ (Claret 2004), we obtain a spin-down timescale of 0.45 Myr. However as noted by Petit et al. (2013), the square-root dependence of this quantity on the wind momentum renders spin-down estimates valid only to a factor of a few.

8 DISCUSSION AND CONCLUSIONS

In this paper we have performed a first thorough analysis of the variability, geometry, magnetic field and wind confinement of the Of?p star CPD $-28^\circ 2561$. Using more than 75 new medium and high resolution spectra, we determined the equivalent width variations and examined the dynamic spectra of photospheric and wind-sensitive spectral lines, deriving a rotational period of 73.41 d. We confirmed the detection of an organized magnetic field via Zeeman signatures in LSD Stokes $V$ profiles. The phased longitudinal field data exhibit a weak sinusoidal variation, with maximum of about...
yield $i + \beta = 135^\circ$, implying that both magnetic poles are presented clearly to the observer during each stellar rotation. Only two other magnetic O-type stars are known to exhibit analogous behaviour: HD 57682 (Grunhut et al. 2012), for which $i + \beta = 139^\circ$, and Piskett’s star (e.g. Grunhut & Wade 2013, and Grunhut et al. 2015, in prep.), for which $i + \beta = 130^\circ$.

The behaviour of the $\lambda 4542$ line can likely be understood as due to the depth of formation and the relative contribution of the magnetosphere to the line profile. Martins et al. (2014) constrained the line formation region of various lines in spectra of non-magnetic O stars, and investigated the variability as a function of depth. They demonstrate that the Balmer lines are formed further out in the wind compared to some weak He i or high ionization He ii or C iv lines. Lines formed closer to the photosphere also exhibit some additional variability that likely arises in the photosphere (and this may qualitatively explain the different behaviour of lines such as He ii $\lambda 4542$). However, the variability of these photospheric lines is weak in comparison to the variability of wind-dominated lines, and cannot contribute appreciably to the variability of features such as H\alpha or He ii $\lambda 4686$. We therefore propose that in the case of CPD $-$28° 2561 (and likely other magnetic O stars), lines clearly formed out in the magnetosphere (H\alpha, H\beta, He ii $\lambda 4686$) are overwhelmingly dominated by the rotational variation of the density of the confined wind plasma.

The dynamic spectra of CPD $-$28° 2561 reveal characteristic differences in the mean velocities and skewness of the two emission maxima per rotation cycle. This phenomenon - first discussed by Grunhut et al. (2012) in the context of HD 57682 - is not reproduced by the orange-slice MHD model. In the case of HD 57682, the shift between the two emission peaks is about 70 km/s; for CPD $-$28° 2561, it is comparable (about 60 km/s). This effect could potentially be explained by a difference in strength and/or geometry of the two magnetic poles at the stellar surface. Such differences are commonly inferred for cooler magnetic stars and are frequently parametrized using a dipole model offset from the stellar centre along its axis. Typically, dipole offsets measured for Ap/Bp stars are 0.0-0.3 $R_*$ (e.g. Wade et al. 1997, 2000b). First 2D MHD simulations computed using reasonable dipole offsets provide promising results. However, their presentation and discussion is outside the scope of this paper. This potential ability of dipole offset to explain the observed emission line variations will be the topic of a forthcoming paper.

The absolute level of H\alpha emission from the magnetosphere of an O star should provide constraints on the rate by which it is fed by radiatively driven wind material through comparison with theoretical models such as those described in Sect. 6. However, as discussed in detail by Grunhut et al. (2012), adjusting the underlying MHD model so that the level of H\alpha emission is reproduced in the high state results in variability between the high and low states that is too small to reproduce the observations. This was interpreted by those authors as a consequence of a magnetosphere consisting of somewhat optically thicker plasma ‘snakes’ (dense plasma clumps falling back toward the star) than predicted by 2D MHD models. Such thicker clumps would lead to higher variability at a given mass-loss feeding rate, and so to a better fit to the observations for the mass-loss feeding rate that reproduces the mean level of emission in the high state.

For CPD $-$28° 2561, we reproduced the EW variation of the H\alpha line by computing synthetic H\alpha equivalent width curves with somewhat too strong absolute emission and then artificially shifting them so that the absolute levels of emission at the extremae were fit. While this procedure resulted in equivalent width curves that
reproduced well the variability, we were not able to simultaneously derive constraints on the mass feeding rate. Recent monitoring of CPD −28° 2561 has been performed with HST/STIS (UV spectroscopy) and XMM (X-ray spectroscopy) [PI: Nazé]. These data will likely help to mitigate the problem of the star’s uncertain wind properties, and will provide important insights into its wind confinement and magnetosphere.

In conclusion, we underscore that the weak longitudinal magnetic fields, complex spectra and faint apparent magnitudes of all Of?p stars challenge the capabilities of current polarimetric instrumentation. While their magnetic fields are securely detected, field topologies are only very roughly characterized. The subject of the present paper - CPD −28° 2561 - along with HD 148937, represent the most challenging examples of the class. The development of spectropolarimetric capabilities such as those of ESPaDOnS, but on 8m-class telescopes, is a vital step required in order to investigate stars such as these at an acceptable level of detail, and more generally to extend our understanding of the magnetic properties of the most massive stars.

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